

KIDs for Millimeter-Wave Cosmology

Adam Anderson
Fermilab Cosmic Day
30 October 2023

Key Question

If CMB-S4 construction and operations last through ~2040, is there anything new left to do in ground-based CMB / millimeter-wavelength surveys?

Key Question

If CMB-S4 construction and operations last through ~2040, is there anything new left to do in ground-based CMB / millimeter-wavelength surveys?

YES! Newer detector technologies (e.g. kinetic inductance detectors) can provide higher-density focal planes. Much more optimal for:

- *Millimeter-wavelength intensity mapping*
- *High-frequency CMB observations*

Key Question

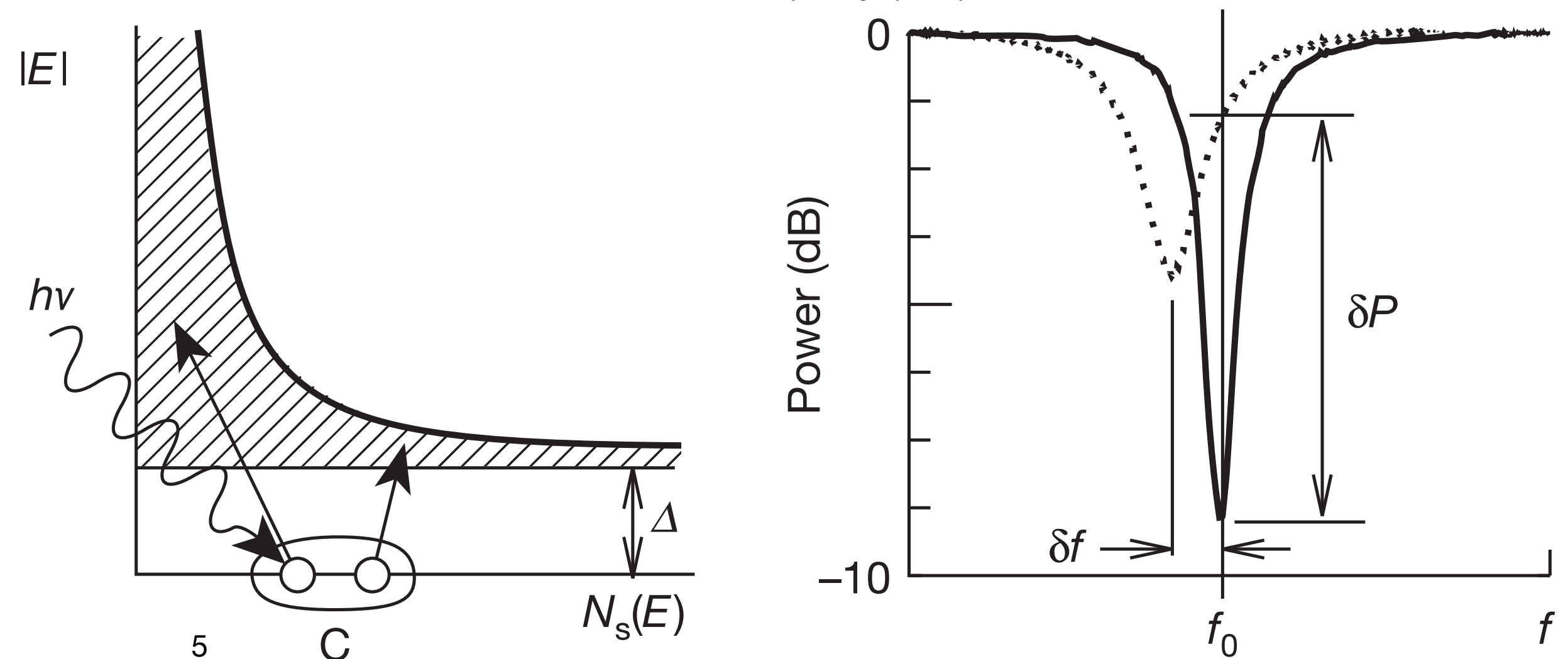
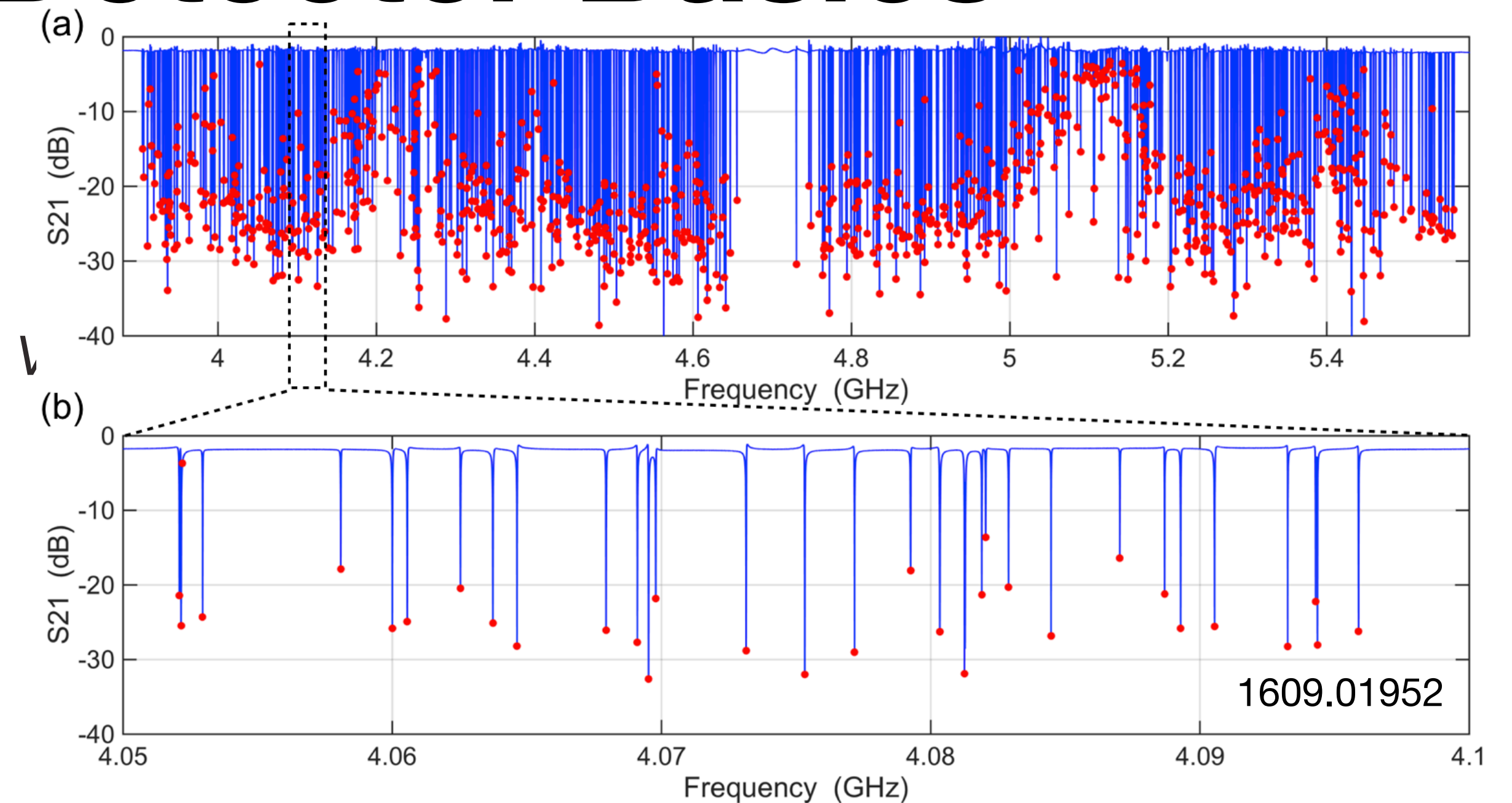
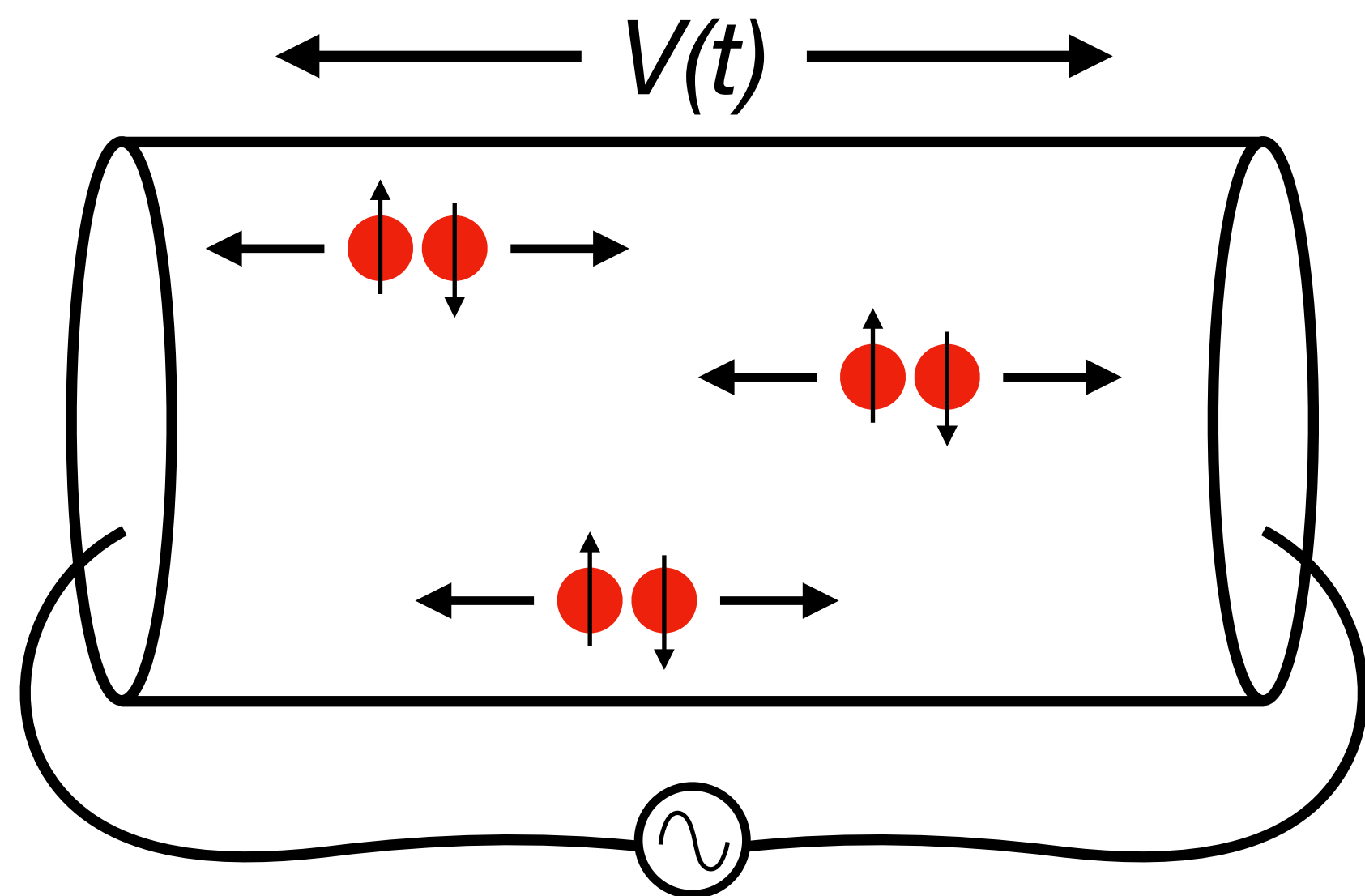
If CMB-S4 construction and operations last through ~2040, is there anything new left to do in ground-based CMB / millimeter-wavelength surveys?

