LuSEE-Night

A Pathfinder Radio Telescope on the Far Side of the Moon

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Fermilab, 17 May 2024

Brookhaven National Laboratory

U.S. DEPARTMENT OF ENERGY

@BrookhavenLab

Overview

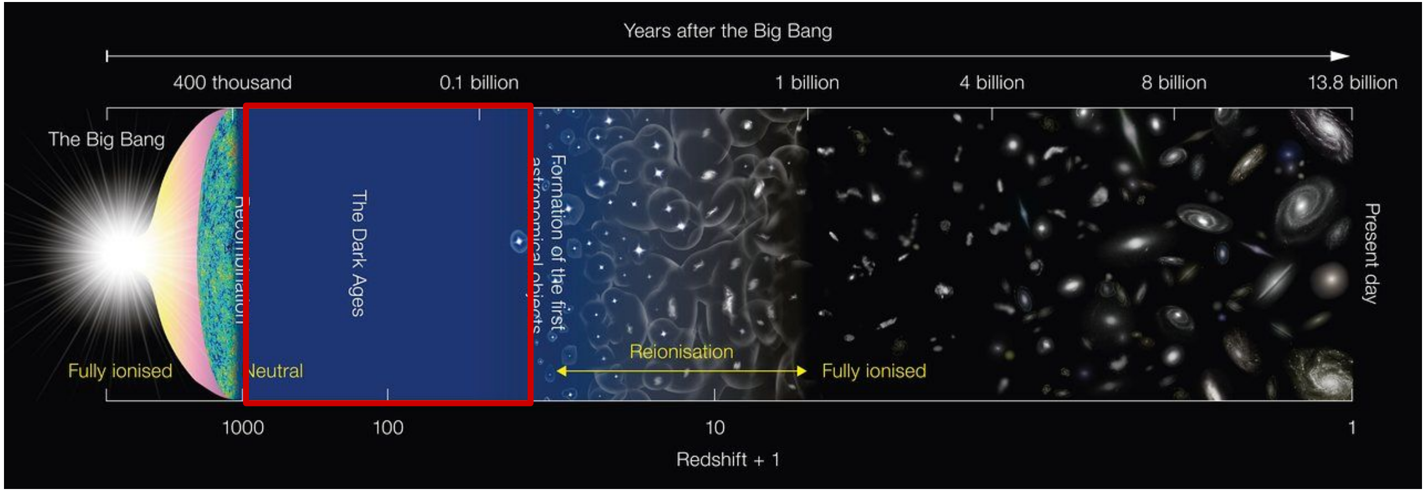
- Motivation: Dark Ages Cosmology

  In fiction, a **MacGuffin** (sometimes **McGuffin**) is an object, device, or event that is necessary to the plot and the motivation of the characters, but insignificant, unimportant, or irrelevant in itself.
- History and context of radio instrumentation on or around the Moon
- LuSEE-Night

The Dark Ages

- Period between CMB and Cosmic Dawn
- No compact radiation sources → linear physics
- No compact radiation sources → no optical radiation to observe

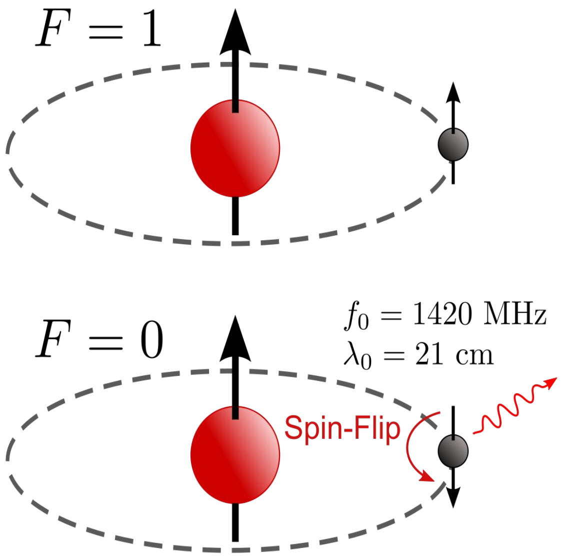
One of the least constrained frontiers of modern cosmology!



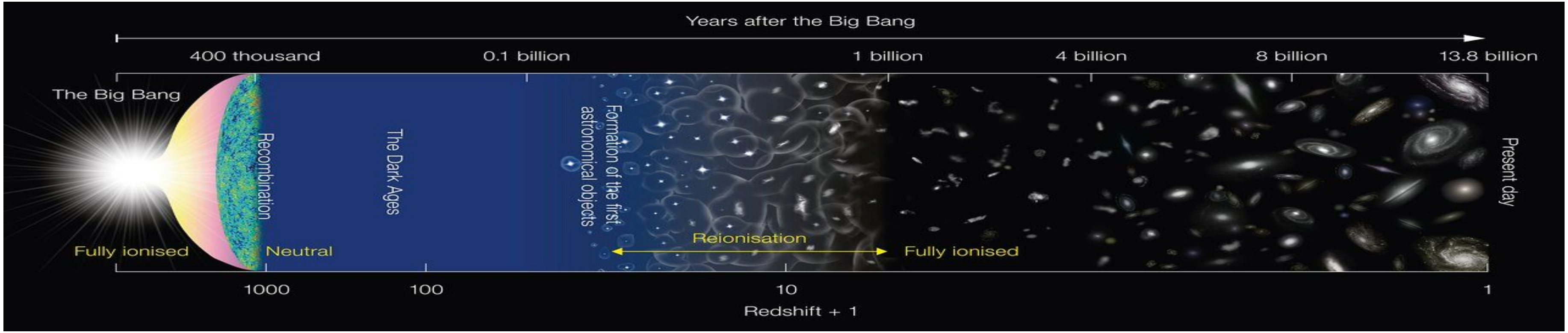
ESO & NAOJ

21cm emission

- Hyperfine transition in neutral hydrogen at ν=1420MHz, λ=21.1cm;
- This is the **only** transition around -- if you see a line at 710MHz, it is a z=1 galaxy;
- (not true in optical)
- Universe is mostly hydrogen (75%)
- Hydrogen is in various states and 21cm emission dominated by different mechanism across the cosmic ages
- NB: you need atomic hydrogen: neither ionized nor molecular hydrogen work



21cm shines differently at different epochs



Dark Ages

$20 \lesssim z \lesssim 150$

- Pristine primordial density field
- We see 21cm in emission / absorption against CMB light from the global density field

Epoch of Reionization

$6 \lesssim z \lesssim 20$

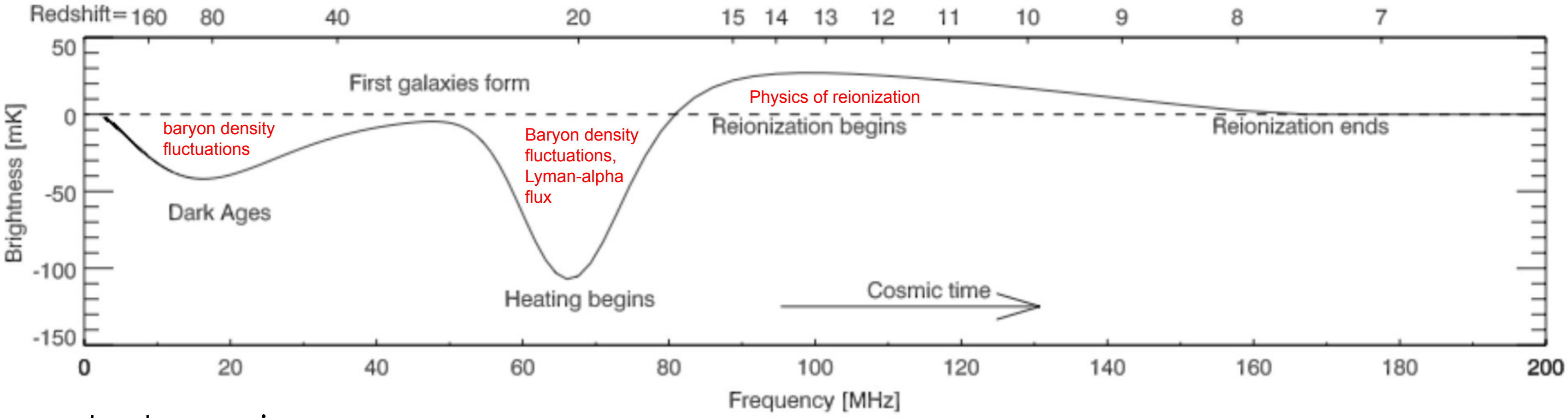
- First stars and galaxies are reionizing universe
- Large bubbles of ionized gas among neutral medium
- Signal driven by astrophysics
- Non-DOE science

Low redshift

$z \lesssim 6$

- Universe is reionized
- pockets of neutral hydrogen in galaxies
- Very similar science to standard galaxy surveys
- We don't aim to go after individual galaxies

Physics of the 21cm monopole



We see hydrogen in contrast with CMB. There can be net absorption or net emission:

$$\delta T_b = \frac{T_S - T_R}{1+z}(1 - e^{-\tau_\nu})$$
$$\approx \frac{T_S - T_R}{1+z}\tau$$

Three processes matter:

- absorption / emission from CMB
- collisions with other particles
- Wouthuysen-Field effect (resonant scattering of Ly-alpha photons)

- **$200 \lesssim z \lesssim 1100$**: The residual free electron fraction left after recombination allows Compton scattering to maintain thermal coupling of the gas to the CMB, setting $T_K = T_\gamma$. The high gas density leads to effective collisional coupling so that $T_S = T_\gamma$ and we expect $\bar{T}_b = 0$ and no detectable 21 cm signal.
- **$40 \lesssim z \lesssim 200$**: In this regime, the gas cools adiabatically so that $T_K \propto (1+z)^2$ leading to $T_K < T_\gamma$ and collisional coupling sets $T_S < T_\gamma$, leading to $\bar{T}_b < 0$ and an early absorption signal. At this time, T_b fluctuations are sourced by density fluctuations, potentially allowing the initial conditions to be probed [32, 22].
- **$z_\star \lesssim z \lesssim 40$**: As the expansion continues, decreasing the gas density, collisional coupling becomes ineffective and radiative coupling to the CMB sets $T_S = T_\gamma$, and there is no detectable 21 cm signal.

(Pritchard and Loeb review, 2011)

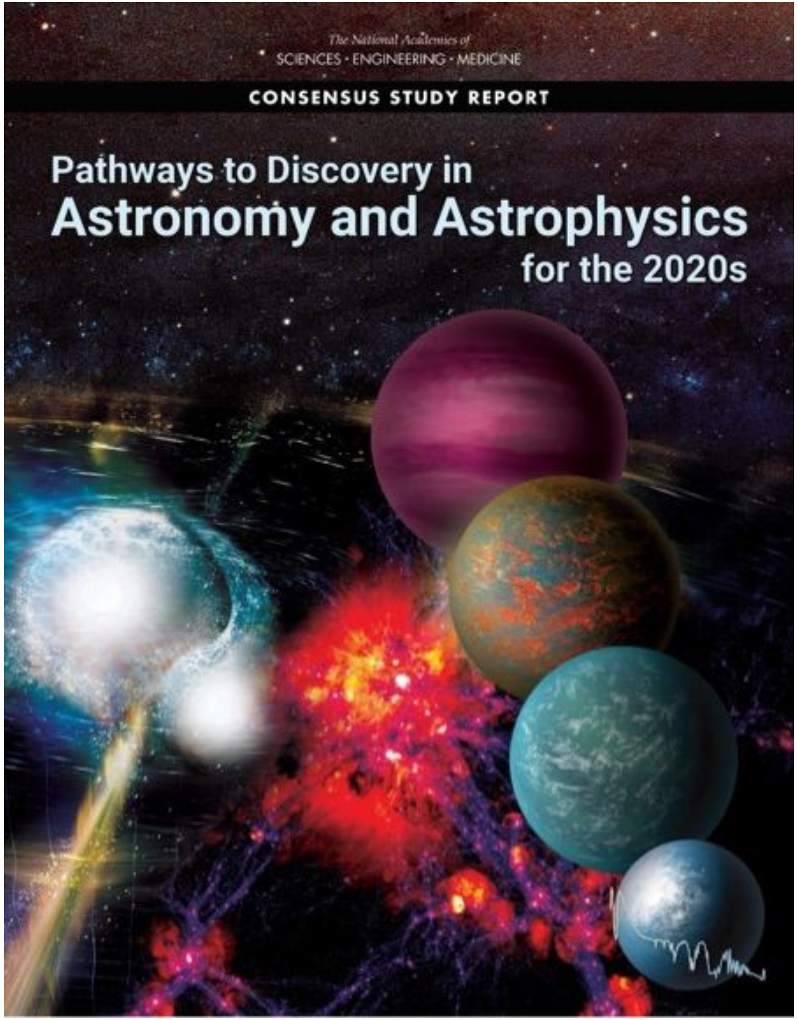
Dark Ages Monopole: A cosmic calorimeter

- The size of the dark ages through is sensitive to any non-standard energy injection or sink in the early universe.
- Measuring it would be confirmation of our understanding of thermodynamics in the post-CMB era
- Deviations would imply something from a large menu of options:
  - dark matter-baryon scattering
  - millicharged dark matter
  - dark matter annihilation
  - axions
  - neutrino decay
  - charge sequestration
  - quark nuggets
  - dark photons
  - interacting early dark energy
  - …

Decadal Survey on Astronomy & Astrophysics,
Panel on Cosmology p.258

DISCOVERY AREA: THE DARK AGES AS A COSMOLOGICAL PROBE

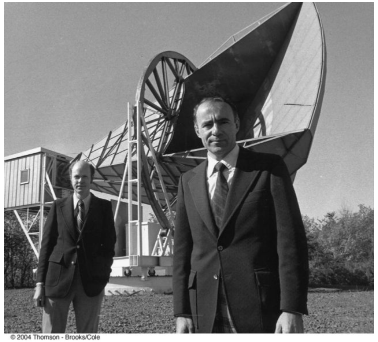
“The panel sees 21 cm and molecular line intensity mapping of the Dark Ages and reionization era as both the discovery area for the next decade and as the likely future technique for measuring the initial conditions of the universe in the decades to follow.