YES! Newer detector technologies (e.g. kinetic inductance detectors) can provide higher-density focal planes. Much more optimal for:

- *Millimeter-wavelength intensity mapping → **SPT-SLIM***
- *High-frequency CMB observations → **SPT-3G+***

Kinetic Inductance Detector Basics

- Cooper pairs in AC potential have nonzero inertia, results in phase shift between I and V. **Looks like an inductance: “kinetic inductance”**.
- Breaking of Cooper pairs by e.g. photons, changes kinetic inductance.
- Create a sensor by coupling the superconductor in an LC circuit so that inductance can be sensed as a change in resonant frequency.

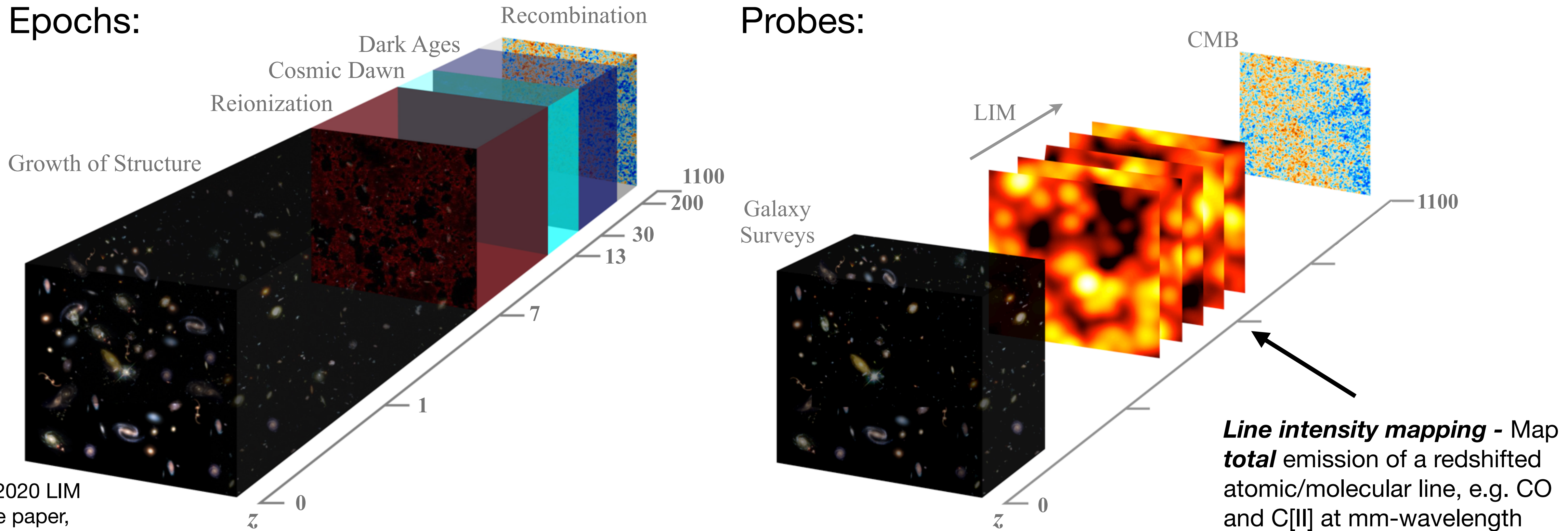


SPT-SLIM: Line Intensity Mapping

Line Intensity Mapping (LIM)

Epochs:

Probes:

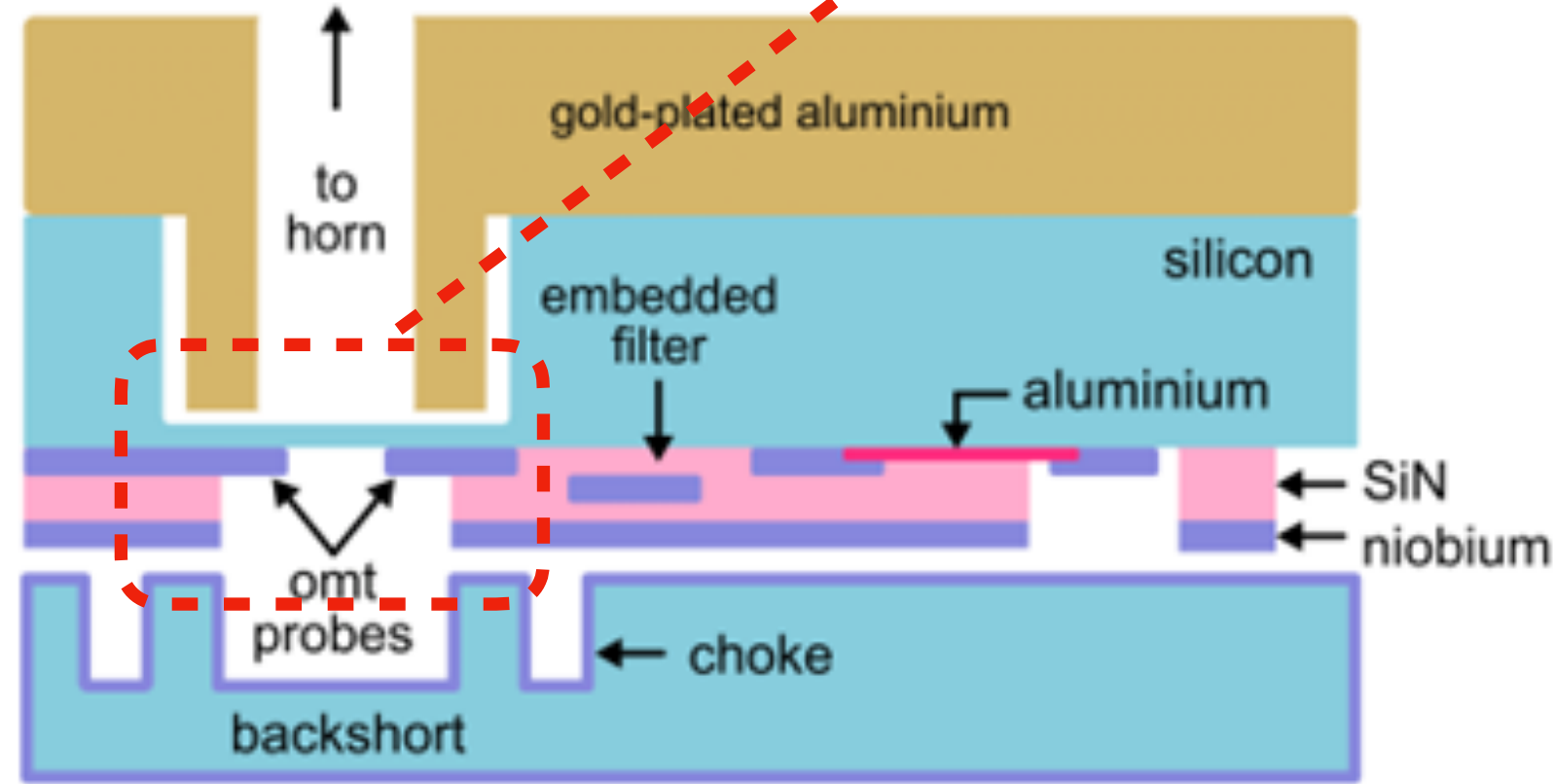


Astro2020 LIM
white paper,
1903.04496

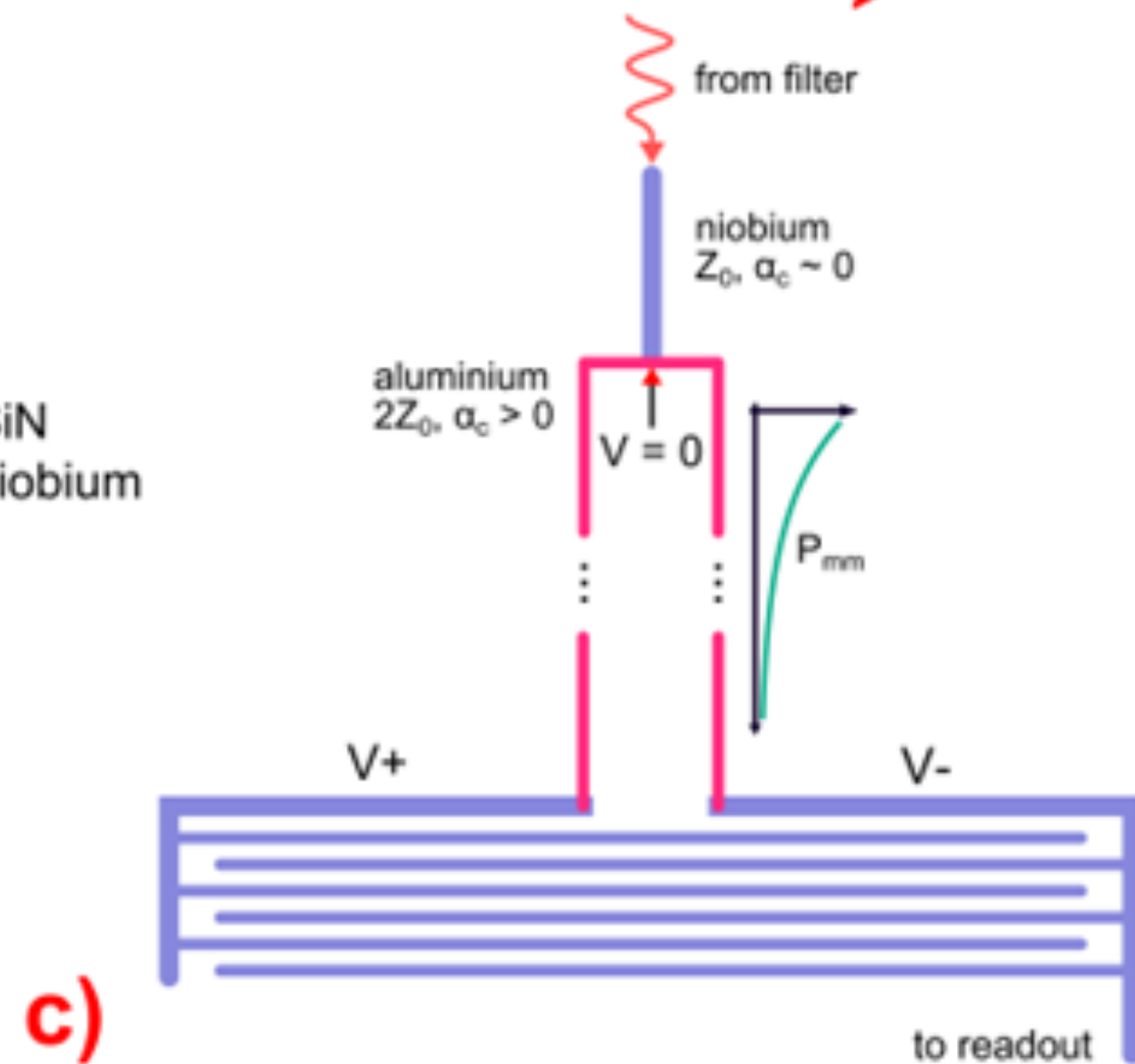
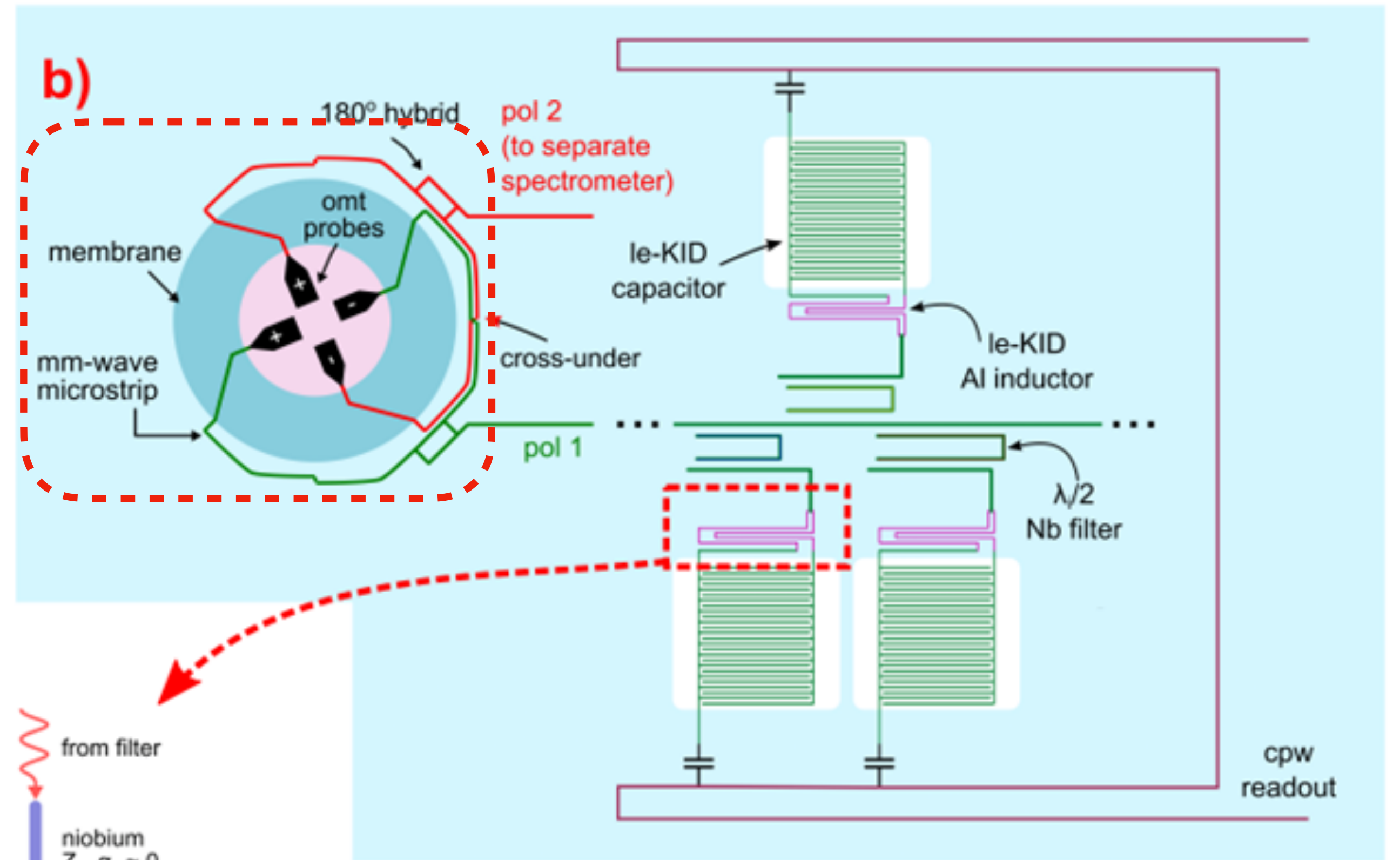
- Large redshift range $z > 3$ that is relatively unexplored: provides information on expansion history of universe, reionization, star-formation, ++
- LIM is efficient: we measure *all* sources and do not need to threshold on galaxies
- Multiple lines available across radio, millimeter, IR

On-Chip Spectrometers

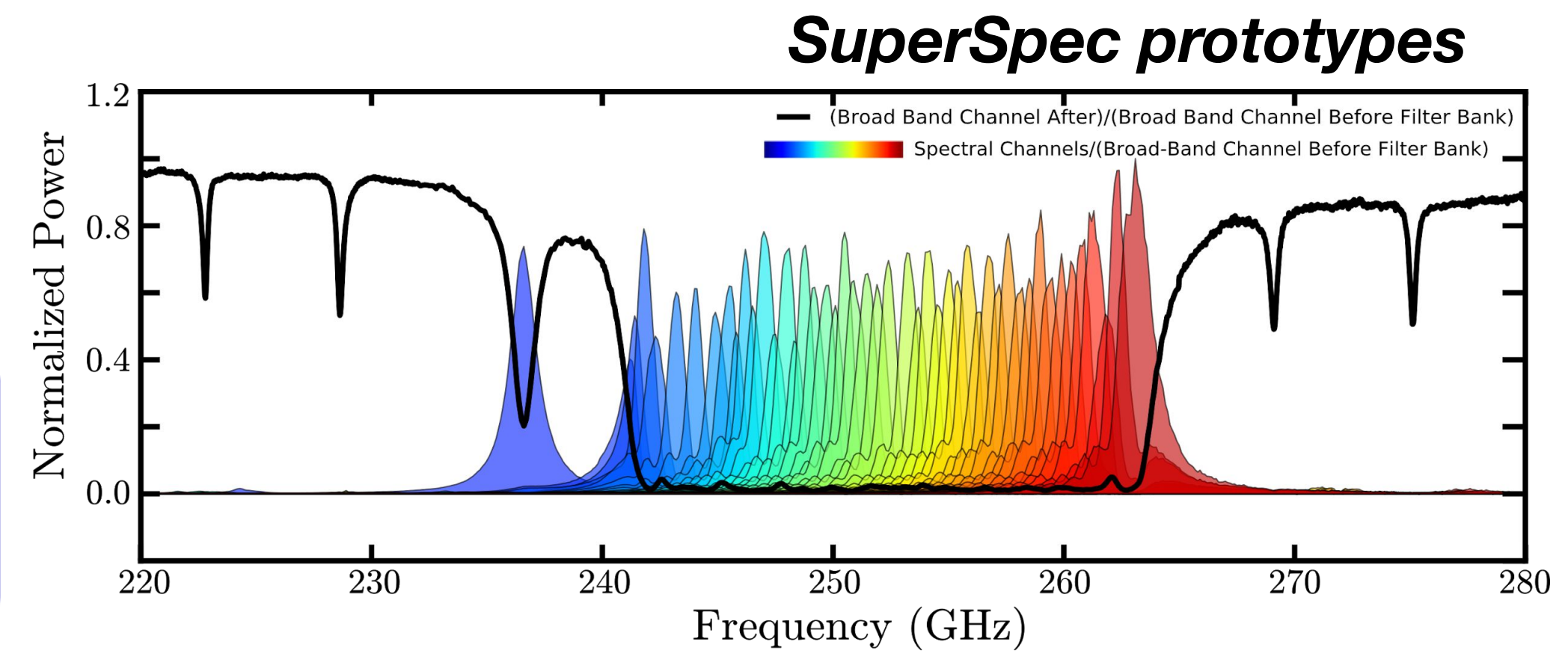
Light from telescope



Cross section



c)