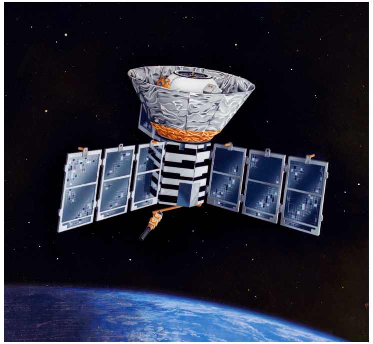
You eat the elephant in small pieces

CMB

- Monopole discovered 1965 by Penzias and Wilson (Nobel Prize in 1978)



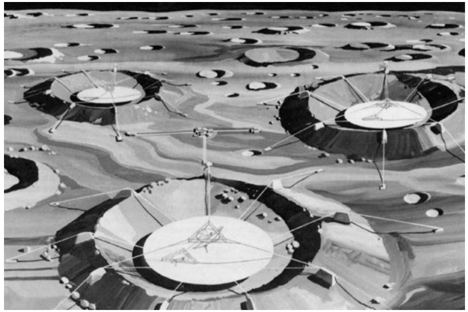
- Fluctuations discovered 1990 by COBE satellite (Smoot’s Nobel Prize in 2006)



Dark Ages

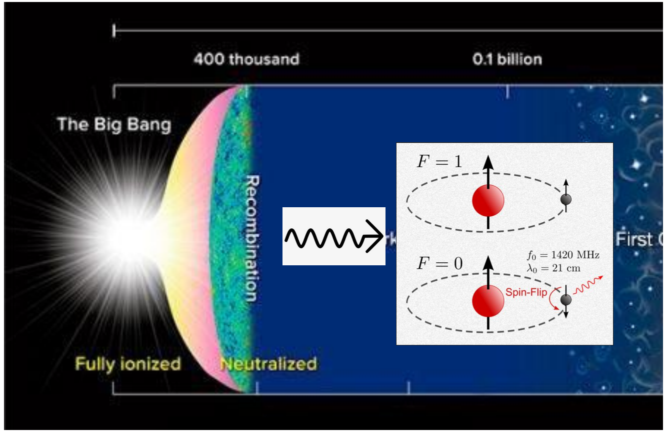
- Monopole detected 2030s (?)

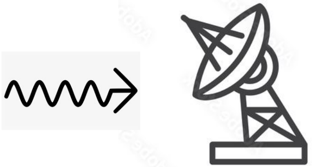
- Fluctuations detected 2060(?)



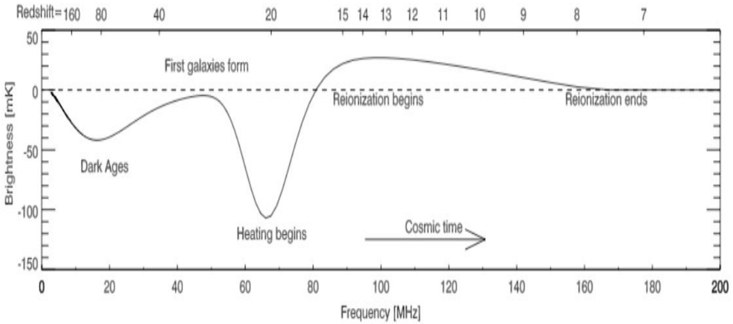
Science Observable: monopole

CMB "backlight" Dark Ages Hydrogen

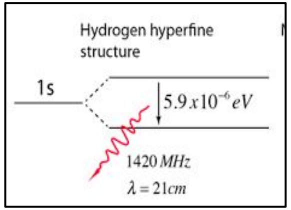




We observe the resulting absorption/emission spectrum



"Constraining the unexplored period between reionization and the dark ages with observations of the global 21 cm signal", J. Pritchard and A. Loeb 2010



Cold hydrogen gas sometimes absorbs, sometimes emits a 21cm photon depending on spin temperature

- ➔ Dark Ages spectrum is a ***cosmic calorimeter***
- ➔ Measures something adding or taking energy from hydrogen fluid

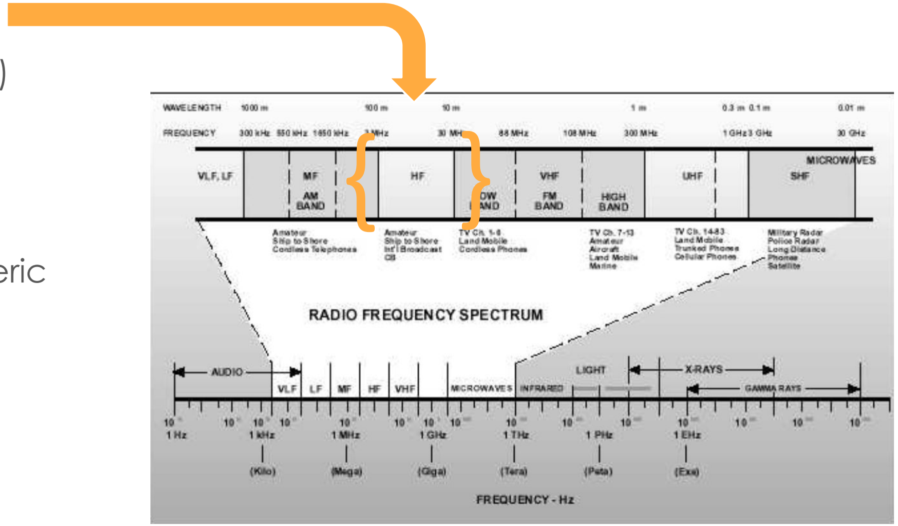
What do we need to measure this?

Objective:

- observe the sky at low radio frequencies (0.5 – 50MHz) where the Dark Ages signal lives

Impossible from Earth:

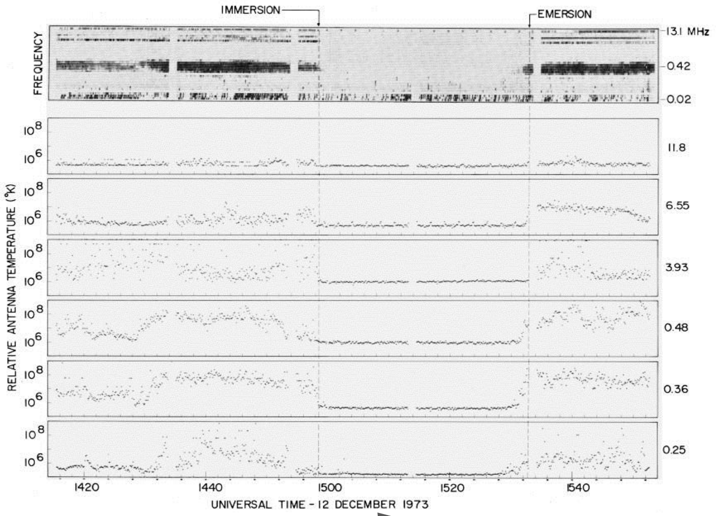
- RF contamination from anthropogenic and atmospheric sources
- Ionosphere absorbs, emits, and refracts at these frequencies

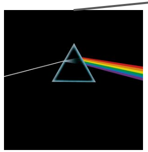


Solution:

- Telescope on the Moon's far side is protected from terrestrial RFI
- The Moon has no atmosphere → no ionospheric distortions
- During the lunar night, solar RFI is also blocked
- Alternatives with their own advantages and disadvantages:
  - Lunar Orbit
  - Outer solar system

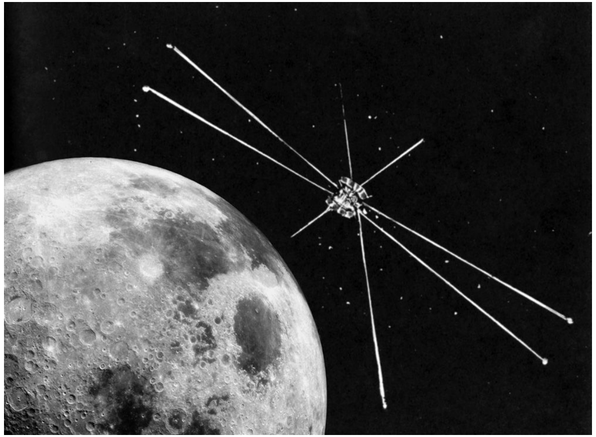
What do we know about low-frequency sky





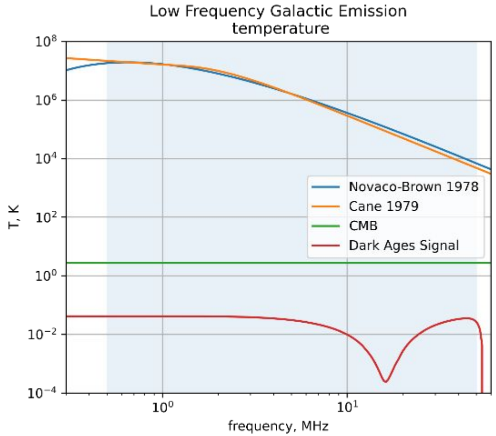
The Dark Side of the Moon by Pink Floyd released in March 1973

- We have empirical evidence that the far side of the moon is radio-quiet
- Caveat: in orbit, not on the surface



Radio-Astronomy-Explorer-2 was the largest spacecraft ever built with four 230m antennas

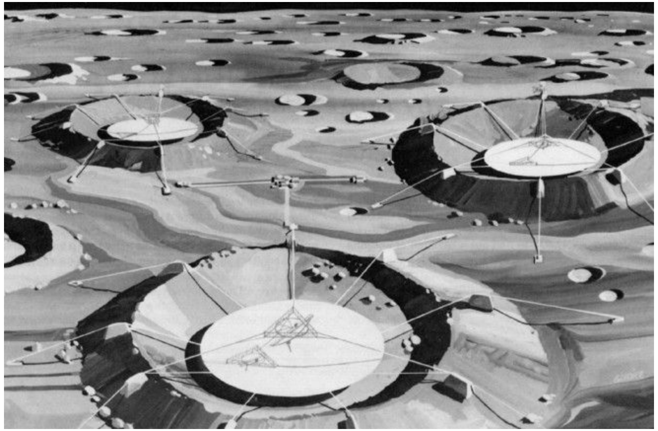
- Last real data on the low-frequency sky from the late 1970s
- Still state-of-the-art dataset to be superseeded



- Dark Ages signal is dominated by orders of magnitude brighter synchrotron emission from Galaxy

How to build a telescope on the Moon?

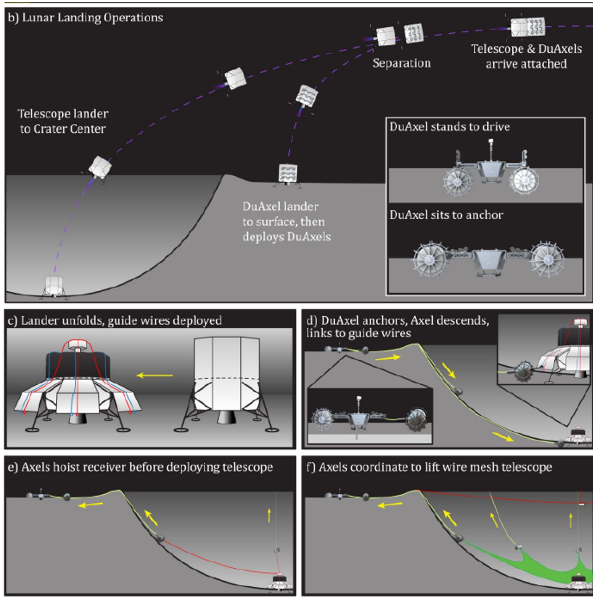
- People were day-dreaming about putting a telescope of the far side of the Moon since 1950s
- Richard Vondrak working at Apollo Science Operations Center proposed using lunar craters to build radio telescopes like the Arecibo Observatory in Puerto Rico.

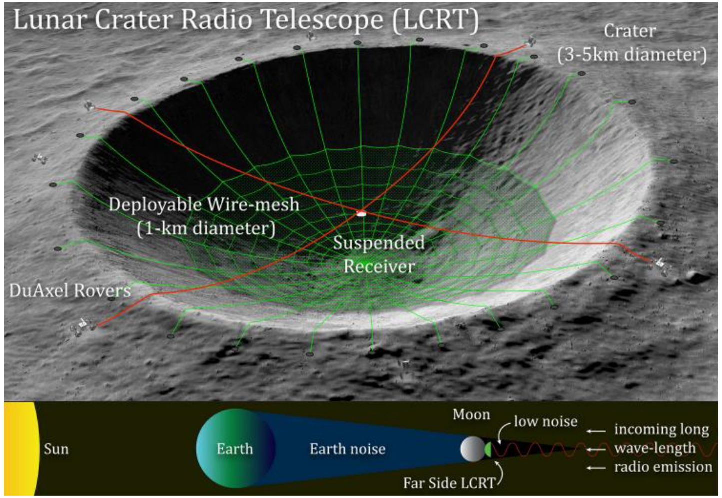


1950s concept from Apollo times: Arecibos in Lunar Crater

Arecibo (now decommissioned)