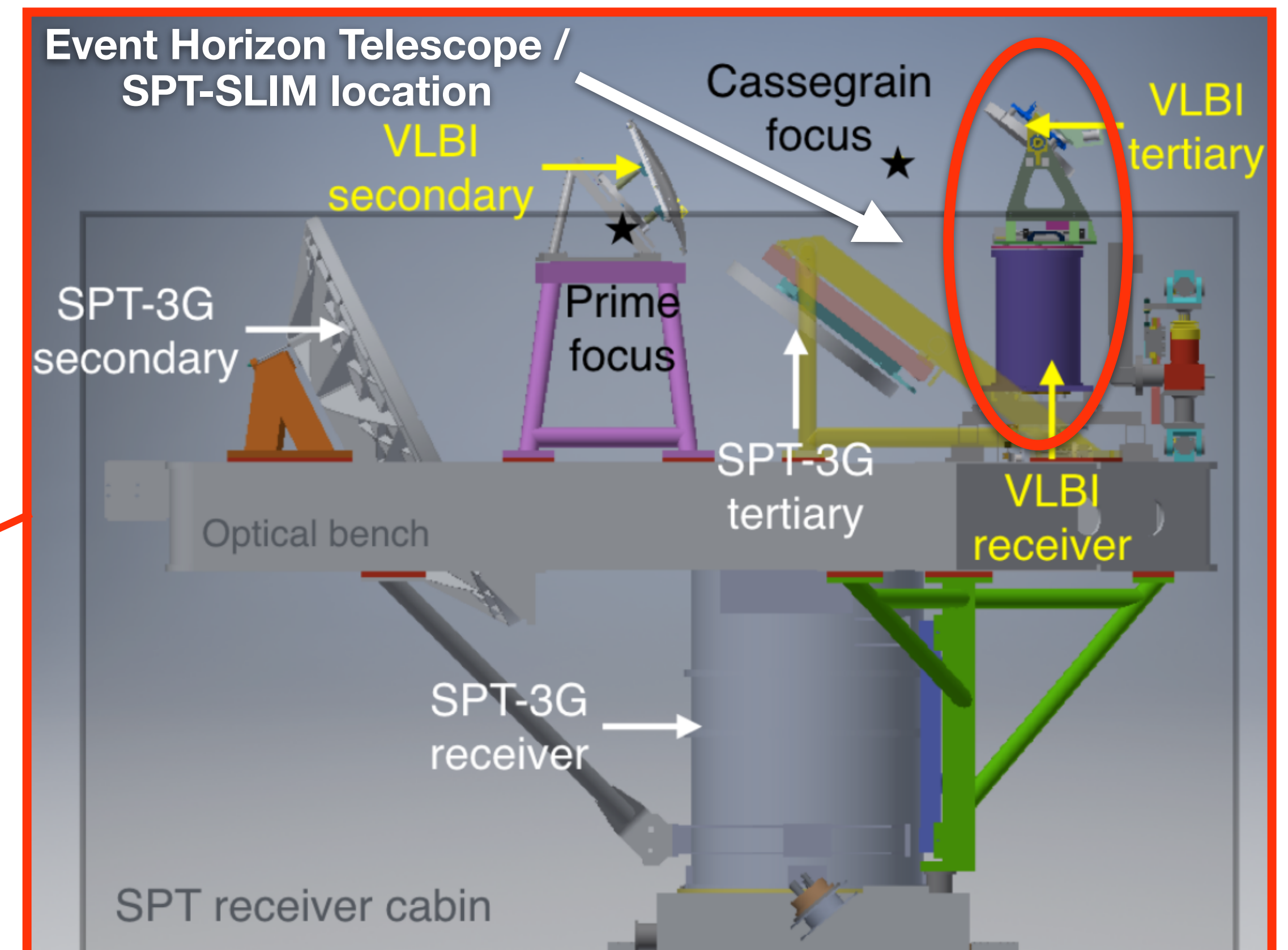
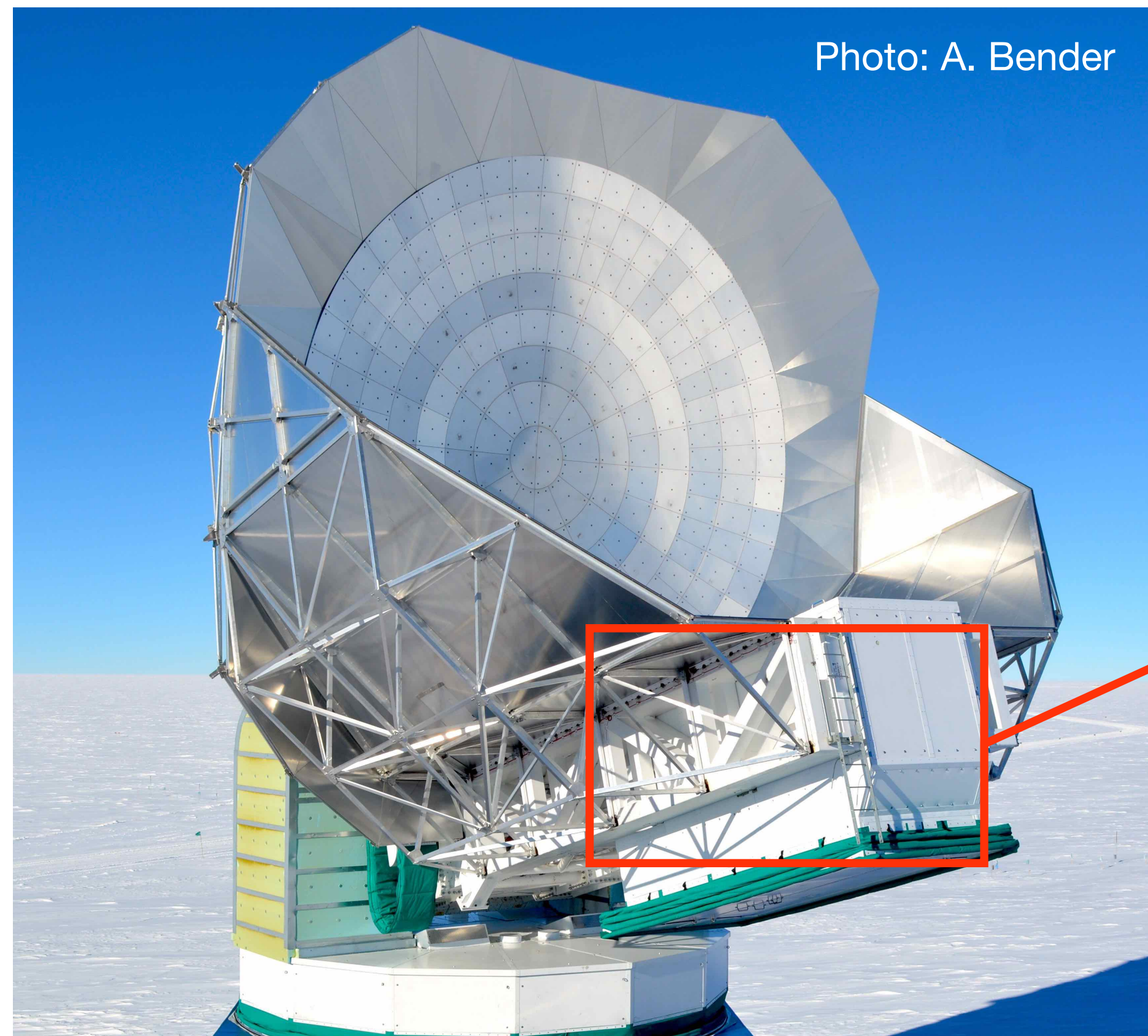


SPT-SLIM Experimental Concept

South Pole Telescope is 10-m CMB telescope observing at 90/150/220 GHz during both austral winter *and* summer

SPT optics include mount point for optional receiver, used by Event Horizon Telescope (EHT) during 2017-present

SPT-SLIM - Replace EHT cryostat with on-chip spectrometers and observe for one summer season



SPT-SLIM: South Pole Telescope Summertime Line Intensity Mapper

Argonne:

T. Cecil
C. Chang
Z. Pan

Cardiff:

P. Barry
G. Robson

Harvard:

G. Keating

Fermilab:

A. Anderson
B. Benson
S. Simon
M. Young

McGill:

M. Adamic
M. Dobbs
M. Rouble

SLAC

K. Karkare
A. Saleem
C. Zhang

U. Arizona:

D. Kim
D. Marrone

U. Chicago:

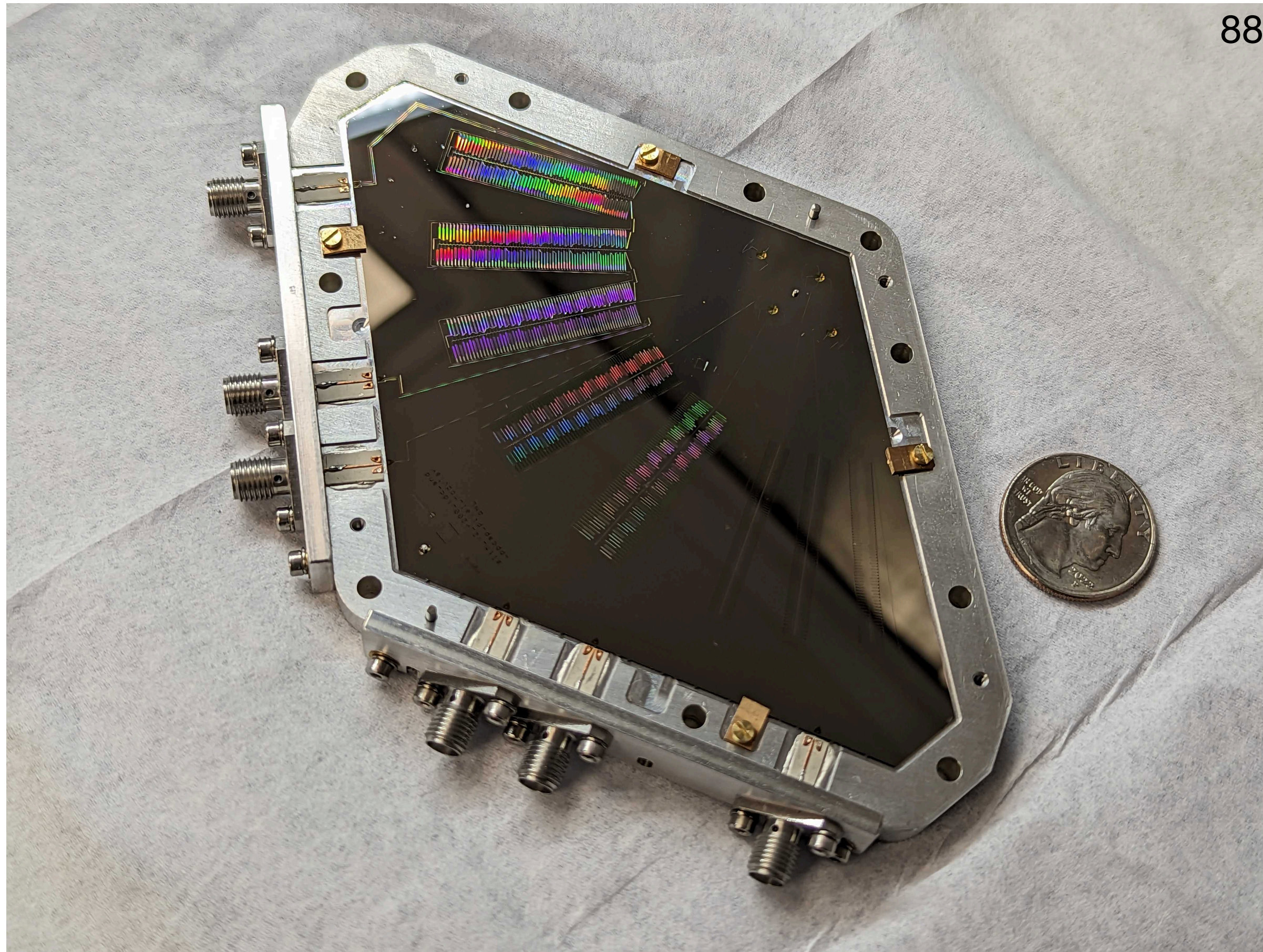
K. Dibert
K. Fichman
T. Natoli
J. Zebrowski

Three-Speed Logic

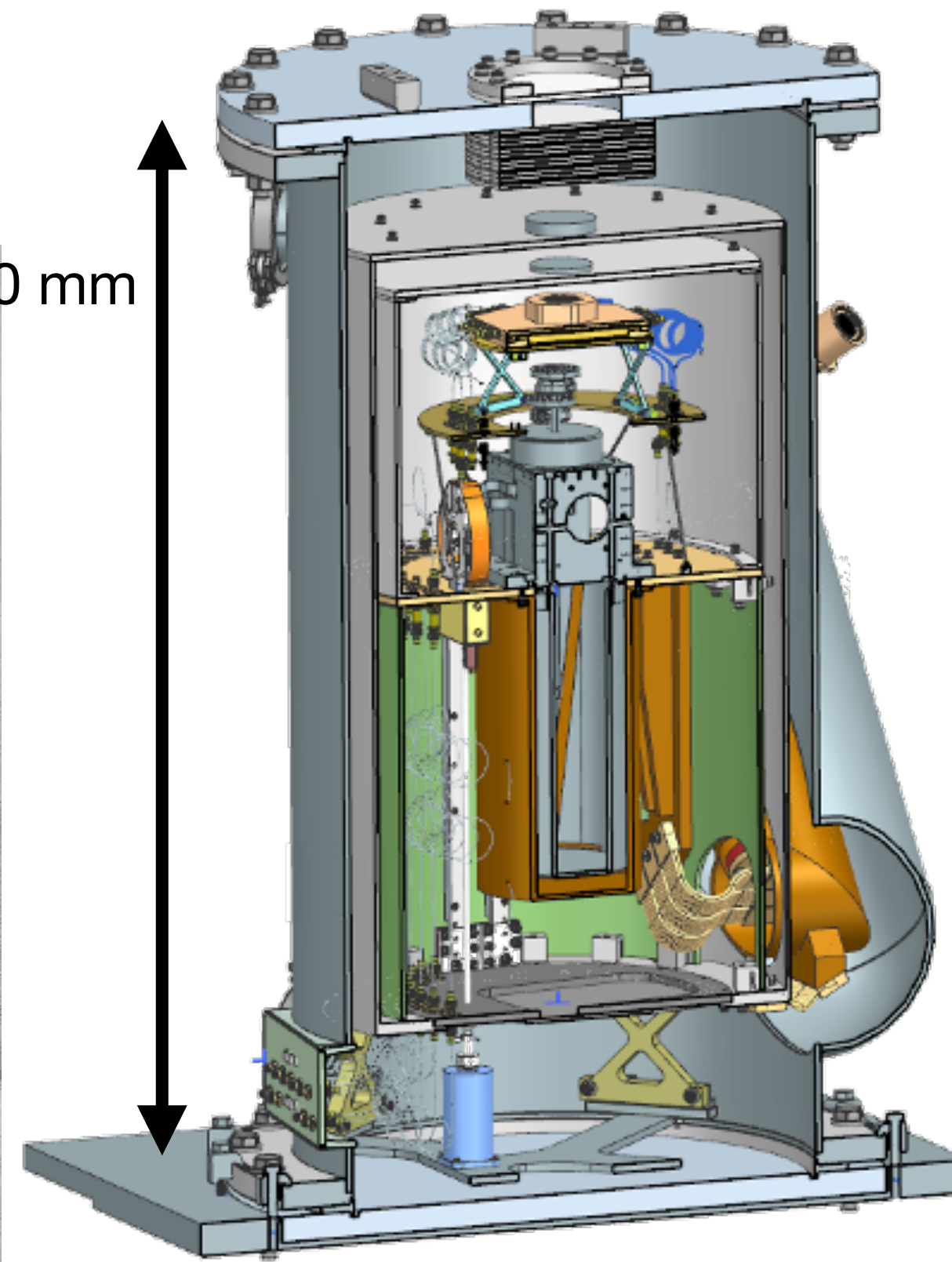
G. Smecher



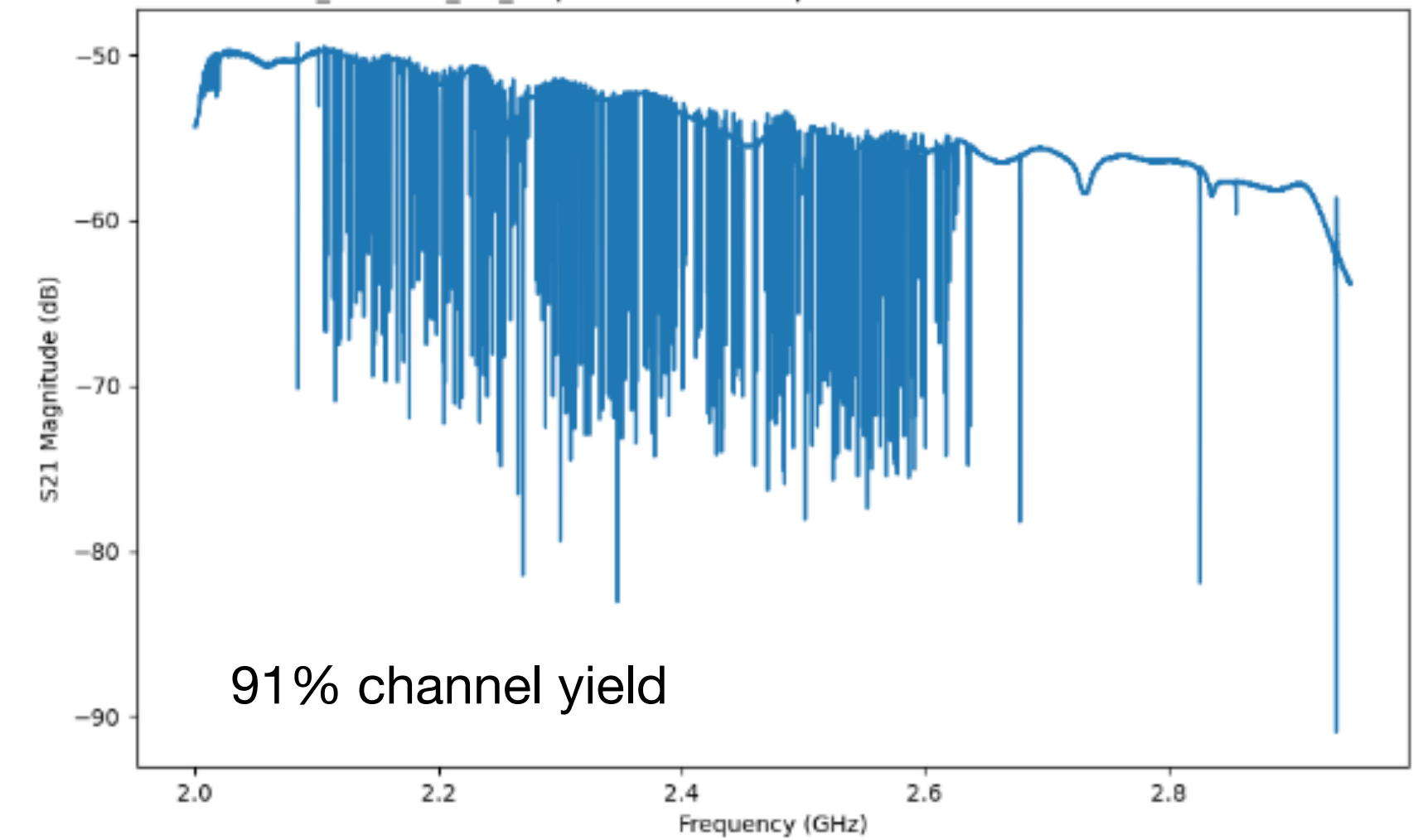
SPT-SLIM Hardware



880 mm



v2_anlw001_idc_stepdown VNA Sweep with ~ -42.0 dB total attenuation



A. Anderson



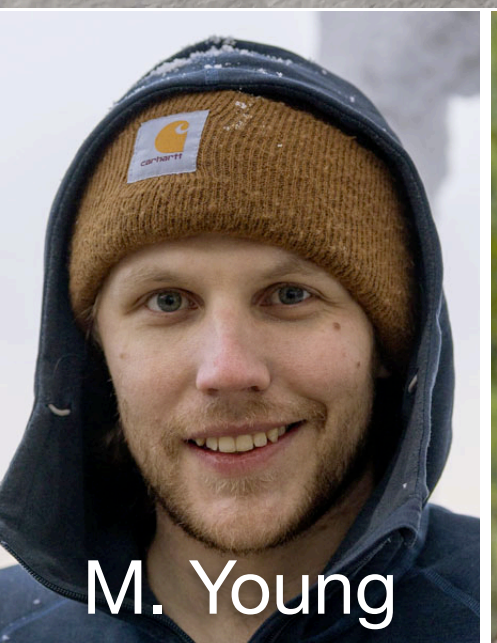
B. Benson



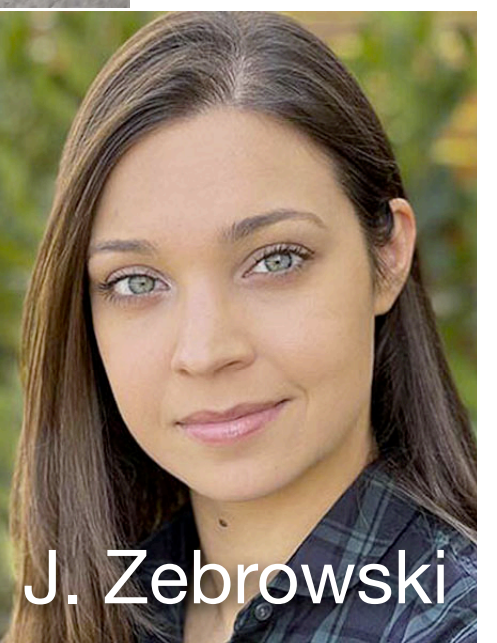
K. Dibert



D. Mitchell



M. Young



J. Zebrowski

SPT-3G+: High-frequency CMB

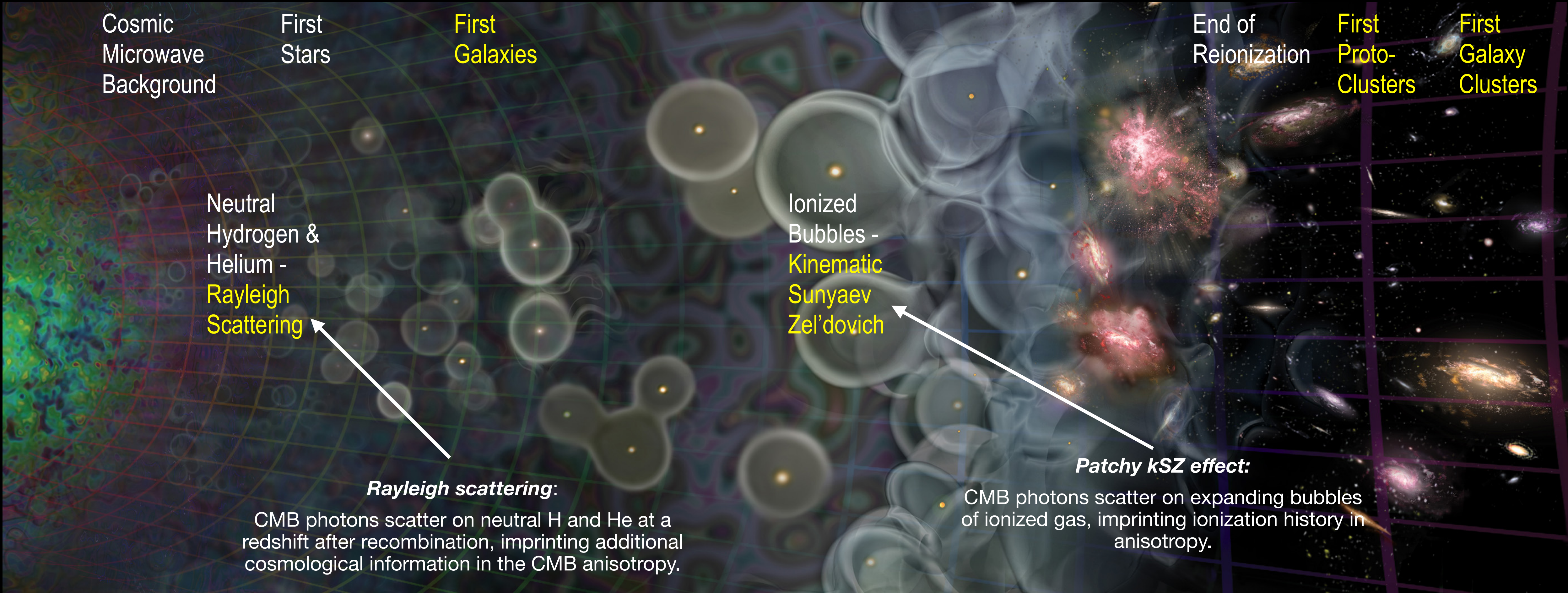
Reionization, Recombination, and the CMB

Recombination

Reionization



Redshift	~1100	~30	~15		~6	~4	~2.5
Age	~0.4 Myr	~100 Myr	~300 Myr		~1 Gyr	~1.6 Gyr	~2.7 Gyr