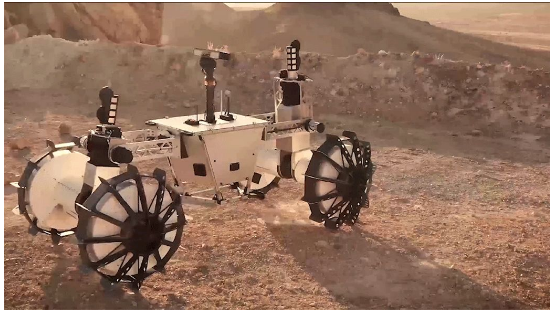
Lunar Crater Radio Telescope



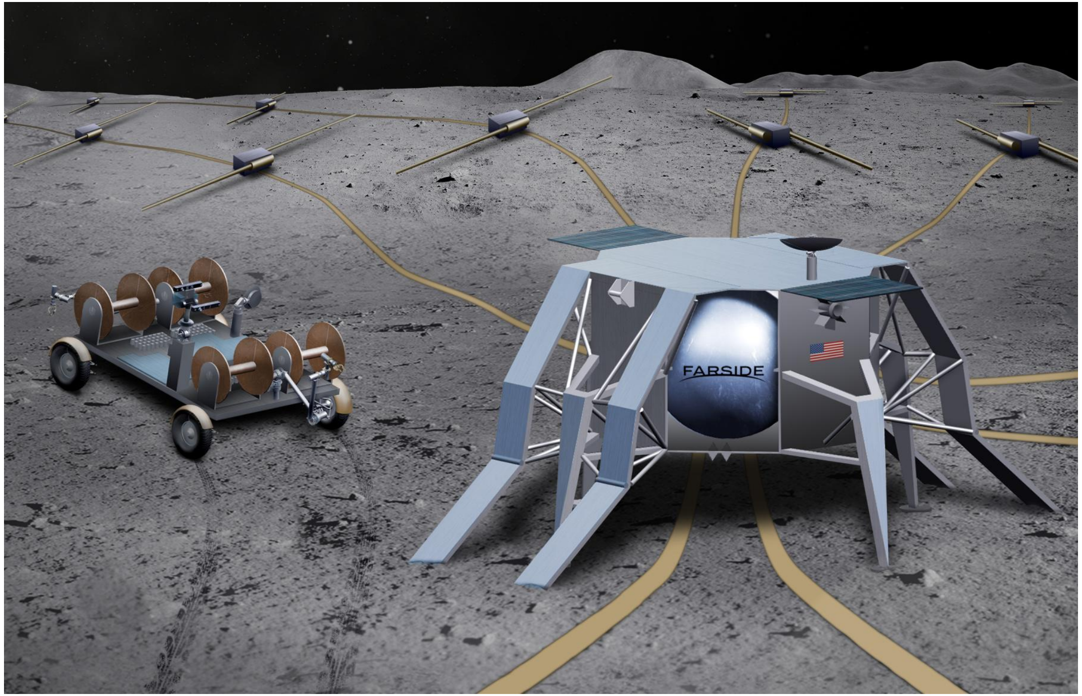


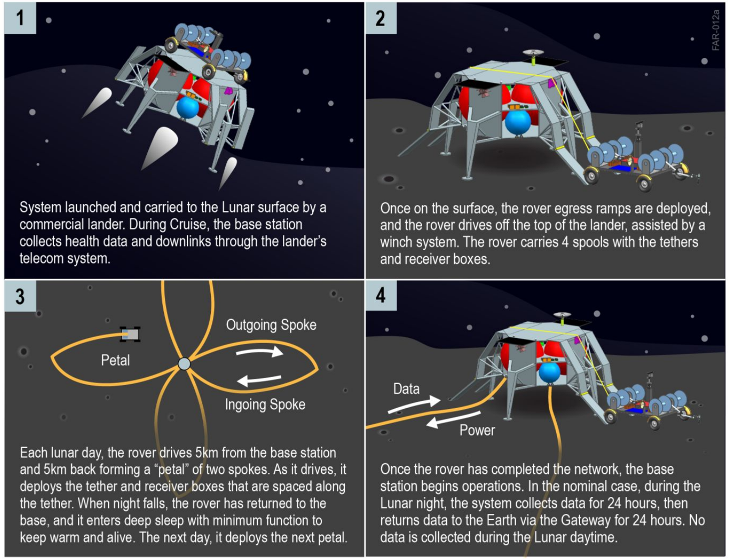
LCRT: modern incarnation of the crater concept
PI: Saptarshi Bandyopadhyay

DuAxels actually exist



FARSIDE:



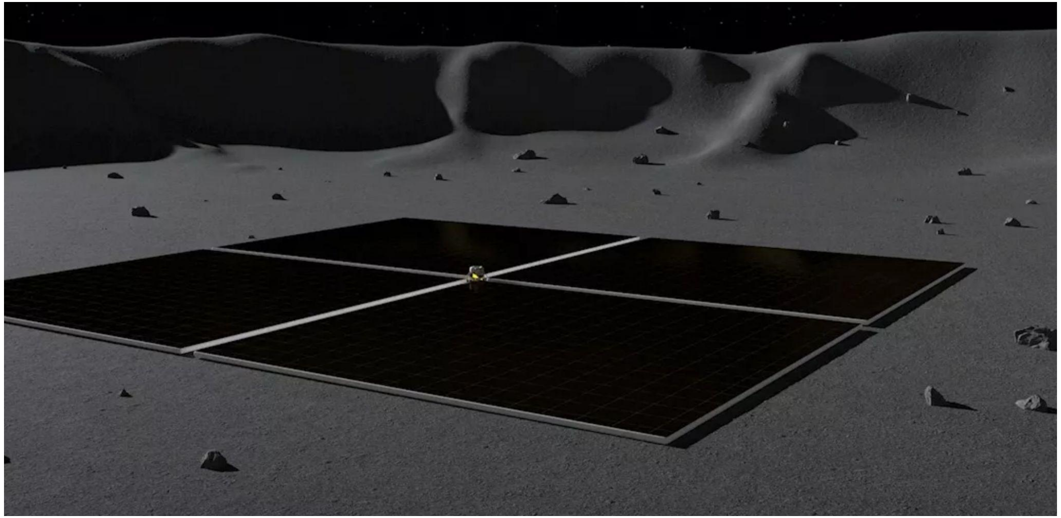


Lands and installs a 128-node interferometric array of dipole antennas

PI Jack Burns

ALO: Astronomical Lunar Observatory

- Inflatable 32x32 array
- 1-80 MHz, total cargo ~1500kg
- launch ~2027-2029
- PI: Marc Klein Wolt
- ESA funded conceptual design level, designed by ASTRON in Netherlands
- See Koopmans' talk



Motivated by the hierarchy problem….

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PAPER

A very high energy hadron collider on the Moon

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Keywords: high-energy physics, future colliders, beyond the standard model, high-temperature superconductors, lunar science, space travel, megastructures

Abstract

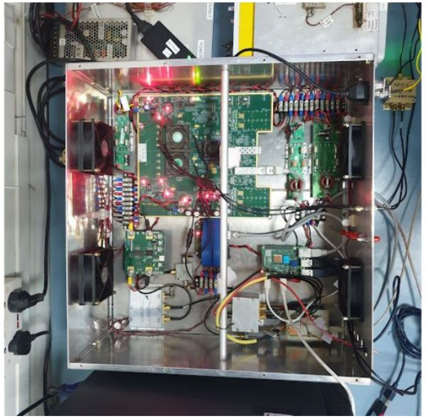
The long-term prospect of building a hadron collider around the circumference of a great circle of the Moon is sketched. A circular collider on the Moon (CCM) of ~11 000 km in circumference could reach a proton–proton center-of-mass collision energy of 14 PeV—a thousand times higher than the Large Hadron Collider at CERN—optimistically assuming a dipole magnetic field of 20 T. Several aspects of such a project are presented, including siting, construction, availability of necessary materials on the Moon, and powering, as well as a discussion of future studies and further information needed to determine the more concrete feasibility of each. Machine parameters and vacuum requirements are explored, and an injection scheme is delineated. Other unknowns are set down. Due to the strong interest from multiple organizations in establishing a permanent Moon presence, a CCM could be the (next-to-) next-to-next-generation discovery machine for high-energy particle physics and a natural successor to next-generation machines, such as the proposed future circular collider at CERN or a super proton–proton collider in China, and other future machines, such as a collider in the Sea, in the Gulf of Mexico. A CCM would serve as an important stepping stone toward a Planck-scale collider sited in our Solar System.

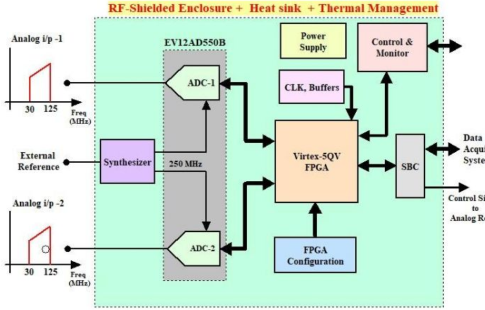

Bayt.com Blog:
Back to Reality: How to Ease into Work After a Relaxing Vacation

PRATUSH

- 40 - 200 MHz (55 - 110MHz baseline design)
- PI Mayuri S Rao
- Currently in design phase with ISRO
- PRATUSH antenna sits right on top of the spacecraft
- All electronics (PRATUSH + satellite) enclosed beneath the antenna
- Observations are made when in the shadow region on the lunar farside
- data is downlinked when in view of the Earth on the nearside.
- All subsystems are individually designed to be maximally smooth

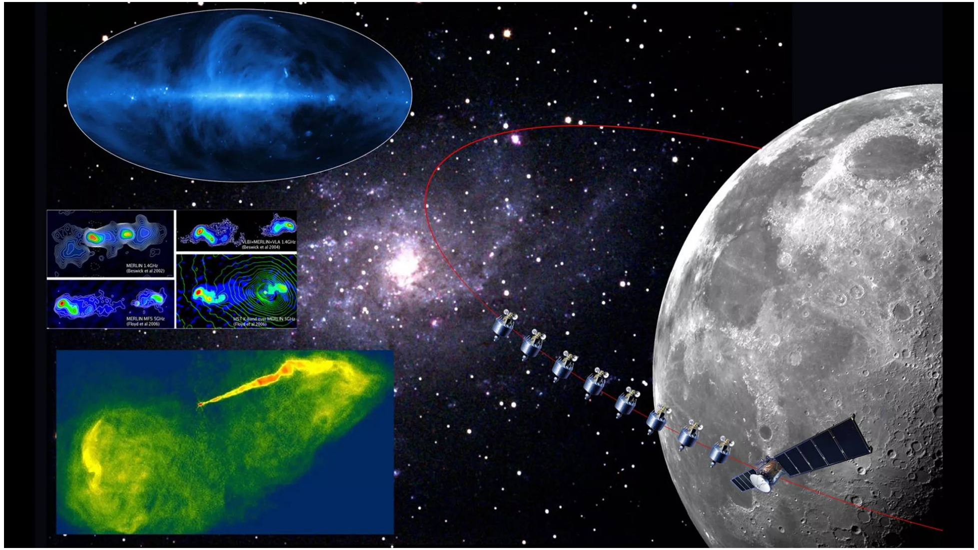
Monocone antenna design on shaped reflector (55 -110 MHz)





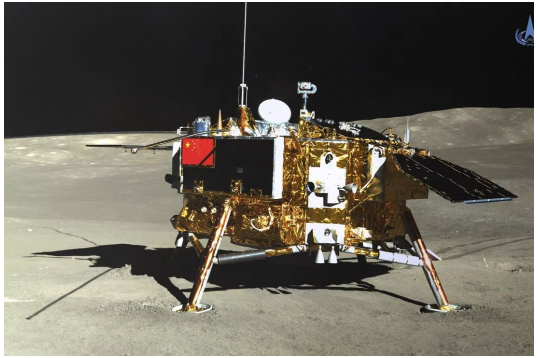
Hongmeng Project:
Discovering Sky at the Longest Wavelength

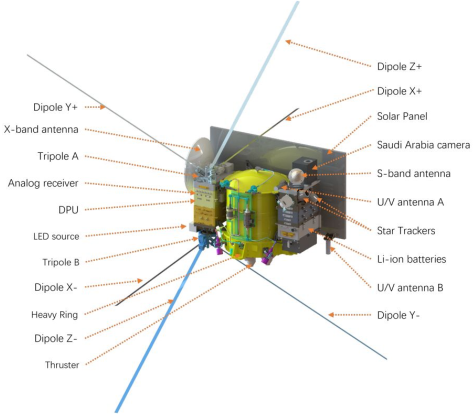
- **Recently selected**
- PI Xuelei Chen
- Mother + 8 daughter satellites
- tunable receiver below 30MHz
- A two satellite pathfinder did not manage to get fringes last year due to failure of one of the satellites

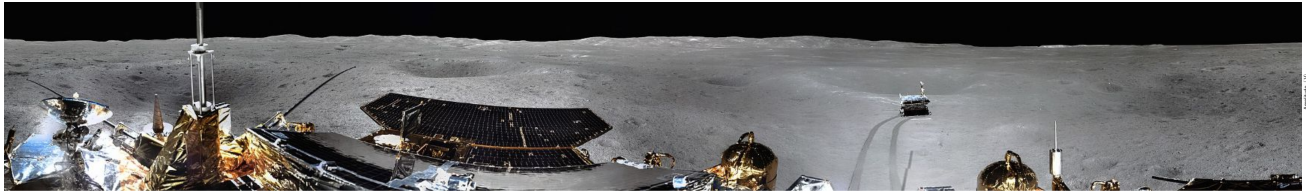


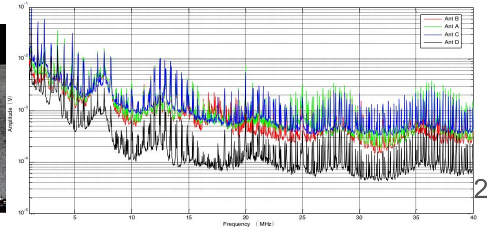
Chang’e-4

- Chinese Lander to the Far Side of the Moon
- Achieved first soft landing on the far side of the Moon, on 3 January 2019
- Carried Low. Frequency Radio Spectrometer (LFRS) on lander
- Carried NCLE (Netherlands-China Low-Frequency Explorer) on relay satellite
- Both instruments believed to be dominated by switching noise
- carried extra antenna to detect that, but no cigar





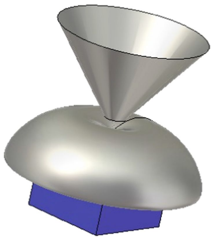




Do you want to be on the Moon or in orbit?

Advantages of orbit:

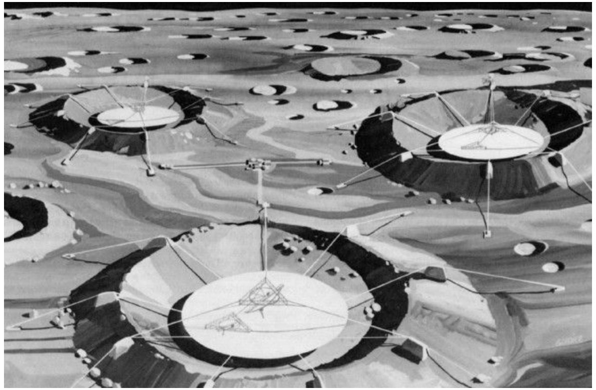
- system to simulate is compact and considerably smaller than wavelength
  - The entire system can be made geometrically simple and symmetric
- power management is easier: can re-charge every few hours
- orbit is cheaper than landing
- easy to implement rotation modulation (e.g. through tumble)



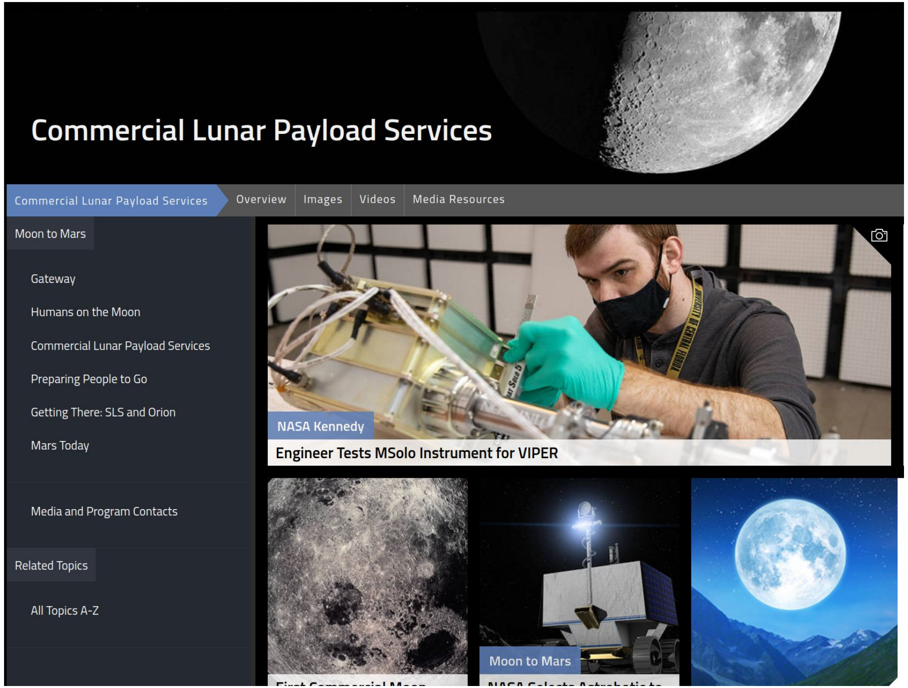
PRATUSH concept

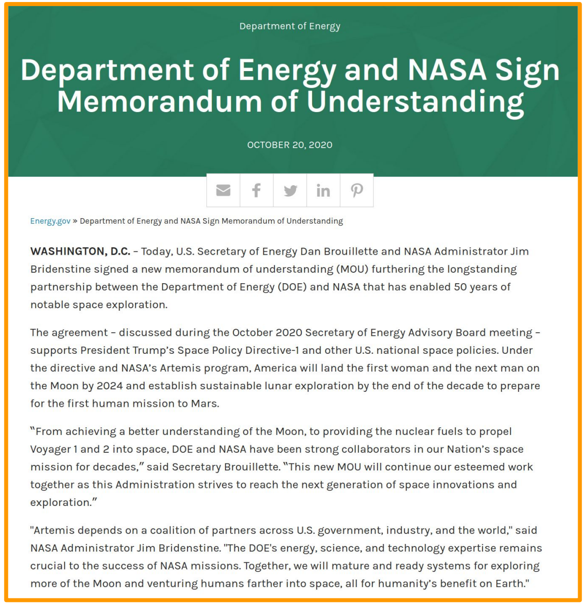
Advantages of lunar surface:

- Long, stable nights
- Stabler baselines for interferometers
- The effective size of the earth = size of the orbit of the most distant satellites >> R_{Earth} - the size of the truly shielded region much smaller
- Potentially quieter environment with less dense plasma



Meanwhile, in the US…





Department of Energy

Department of Energy and NASA Sign Memorandum of Understanding

OCTOBER 20, 2020

Energy.gov » Department of Energy and NASA Sign Memorandum of Understanding

WASHINGTON, D.C. – Today, U.S. Secretary of Energy Dan Brouillette and NASA Administrator Jim Bridenstine signed a new memorandum of understanding (MOU) furthering the longstanding partnership between the Department of Energy (DOE) and NASA that has enabled 50 years of notable space exploration.

The agreement – discussed during the October 2020 Secretary of Energy Advisory Board meeting – supports President Trump's Space Policy Directive-1 and other U.S. national space policies. Under the directive and NASA's Artemis program, America will land the first woman and the next man on the Moon by 2024 and establish sustainable lunar exploration by the end of the decade to prepare for the first human mission to Mars.

"From achieving a better understanding of the Moon, to providing the nuclear fuels to propel Voyager 1 and 2 into space, DOE and NASA have been strong collaborators in our Nation's space mission for decades," said Secretary Brouillette. "This new MOU will continue our esteemed work together as this Administration strives to reach the next generation of space innovations and exploration."

"Artemis depends on a coalition of partners across U.S. government, industry, and the world," said NASA Administrator Jim Bridenstine. "The DOE's energy, science, and technology expertise remains crucial to the success of NASA missions. Together, we will mature and ready systems for exploring more of the Moon and venturing humans farther into space, all for humanity's benefit on Earth."

- Increased desire for NASA/DOE cooperation as part of Artemis program
- CLPS looking for "oven-ready" payloads

CLPS - Commercial Lunar Payload Services

- advent of SpaceX taught NASA that a lot of money can be saved with commercial providers
- a similar program is developed to support Artemis program:
- Now at ~11 eligible vendors that bid on contracts
    - Number of vendors is expected to settle around ~a few.
- awarding 1-4 missions every year
- actively looking for payloads through PRISM (Payloads and Research Investigations on the Surface of the Moon) calls and otherwise
- they view payloads as essentially an exercise in how to run the contracting process
- ~~first launches expected in June 23 (from awards in 2019)~~ First launches finally happened in early 2024!

Landers selected in the first round: Peregrine (Astrobotic Technology), Nova-C (Intuitive Machines), Z-01 (OrbitBeyond)

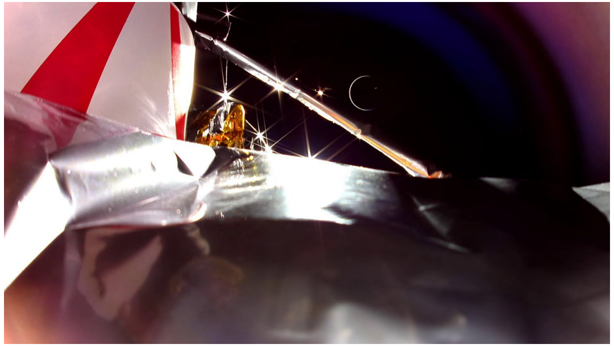
What happened to first CLPS launches.

First CLPS launch:

- January 2024
- Astrobotic, Peregrine
- Launched on an untested ULA rocker
- Rocked performed flawlessly, but mission was doomed almost immediately after separation due to a stuck valve

Second CLPS launch:

- February 2024
- Intuitive Machines, Odysseus,
- Managed to land… sideways
- Carries a radio-instrument called ROLSES
- Many things went wrong with this mission, alm from the get-go
  - none of them was totally catastrophic
  - the company learned an enormous lot



NASA-DOE LuSEE partnership

- UC Berkeley's Space Sciences laboratory (SSL) was chosen to provide a payload for a 2024 CLPS mission.
- Original LuSEE (Lunar Surface Electromagnetics Experiment) targeted plasma science, but incorporated a 20MHz RF spectrometer:
  - concept heavily relied on reusing existing spare parts from Parker Solar Probe
- A synergy with DOE's 21cm cosmology program was realized, and the mission was extended to include a dedicated, multi-night instrument dedicated to Dark Ages science
- Original concept split into 2+1 parts:
  - LuSEE-Lite - focus on day time solar science on CP-12
  - LuSEE-Night - focus on night-time radio science on CS-3
  - Far field calibration "service provision" on CS-4

Parker Solar Probe

Overarching Science Goals

LuSEE-Night is a pathfinder for the Dark Ages Science from the far-side of the Moon:

Goal 1: ***Establish the Lunar surface as a viable observatory for low-frequency radio astronomy;***

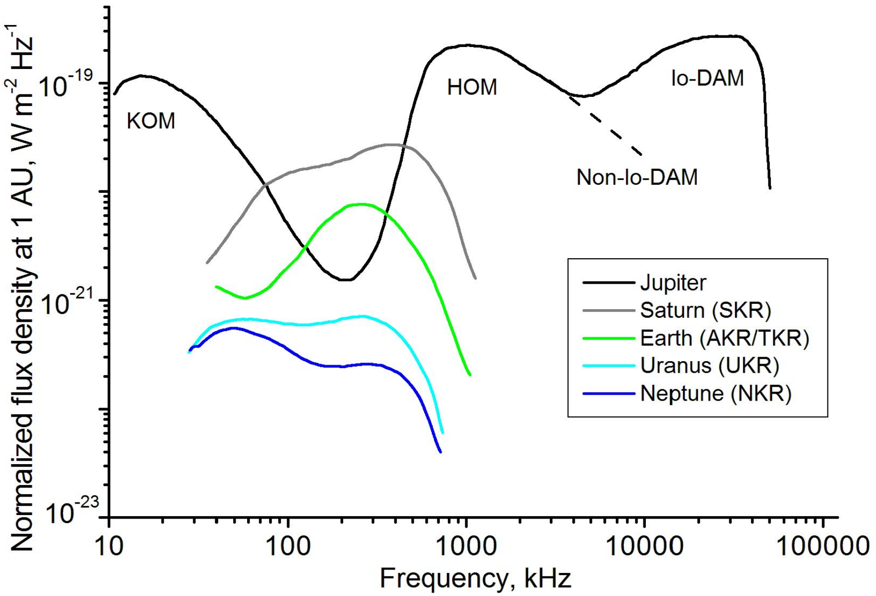
Goal 2: ***Perform the most sensitive observations of the radio sky at 1-50MHz with 20% absolute calibration;***

Goal 3: ***Quantify the systematic effects affecting global spectrum measurement accuracy, and investigate methods to mitigate them;***

Goal 4: ***Constrain the presence of a non-smooth monopole component at the 10^{-3} level compared to foregrounds.***

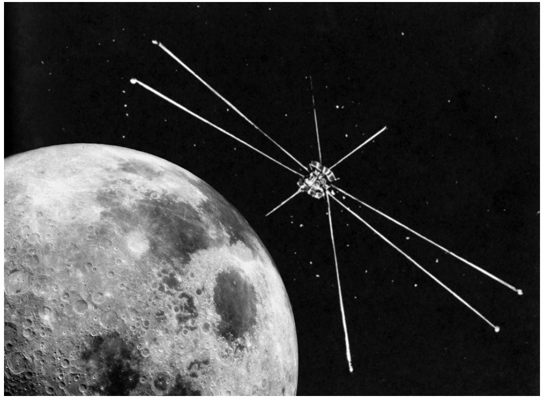
Secondary Science Goals

- Sun bursts during the day
- decametric radio emission from planets
- moon’s exosphere:
  - expected to have plasma frequency at around 30kHz during the day, falling to about 3kHz during the night
- Blind search across frequency

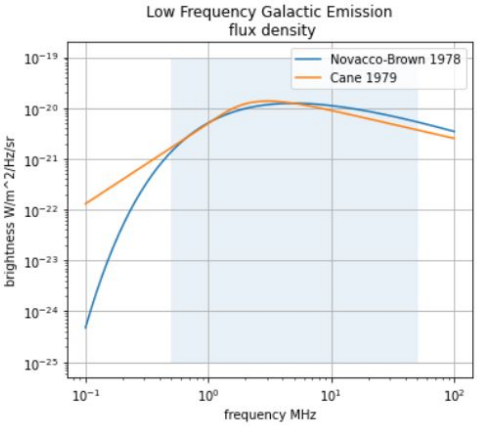


LuSEE-Night: A Dark Ages Pathfinder

- The main problem is that the Dark Ages signal is buried under order of magnitude brighter emission from foregrounds
- Currently the most informative data on low-frequency radio sky are from 1970s
- Our main goals are to update knowledge with modern instrumentation and pave the way for subsequent instrumentation
- We need to understand the limiting systematics on the Moon's surface:
  - how stable it really is?
  - how important are the effects of the regolith?
  - how important are transients, micrometeorites and similar?
  - etc



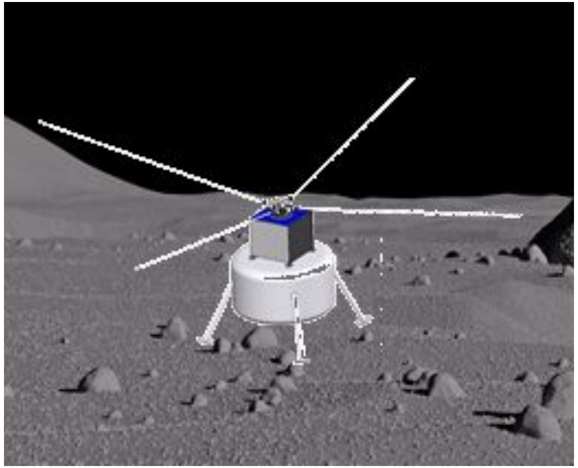
Radio-Astronomy-Explorer-2 was the largest spacecraft ever built with four 230m antennas



Data from 1970s:
still state of the art

LuSEE-Night instrument

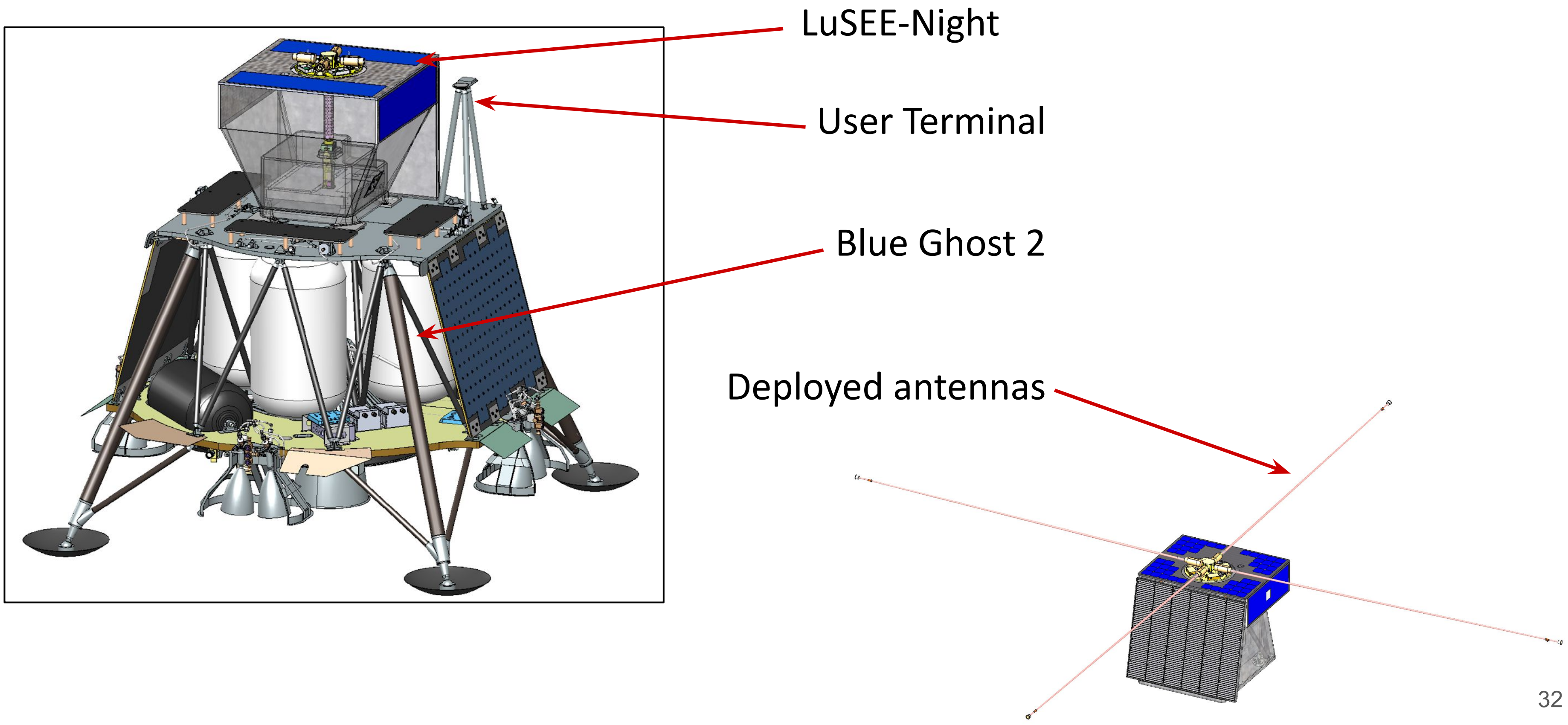
- Instrument design largely driven by the project timeline
- We need to optimize within the limits of:
  - max heritage technology
  - realistic improvements given budget
  - budget includes $, power, time-to-develop
- Basic design includes:
  - 4 STACER antennas
  - ability to actuate in azimuth
  - independently amplified
  - all possible cross-correlations measured and integrated with full bandwidth spectrometer
  - 4 real auto-correlations
  - 6 complex cross-correlation quantities, including two pseudo-dipoles