Neutral Hydrogen & Helium - Rayleigh Scattering

Rayleigh scattering:
CMB photons scatter on neutral H and He at a redshift after recombination, imprinting additional cosmological information in the CMB anisotropy.

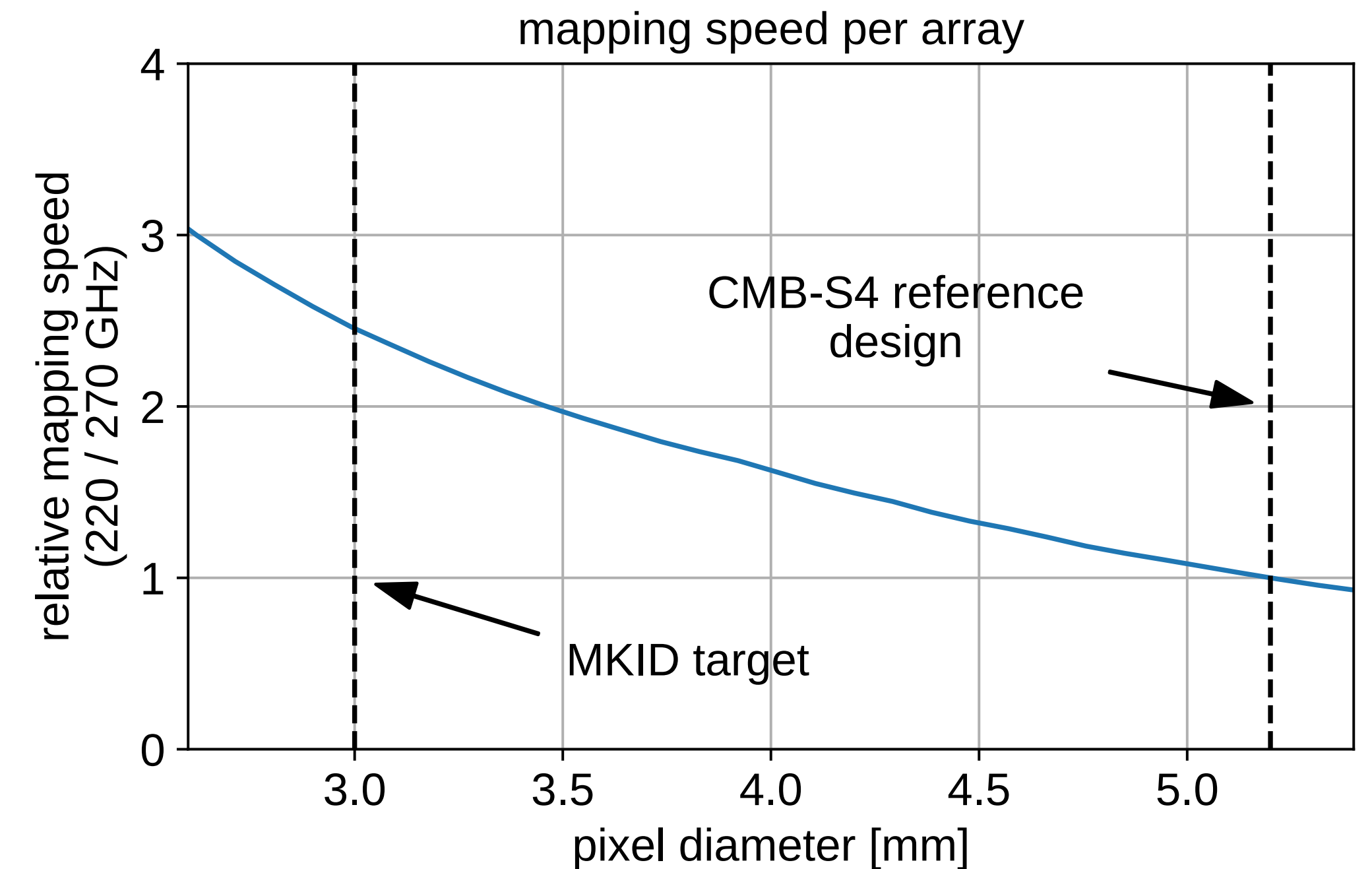
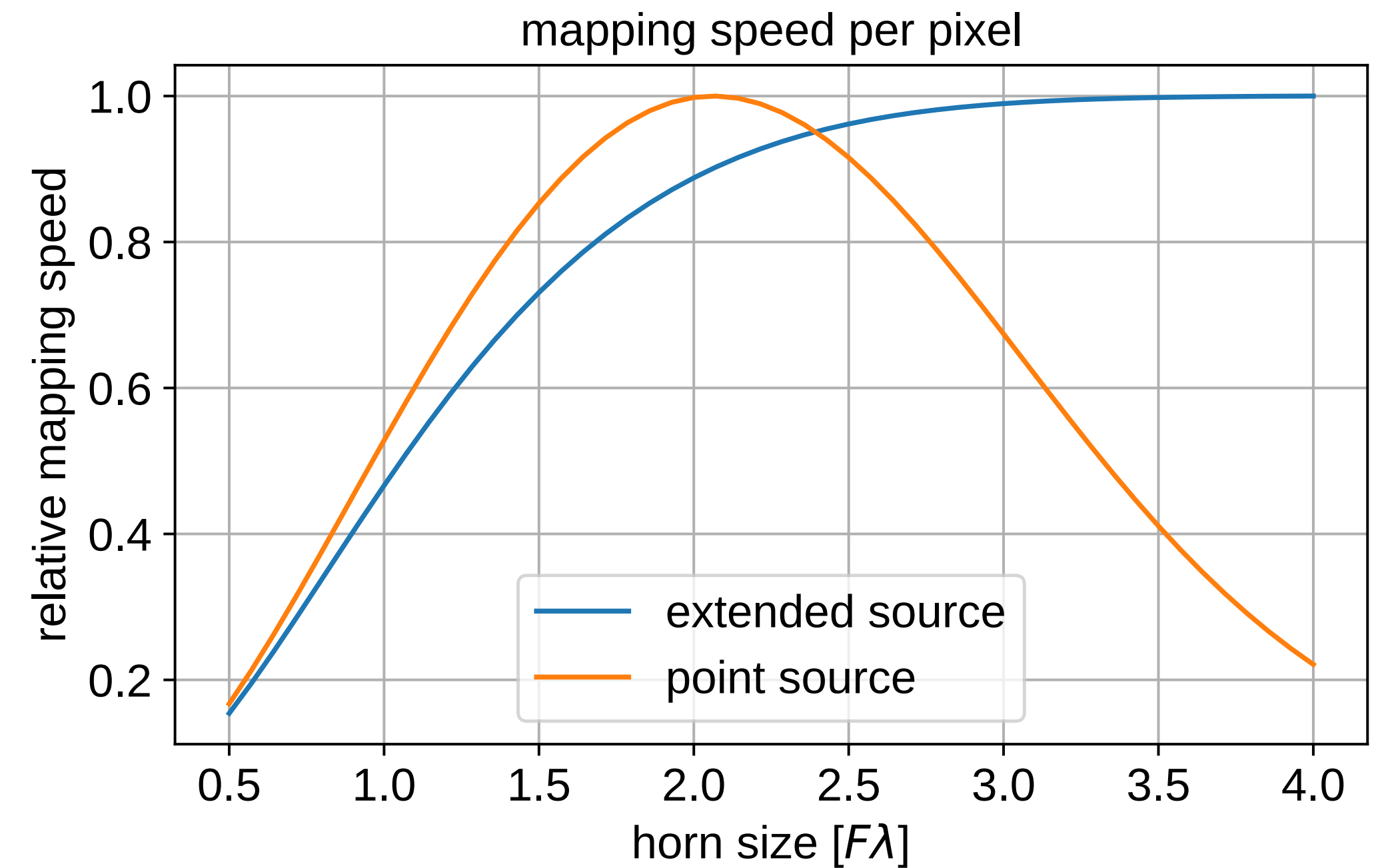
Ionized Bubbles - Kinematic Sunyaev Zel'dovich

Patchy kSZ effect:
CMB photons scatter on expanding bubbles of ionized gas, imprinting ionization history in anisotropy.

Secondary CMB anisotropies and sources targeted by upcoming CMB experiments (Simons Observatory, CCAT-prime, CMB-S4, ...)