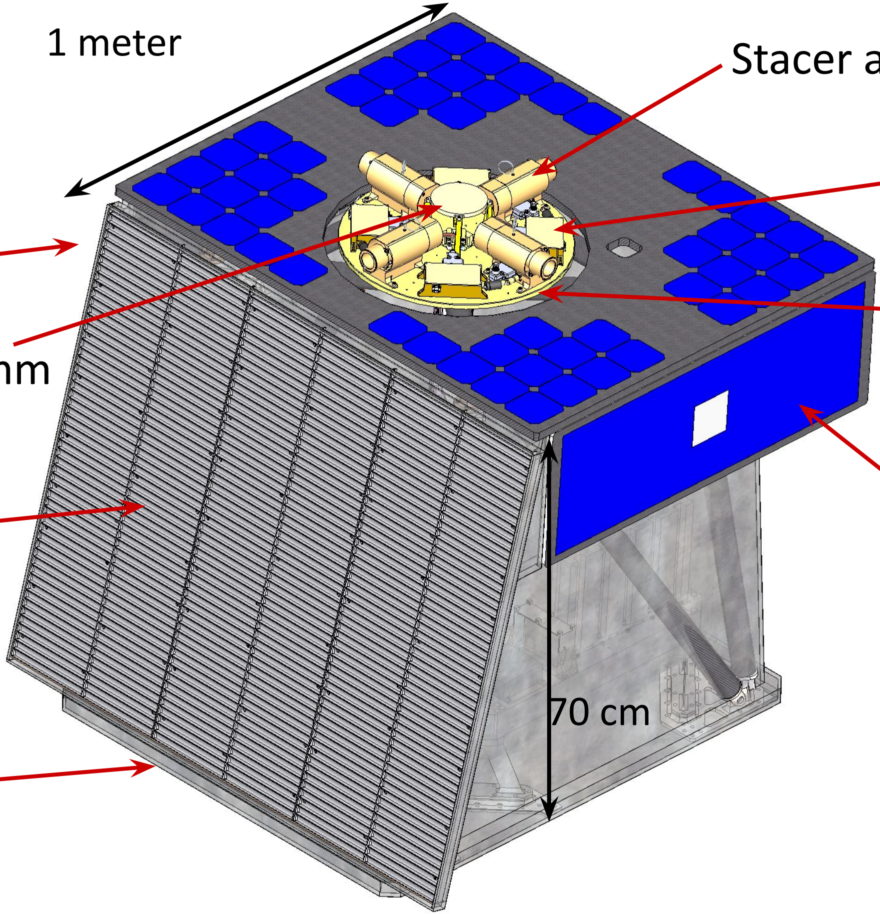
LuSEE-Night Mission.

- Firefly Aerospace has been selected as the CLPS mission provider
- We will launch of CLPS CS-3 mission == Firefly Blue Ghost 2 mission
- Firefly Blue Ghost 1 mission will be the CLPS flight #3 and launch in second part of C24
- Launch 2 Dec 25, landing 11 Jan 26
- We are the main payload, in addition:
    - “User Terminal” – a communication testbed
    - two private payloads

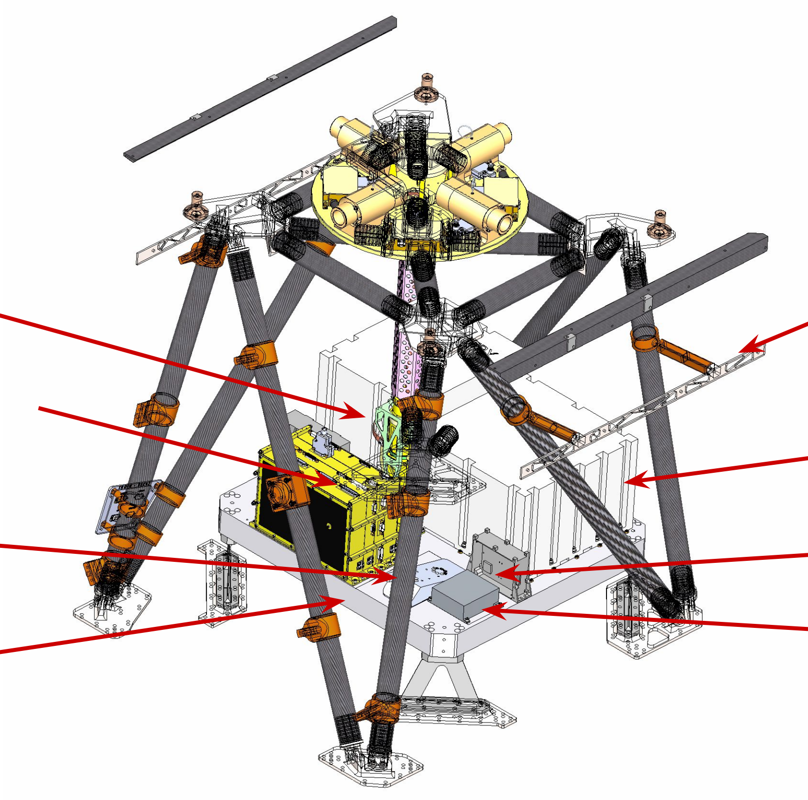
LuSEE-Night payload on Blue Ghost 2



LuSEE-Night payload



LuSEE-Night payload

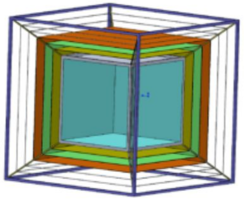


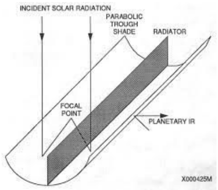
Challenges of a radio telescope on the Lunar far-side

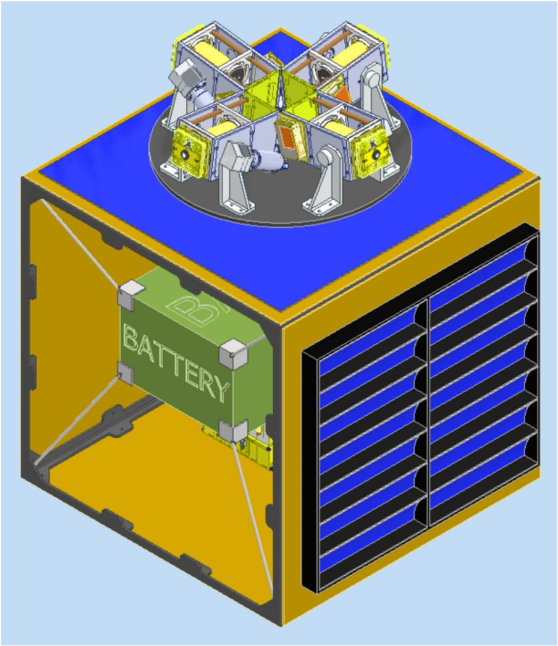
- Thermal management:
  - how to keep cool during the day
  - how to keep warm during the night
- Power management
- Data transfer
- Antenna performance
  - modulation
  - calibration

Thermal management

- A major consideration of the design
- Without atmosphere, everything is radiativelly coupled
- We copied Farside Seismic Suite concept
- During the night:
  - Sky at 3K, lunar surface at 100K
  - box-in-a-box concept with sufficient thermal insulation:
    - spacerless mylar insulation
  - ~20W battery dissipation keeps ut sufficiently warm passively
- During the day:
  - a variable conductance heatpipe
  - a parabolic radiator facing 3K







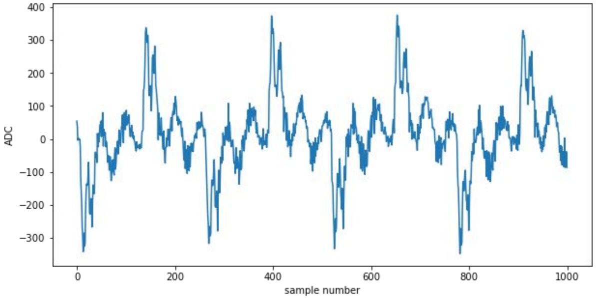
Power and Battery subsystem

Battery

- Our battery use is very atypical for space use:
  - We will recharge at most 20 times
- 50% of payload weight (~40kg) in battery
- 6500Wh with margin

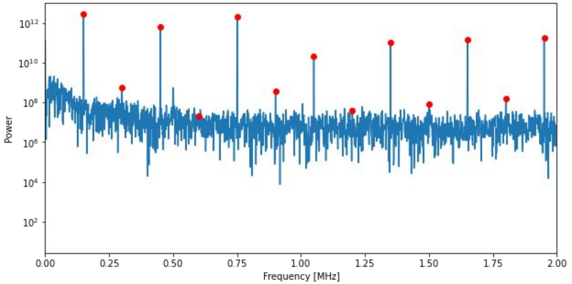
Picket-fence power supply

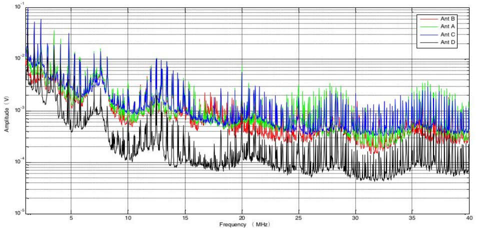
- picket-fence RFI management: all switching power supplies synched to the same master clock that also drive ADC
- All RFI confined to a base frequency + harmonics
- very strong requirement on lander complete death after first sunset
- Spikes can be removed very efficiently with DSP in spectrometer with negligible signal loss
- Self-RFI plagued the Chang’e-4



raw-signal from the parker solar probe

Cannot resolve width of these lines (<1Hz)





Chang’e-4 data dominated by platform RFI

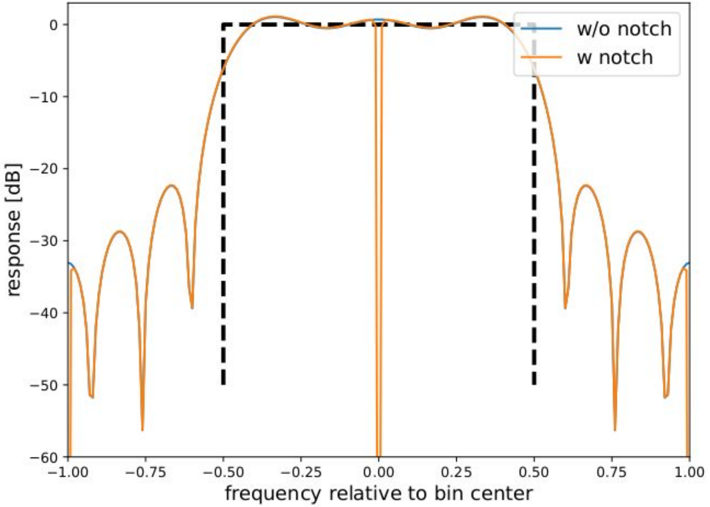
Spectrometer design

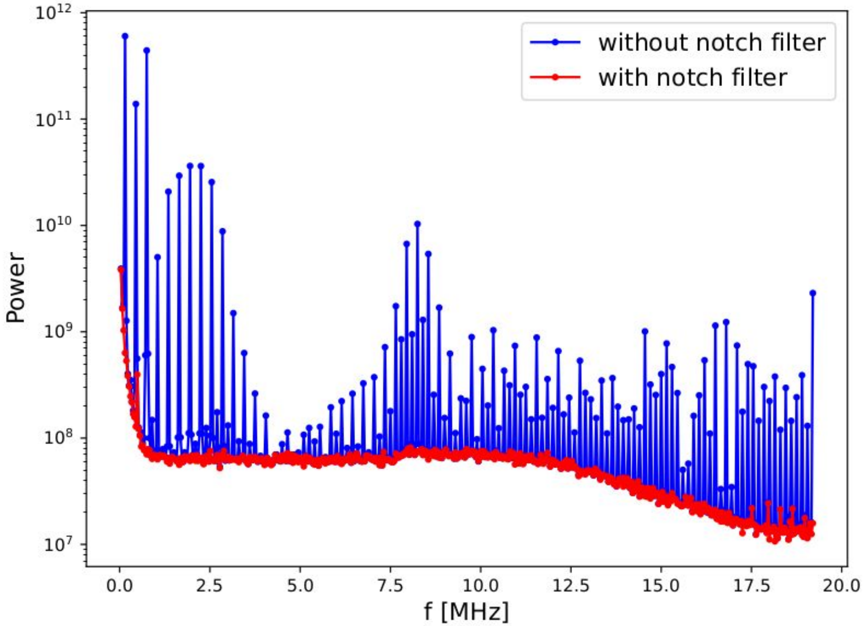
- Our design strongly based on the Parker Solar Probe Radio Frequency Spectrometer
- Updated FPGA, bandwidth, number of channels
- Nominally 2048 channel polyphase filter-bank
- But fancy:
  - notch filter
  - x5 spectral zoom region at low frequency
  - x64 single bin selectable zoom:
    - 400 Hz resolution, can resolve Farady rotation
  - calibrator support