Mapping Speed of Arrays

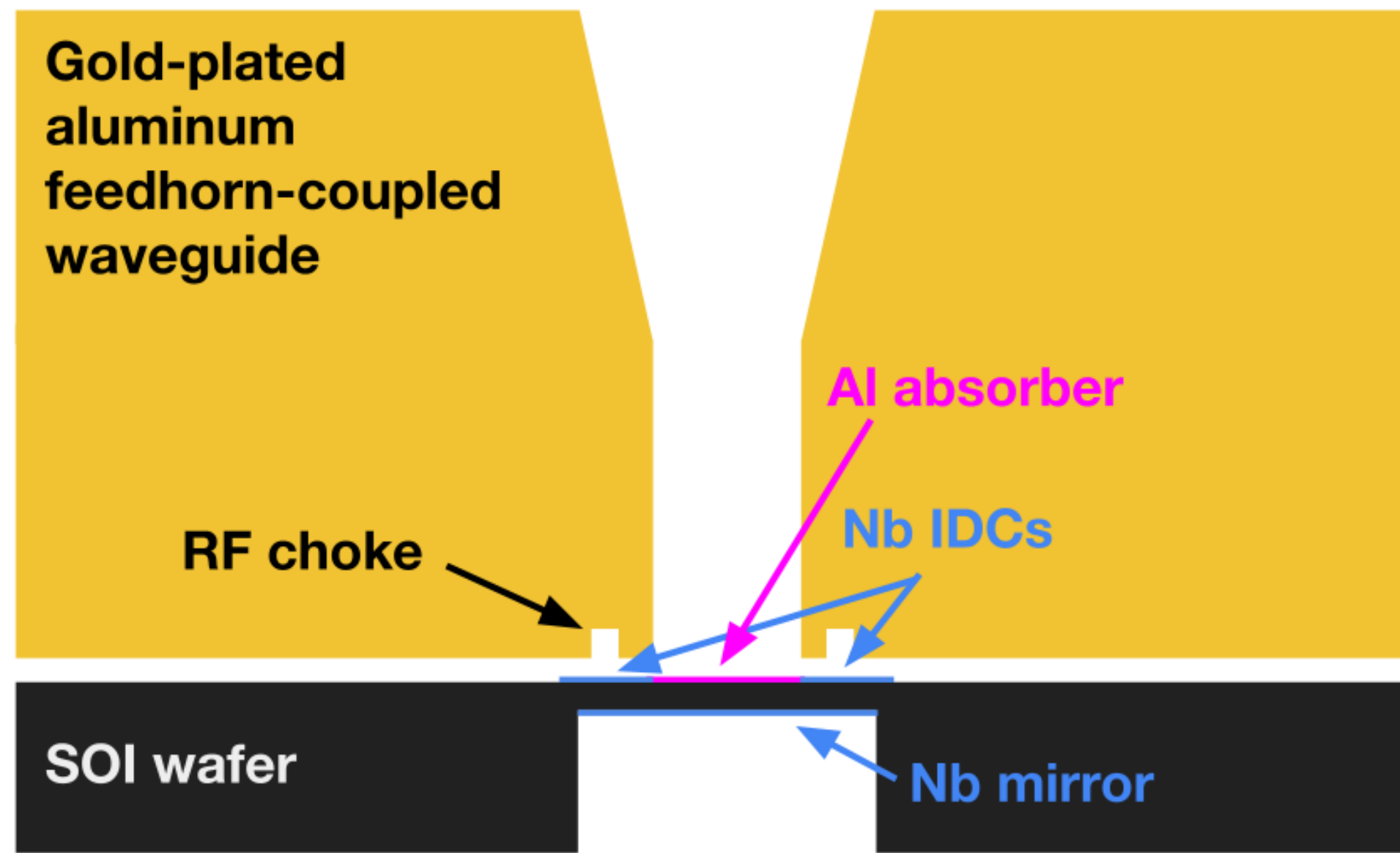
- Mapping speed *per pixel* optimized for pixel sizes $> 2F\lambda$
- But smaller pixels enable more uncorrelated detectors per array, so mapping speed *per array* tends to be maximized for small pixels down to $0.5-1F\lambda$
- Improvement of KIDs vs. TESs due to detector density is greatest at frequencies $> 150\text{GHz}$
- **Example:**
 - CMB-S4 220/270 GHz dichroic band is limited to ~ 2000 detectors / wafer
 - $> 3x$ increase in sensitivity possible by moving to denser arrays with smaller pixels