Spectrometer parameter	Experiment	
	Parker RFS (HF channel)	LuSEE-Night
Frequency range	0.01 – 19.2 MHz	0.5 – 50.0 MHz
Digitizer channels	2	4
Sampling clock	39 MSPS	100 MSPS
Duty factor	~1%	~100%
FPGA	RTAX4000	RTG4
Correlation products	2	6
Transient capture	No	Yes

Current devel version runs full 51.2MHz bandwidth, 4-input, 2048 channel spectrometer on <6W

Notch filter to remove picket-fence





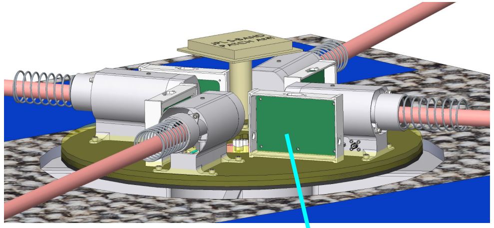
Antennas: science requirements

1. Smooth spectral response
2. Sensitivity
3. Polarization-sensitive
    - To separate polarized foregrounds from unpolarized Dark Ages
4. Manage effect of lunar regolith
5. Mitigate systematics

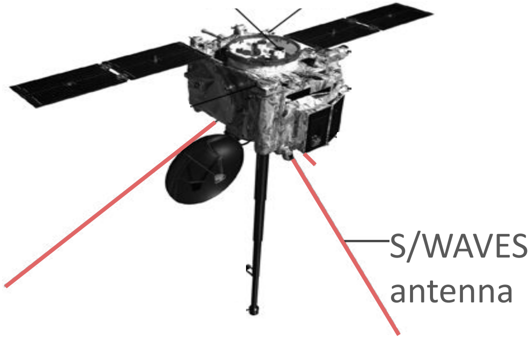
Mission requirements:

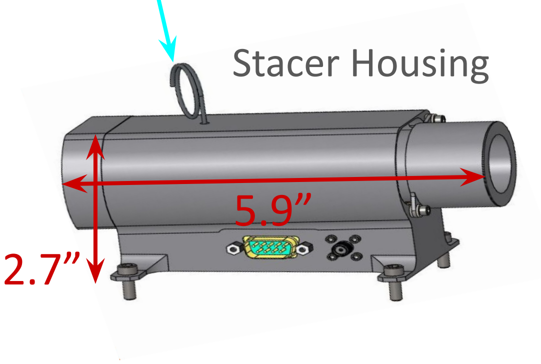
- High TRL
- Light, compact, robust
    - Antenna subsystem < 7.8kg

Four 3-meter monopole antennas



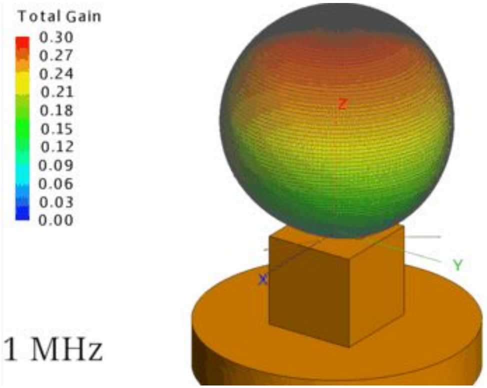
STEREO/WAVES



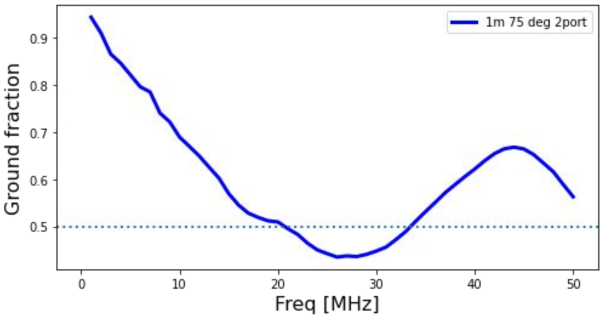


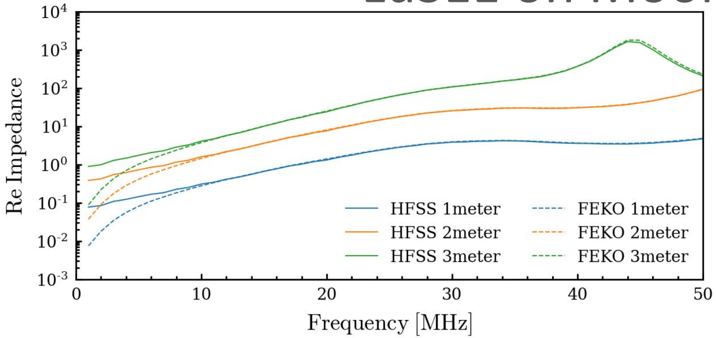
Modeling LuSEE on the lunar far-side

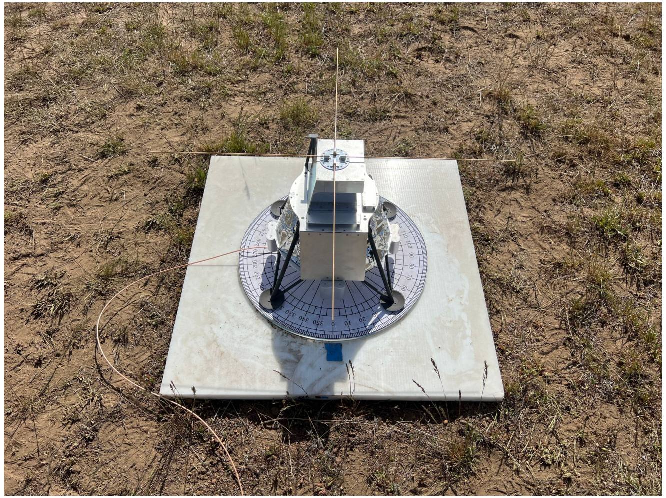
- Optimizing antenna design within mechanical constraints is challenging
- Chromaticity unavoidable
- Beams affected by lander and regolith.
- Fraction of power absorbed by (dielectric) regolith significant
- Have done significant work comparing results with three simulation packages (CST, HFSS, FEKO)

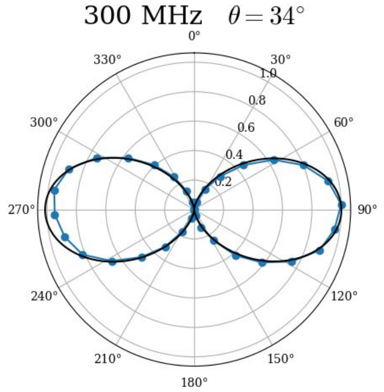


LuSEE on Moon



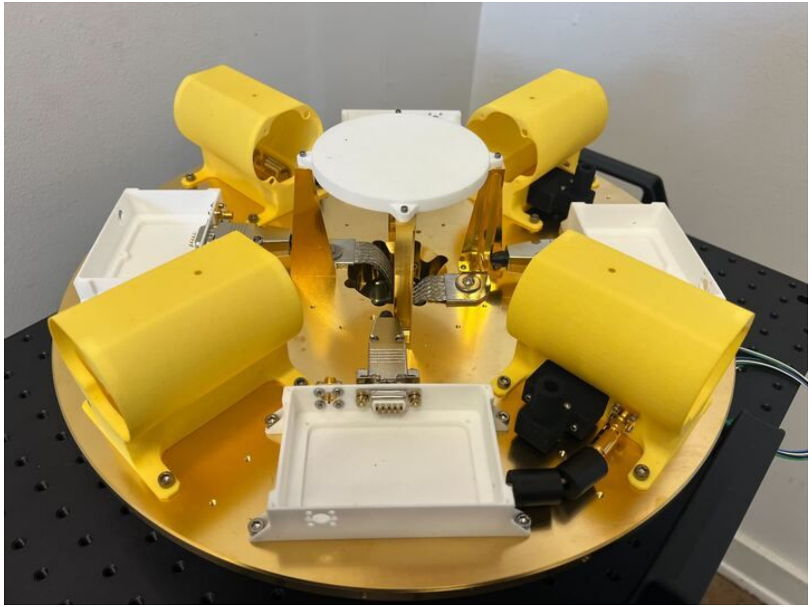


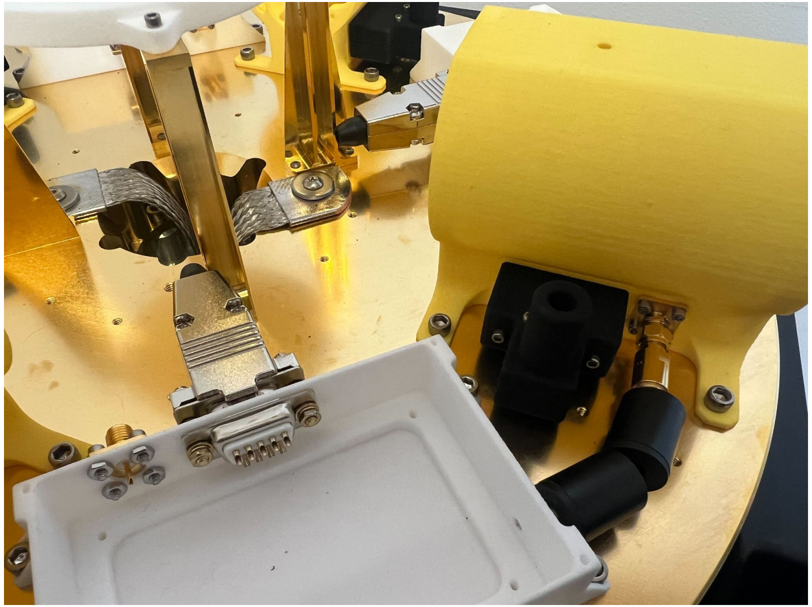




Antenna testing with a scale model. Kaja Rotermund, LBL

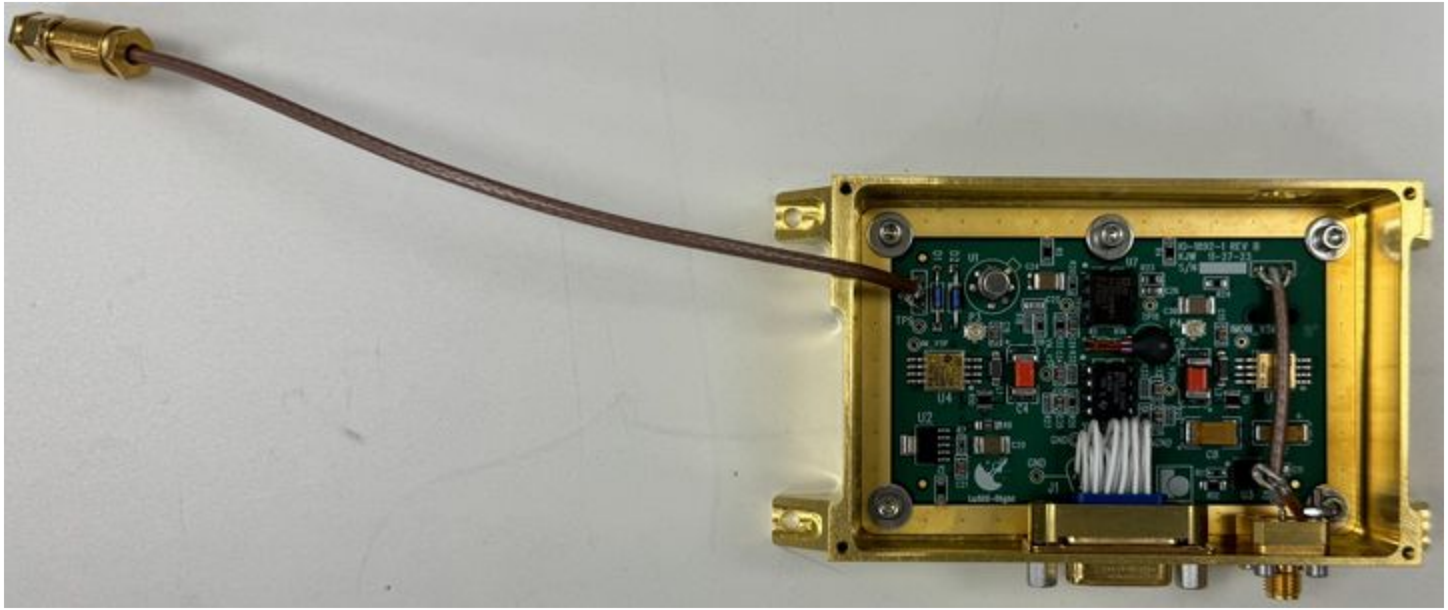
It is becoming real





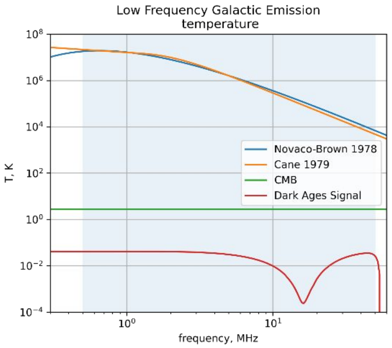
Test-fitting flight-model turntable
with 3D printed components

It is becoming real



Flight-model pre-amplifier

Signal to noise considerations



- Signal to noise is governed by the radiometer equation:

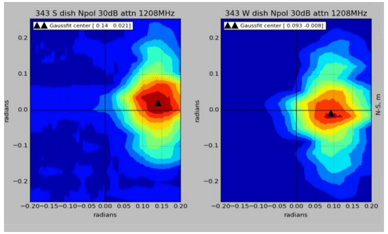
$$\Delta T = \frac{T_{\text{sky}} + T_{\text{noise}}}{\sqrt{Bt}}$$

- LuSEE-Night is a 4 element interferometer, so O(1) factors better
- Square root factor is a experimenter's greatest foe
- The realized LuSEE-Night noise is below initial requirements
- Table below assumes a perfect sky model is given (which it is not), so for illustrative purposes only

T_{noise} / T_{sky}	**0 (~realized)**	**0.5 (~design)**	1	2	10
Time to integrate to CMB **(mins)**	**8**	**20**	30	70	1000
Time to integrate to Dark Ages **(lunar cycles)**	**2**	**4**	7	17	225

Far-field Calibrator

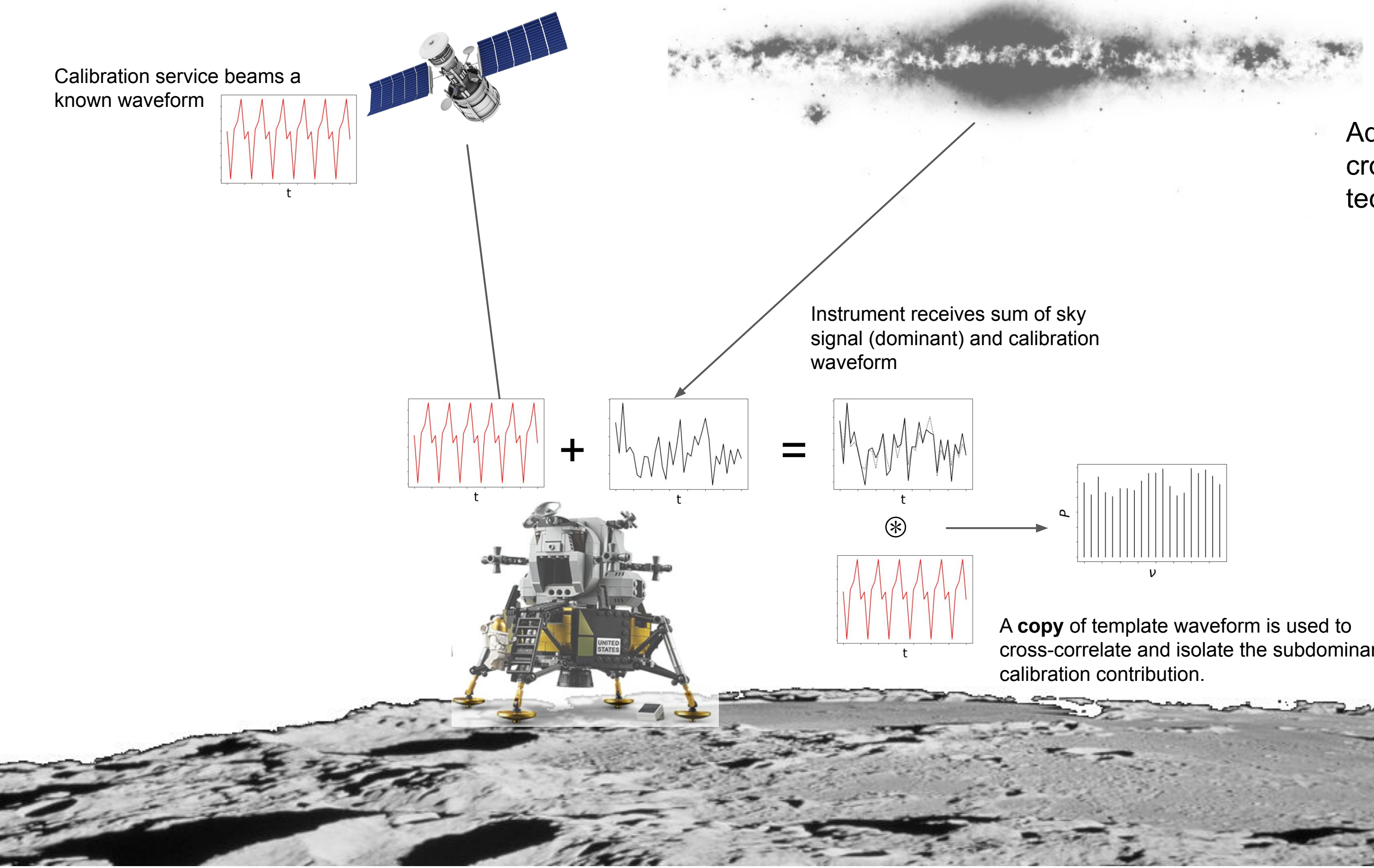
- Field of 21cm is a study of calibration methods
  - Foregrounds many times brighter
- Precision calibration of radio telescope beams at sub-percent precision remains to be demonstrated
- Many attempts with drones, but nobody has reached required precision
- **A true far-field, power stabilized calibrator is almost certainly necessary for dark ages work:**



Basic Concept

Calibration service beams a known waveform

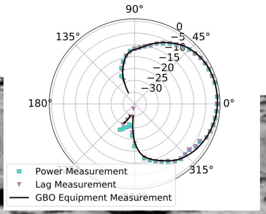
Instrument receives sum of sky signal (dominant) and calibration waveform



A **copy** of template waveform is used to cross-correlate and isolate the subdominant calibration contribution.

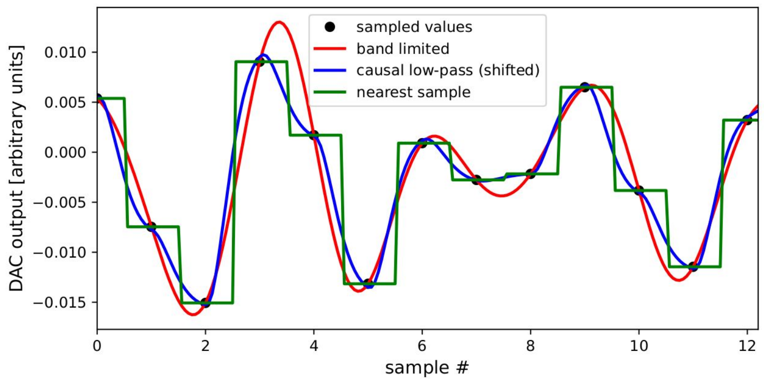
Advantages of cross-correlation technique:

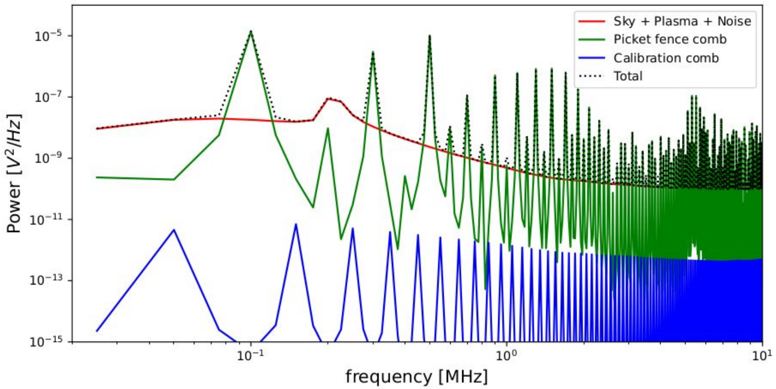
- sky background does not cross-correlate
- only one copy of template attenuated by the distance (1/r scaling), allowing high SNR recovery
- See arXiv:2201.11806 for similar approach on the ground



Calibration scheme

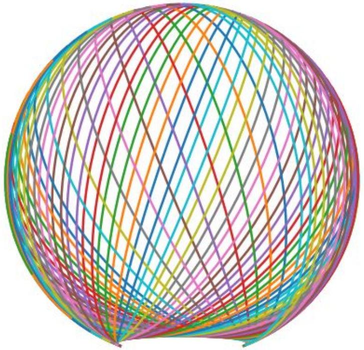
- A pseudo-random waveform of length N is generated so that its Fourier components have random phase and unit argument
- If this waveform is repeated, it generates a frequency comb
- We are detecting this comb, designed to be slotted between our picket-fence comb
- Main issues:
  - relative clock drifts
  - DAC/ADC frequency aliasing



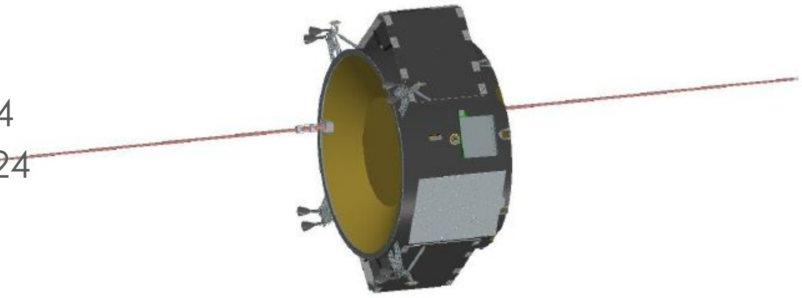


Calibrator status

- CLPS CS-4 mission is for *provision of calibration service*
- Firefly has been selected as a provider of CS-4 calibration service
- Firelfy has subcontracted Vulcan Wireless to develop the payload
- The work is progressing well:
  - PDR for CS-4 occurred end of January 2024
  - FDE for CS-4 scheduled for end of June 2024



Sample satellite coverage



HFSS modelling of calibrator

Wiener deconvolution map reconstruction

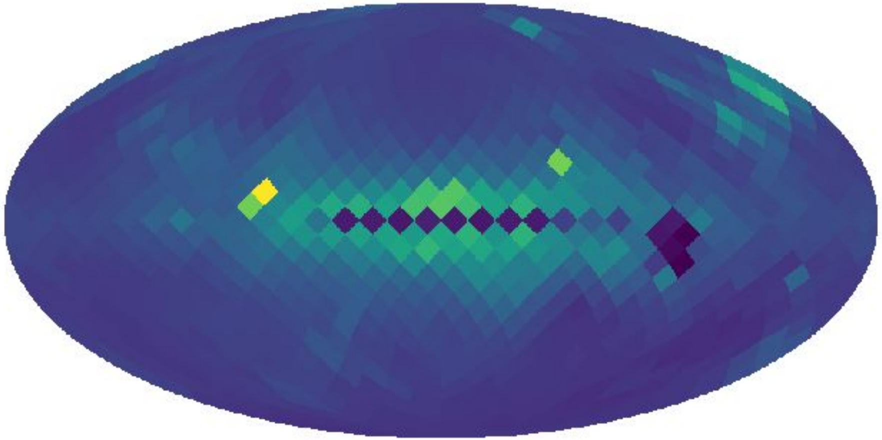
- We measure 16 data products
- Each of these data is a different beam integrated over the visible sky
- However, after taking into account rotation of the Moon and the carousel, we have measurements of a number of different linear combinations of the sky

$$d = \mathrm{A}m + n$$

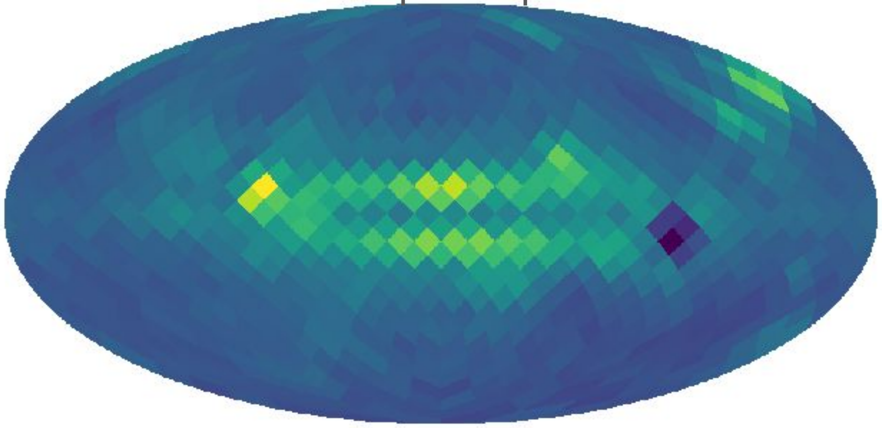
- Wiene deconvolution allows us to reconstruct the low-resolution map of

  the sky

- In collaboration with Hugo Camacho (BNL)

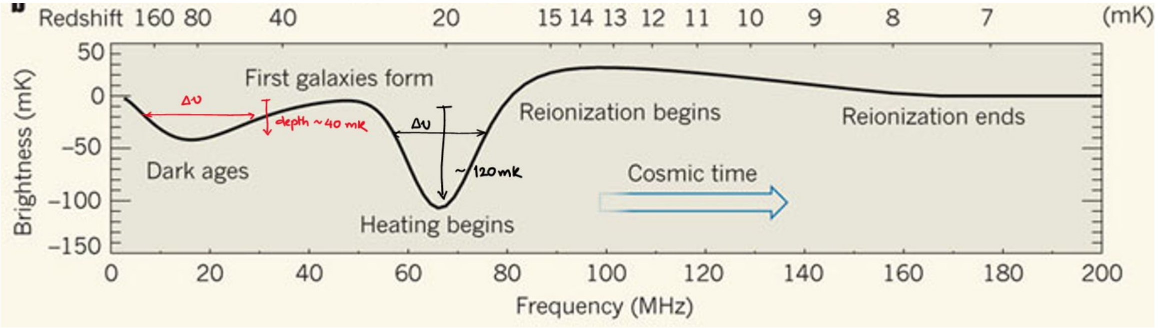


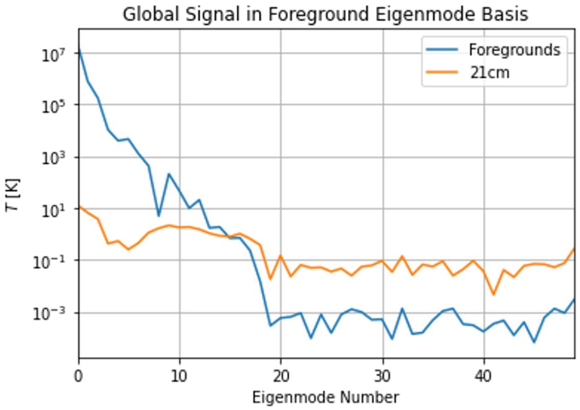
input map



reconstructed map

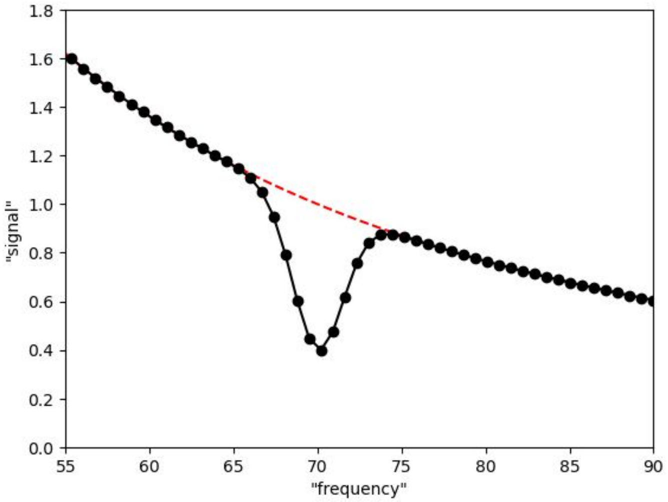
Extracting cosmologically interesting signal





- Traditionally, we rely on spectral smoothness
- Does not work in Dark Ages signal:
  - feature is just to broad
  - Δυ ~ υ
- Instead, assume that fluctuations around the mean foreground shape are drawn from the same distribution as the foreground mean

Smoothness separation



- We want to identify a presence of a sharp feature on the smooth background
- The process:
  - fit a smooth "foreground" away from the expected peak position
  - subtract the smooth component
  - what remains is is the signal of interest
- For this to work, **the uncertainty in background determination should be subdominant:**
  - i.e. it shouldn't matter what functional form I use for fitting the background
- For smooth, e.g. power-law like foregrounds

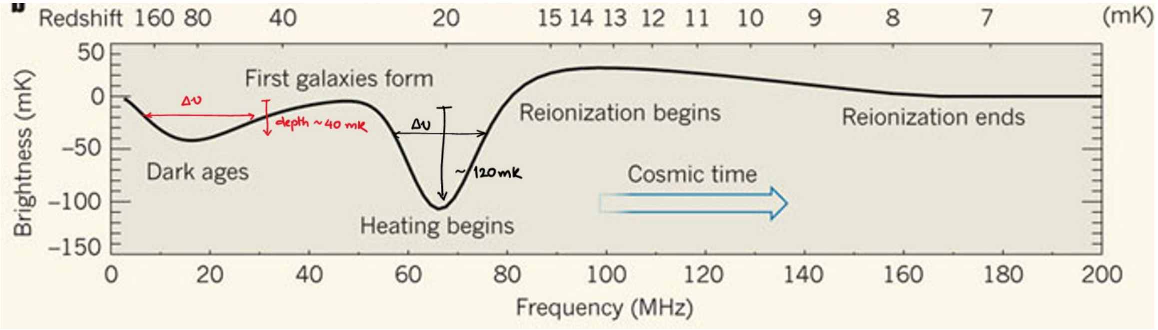
$$\frac{\Delta F}{F} \sim \frac{\Delta \nu}{\nu}$$

- And foreground fit uncertainty over the feature width will scale the same
- Therefore, the "smoothness separation" paradigm only works when

$$\frac{S}{F} \gtrsim \frac{\Delta \nu}{\nu}$$

- Feature should be either very localized, or fractionally large
- Otherwise, the answer will depend on the foreground functional form

Will the smoothness separation work for the monopole?