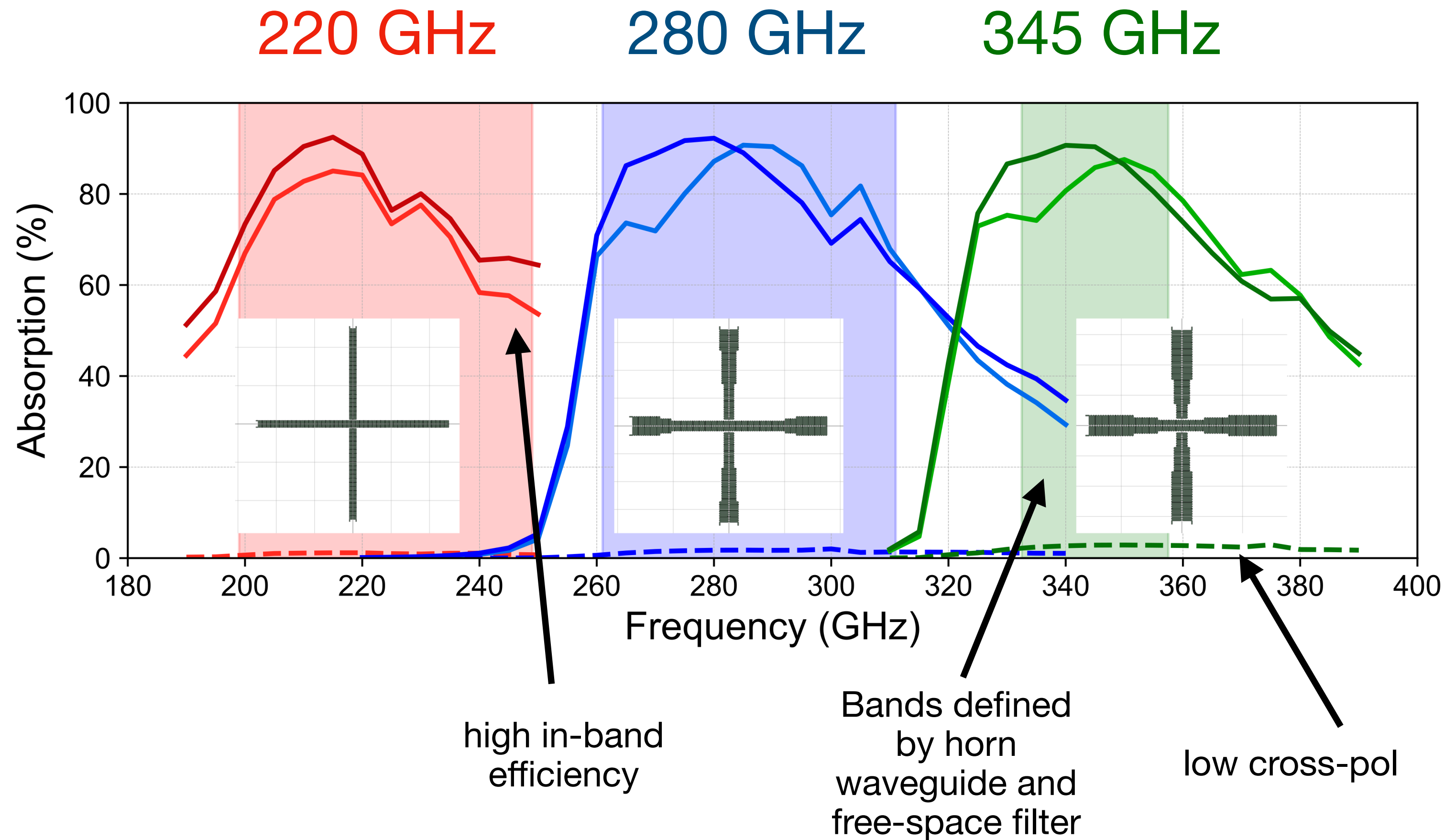
SPT-3G+ Detector Architecture



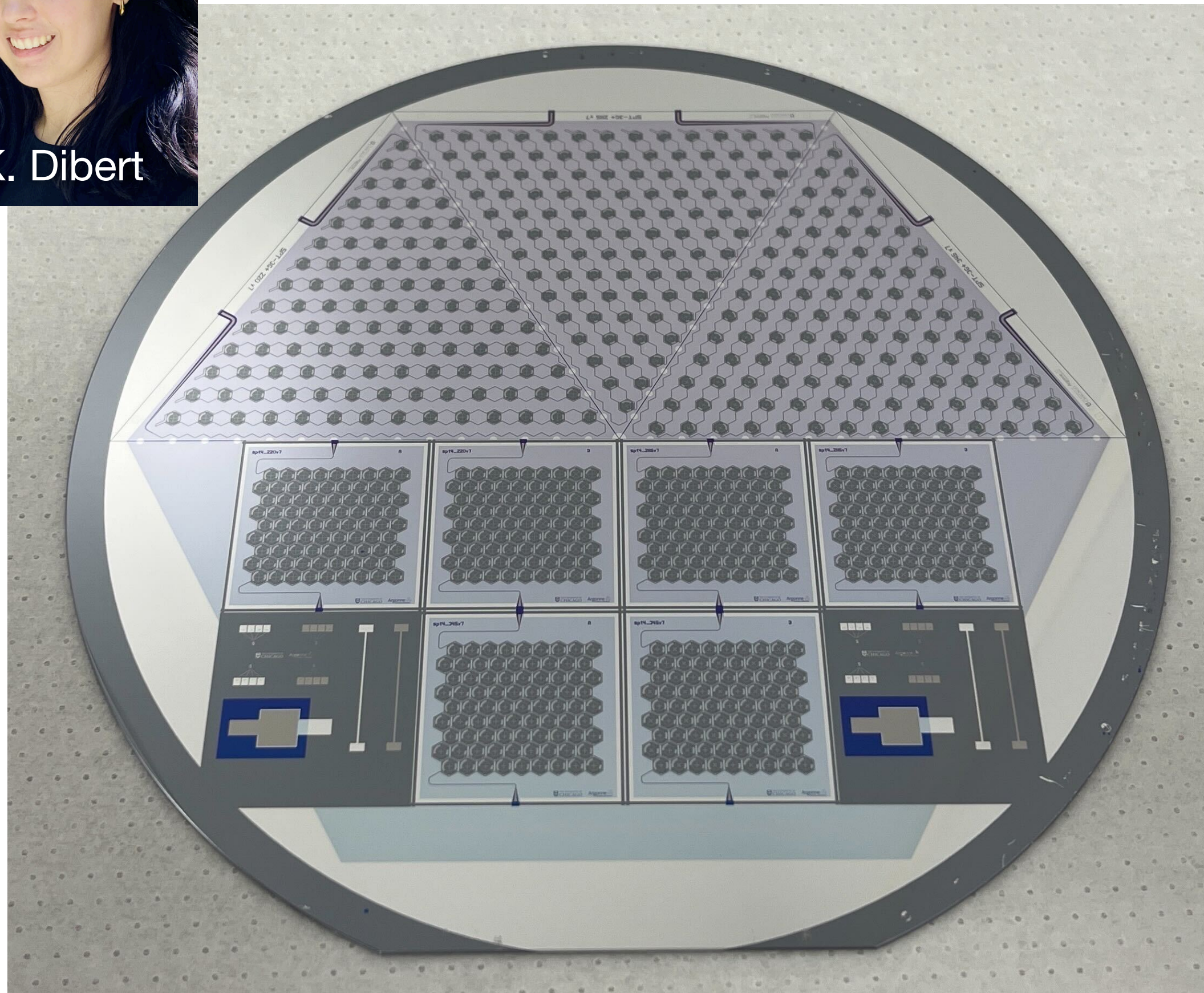
Bandpasses



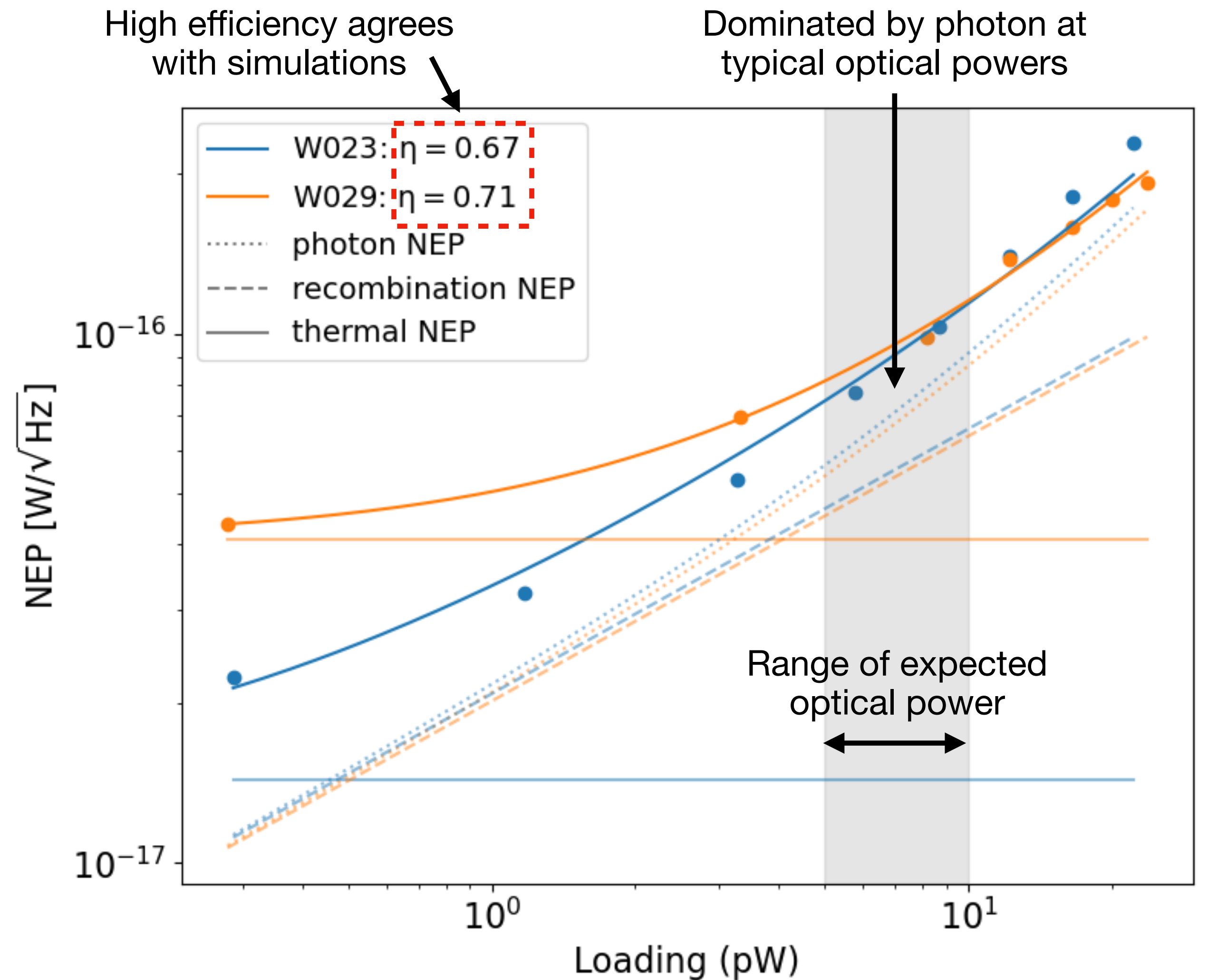
K. Dibert, P. Barry



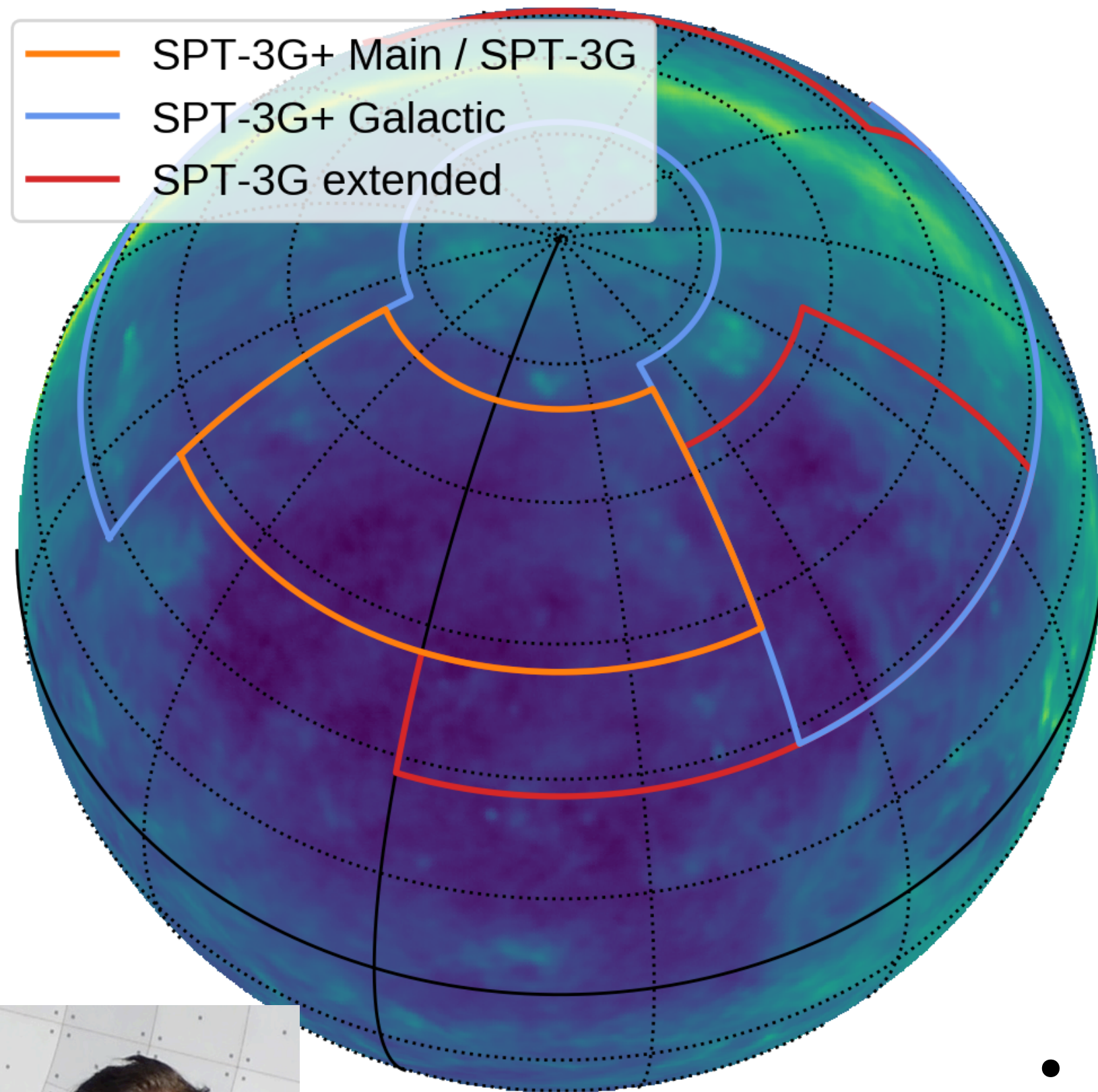
SPT-3G+ Detector Prototyping and Testing



220 GHz detector test wafer fabricated at Argonne
(Karia Dibert, Tom Cecil)



SPT-3G+ Survey Concept

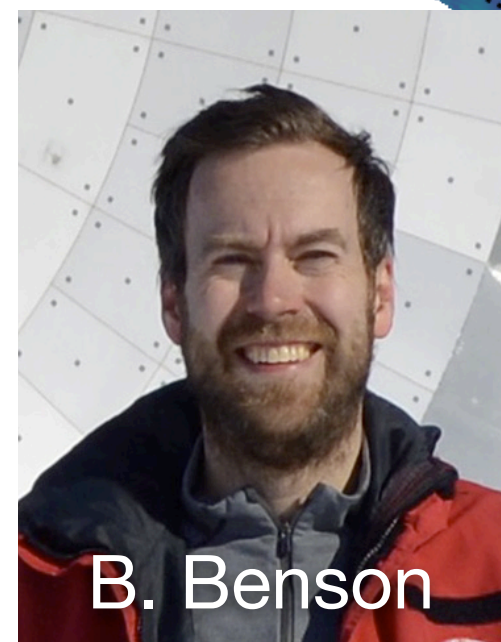


3G

3G+

Band	Survey Depth [$\mu\text{K-arcmin}$]			Resolution [arcmin]
	Main (1500 deg ²)	Extended (3000 deg ²)	Galactic (7000 deg ²)	
95 GHz	3.0	9	—	1.5
150 GHz	2.2	9	—	1.2
220 GHz	8.8	13	—	1.0
220 GHz	2.9	—	14	0.8
285 GHz	5.6	—	28	0.6
345 GHz	28	—	170	0.5

- Reuse the existing SPT primary and secondary optics, with a new cryostat containing KID-based CMB detectors at 220, 285, and 345 GHz.
- Continue to observe existing SPT “Main” field, but add 7000 deg² “Galactic” survey during the austral summer months



B. Benson

Conclusions

- Pushing the *density frontier* of mm-wavelength detectors using KIDs opens up new science opportunities beyond CMB-S4.
- Fermilab, together with Chicago and Argonne, is leading the development of two new cameras based on high-density KIDs:
 - SPT-SLIM - Pathfinder for mm-wavelength line intensity mapping
 - SPT-3G+ - High-frequency CMB
- SPT-SLIM deploys next year!