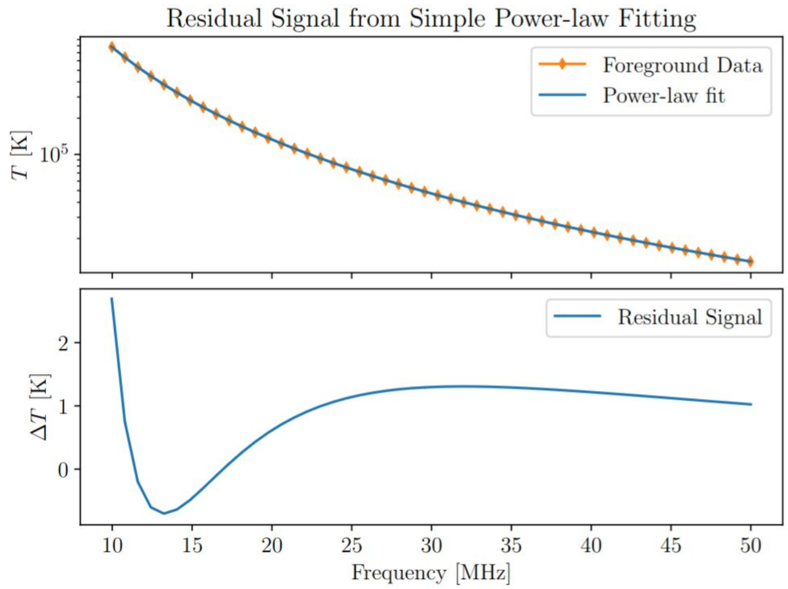


Signal	S/F	$\Delta\nu/\nu$	Does it work?
BAO	~0.3	~0.2	Yes, barely (but we have no-wiggle theory)
Cosmic Dawn	~10^{-3}	~0.1	No.
Dark Ages	~10^{-5}	~1	Nooo!!!

Ergo the incessant EDGES fights

To drive the point home...

- Take a sum of two power-law with indices -2.54, -2.56
- Fit with a single power law (get spectral index -2.55)
- Look at the residuals
  - residual feature is no broader than the Dark Ages signal
  - Yes, you can absorb it, but then so you will absorb the Dark Ages signal...



What is the way out?

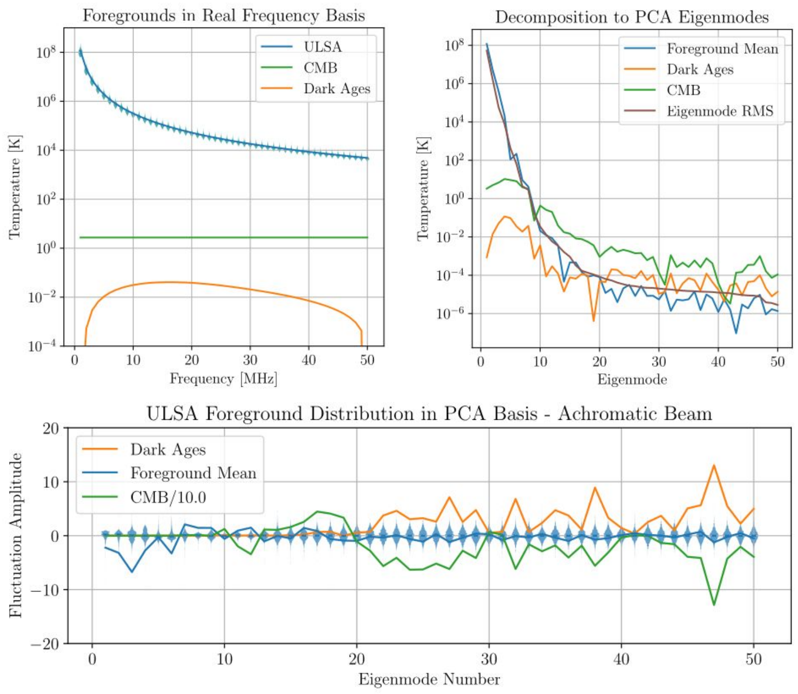
- You need to add more information:
  - more actual data – spatially resolved spectra
  - priors (theoretical or otherwise)
- Approach 1:
  - write down a complete Bayesian model for your data, your instrument and any additional information
  - perform an optimal Bayesian analysis
  - you need to describe data at 10^{-5} level:
    - will likely need many thousands, potentially a million parameters
    - need to go beyond’s
  - see Rapetti’s talk
- Approach 2:
  - use foreground fluctuations to build a data-driven model of foregrounds
  - very different set of assumptions

Pund and Slosar, arXiv:2310.06134

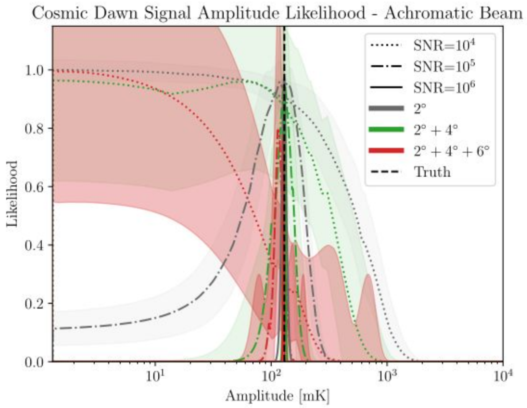
- Replace “signal is smoother than foregrounds” with “foreground fluctuations around the mean are generated by the same process as the foreground mean”
- Defensible assumption, but it is a physical one
  - if foregrounds are generated by some local process, the beam will modulate them
- Chromaticity still hurts:
  - spreads foregrounds over larger number of “shape” space
- Transfer function:
  - hurts less: need to know beam to percent rather than ppm
- Gain fluctuations:
  - hurt more: they make signal appear as a foreground
  - percent level control is sufficient

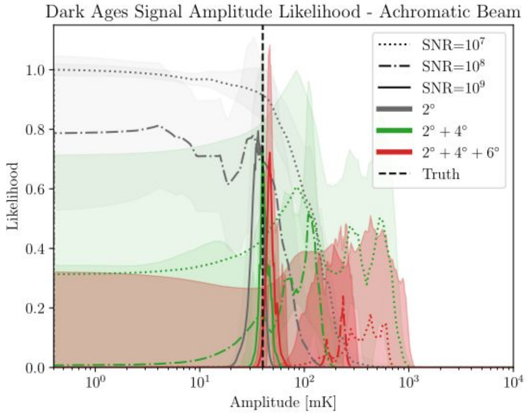
New approach

- Simulation using a Cong et al ULSA maps
- 2 degree resolution instrument with galactic cut
- Fit 3 parameter model; dark ages signal scaled by:
  - amplitude
  - central frequency
  - width
- First 20 SVD modes lost to foregrounds, but
  some information remains

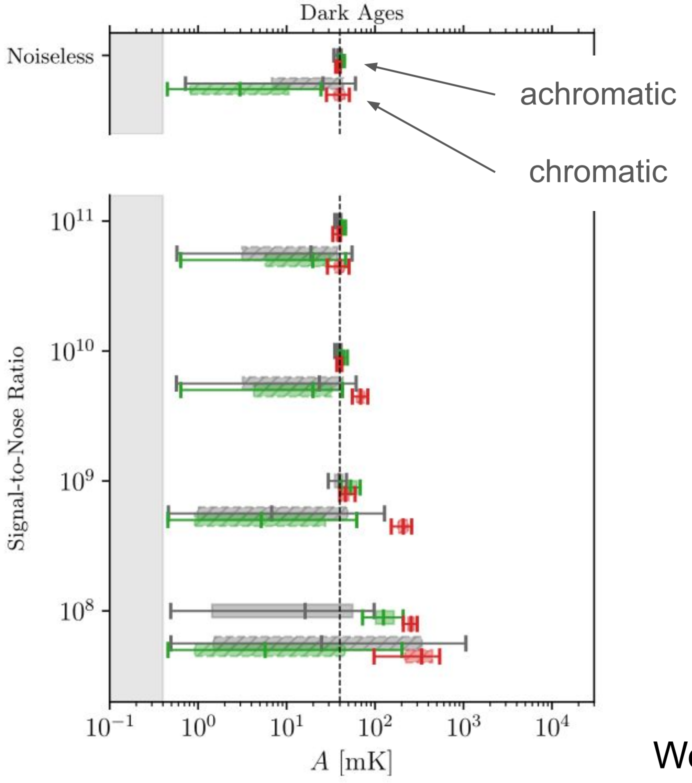


Some results

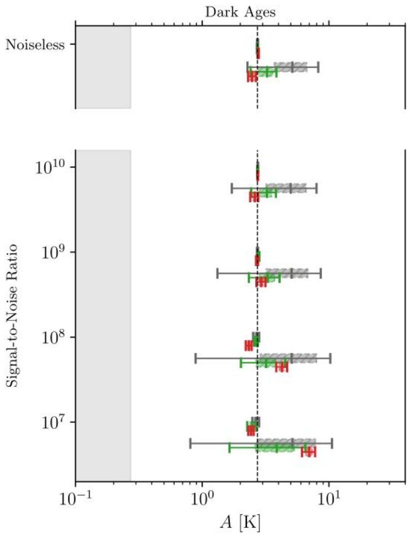




Existence proof!



Chromaticity still hurts!

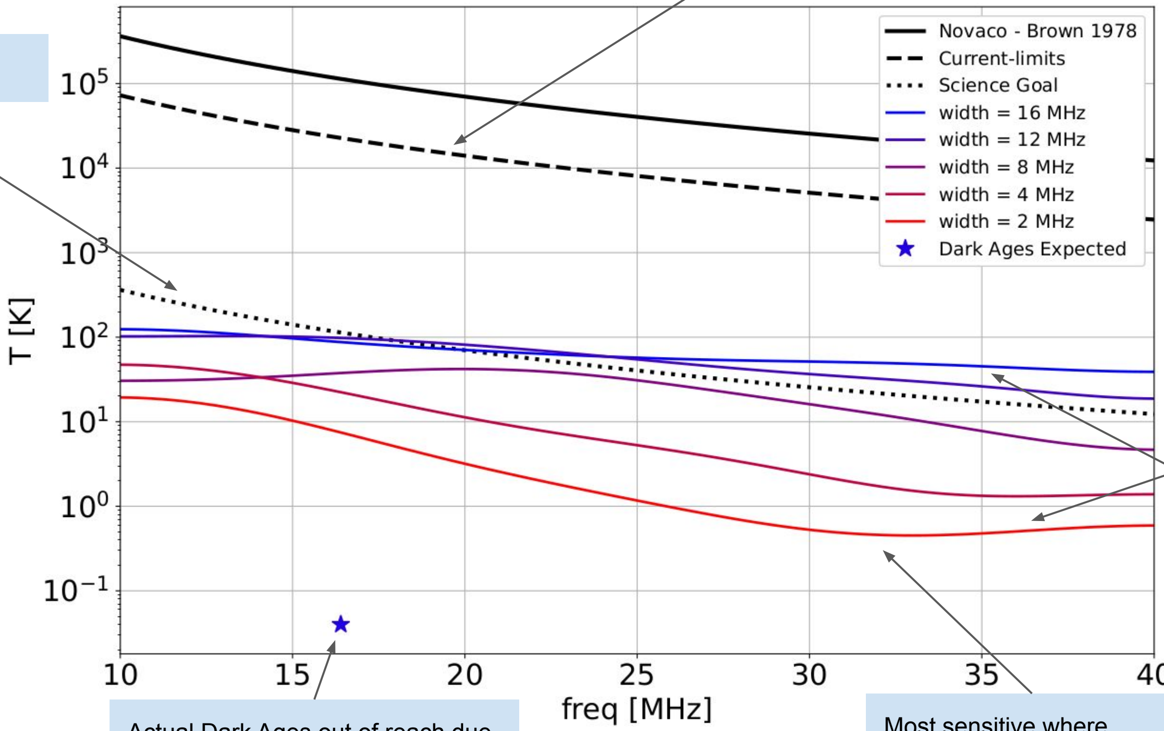


We can perhaps detect CMB!

Sky Monopole Sensitivity plot

Current limits: 20% variations around Novaco - Brown

10^{-3} science goal #4



Narrow features are easier to detect since foregrounds are smooth

Actual Dark Ages out of reach due to chromaticity of dipole antennas

Most sensitive where radiometer is most sensitive

Conclusions

- If many things go to plan
  - LuSEE-Night will establish the viability of radio observations from the Lunar surface
  - LuSEE-Night will produce the most accurate sky map at 1-50MHz
- Project fully funded and manifested on CLPS CS-3 mission with calibration service as CS-4 mission
- Firefly Aerospace chosen as the provider on both
- For the first time in my life I hope for project delay…