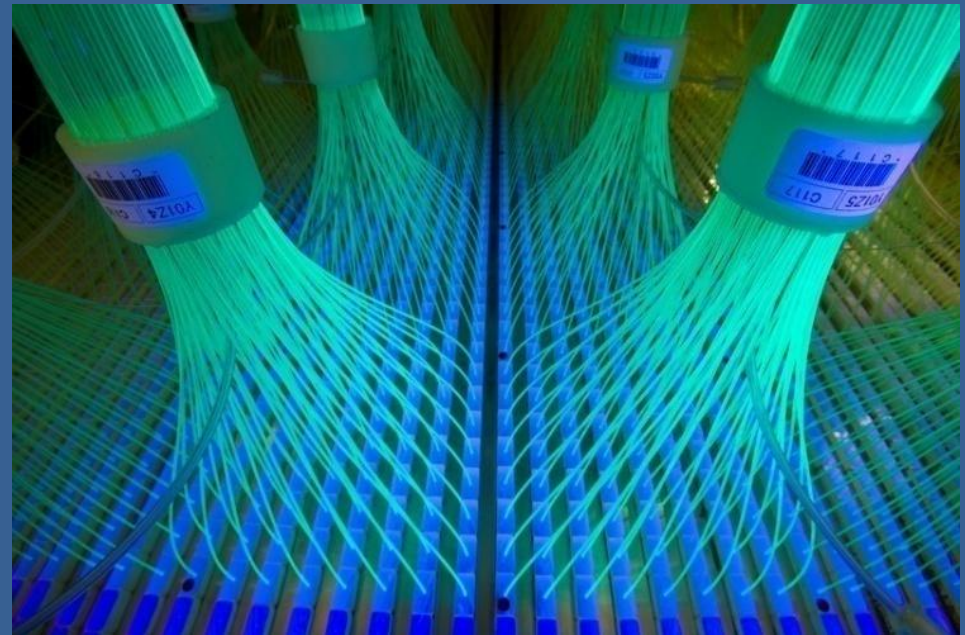




# An Update on the Progress of the SciBooNE Experiment



Robert Napora  
Purdue University Calumet  
May 30<sup>th</sup> 2007

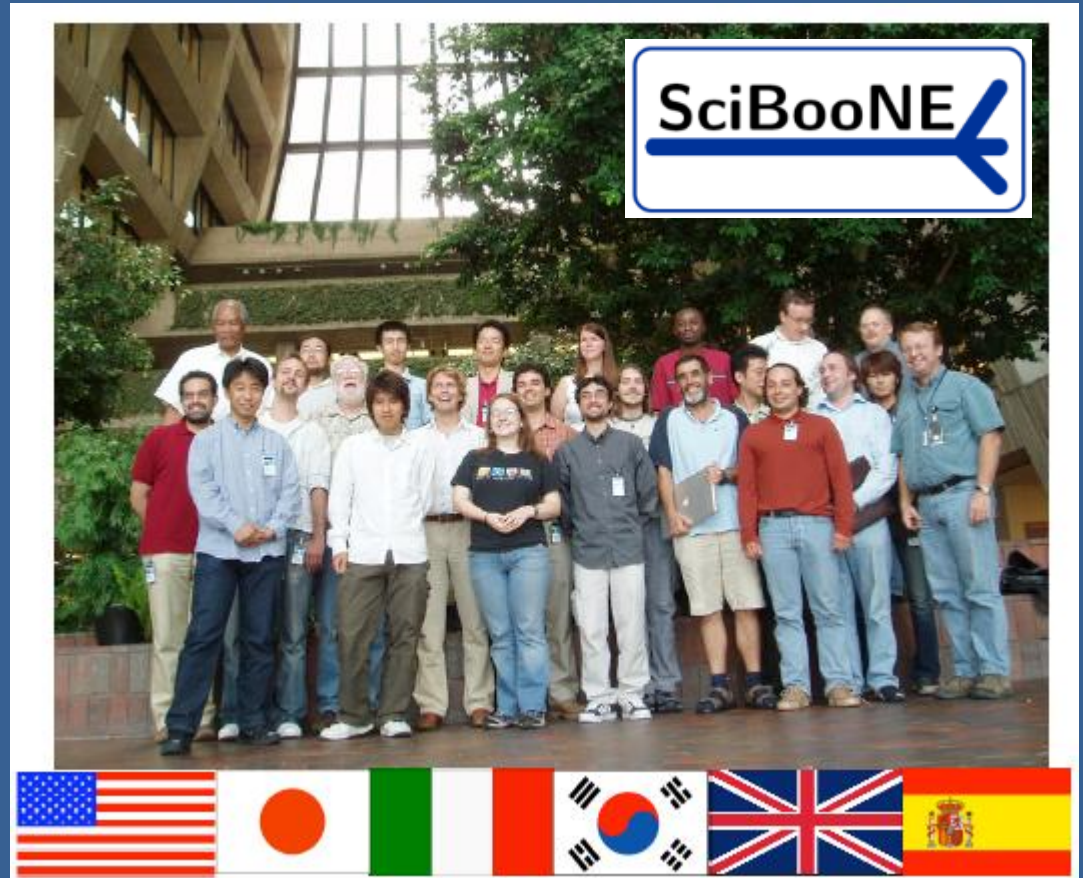


# Outline

- Introduction to SciBooNE
- SciBooNE Detector
- SciBooNE Physics
- SciBooNE Status/Timeline
- Summary

# The SciBooNE Collaboration

- Universitat Autònoma de Barcelona
- Chonnam National University
- University of Cincinnati
- University of Colorado
- Columbia University
- Dongshin University
- Fermi National Accelerator Laboratory
- High Energy Accelerator Research Organization (KEK)
- Imperial College London\*
- Indiana University
- Institute for Cosmic Ray Research
- Kyoto University\*
- Los Alamos National Laboratory
- Louisiana State University
- Purdue University Calumet
- Università degli Studi di Roma and INFN-Roma
- Saint Mary's University of Minnesota
- Seoul National University
- Tokyo Institute of Technology
- Universidad de Valencia

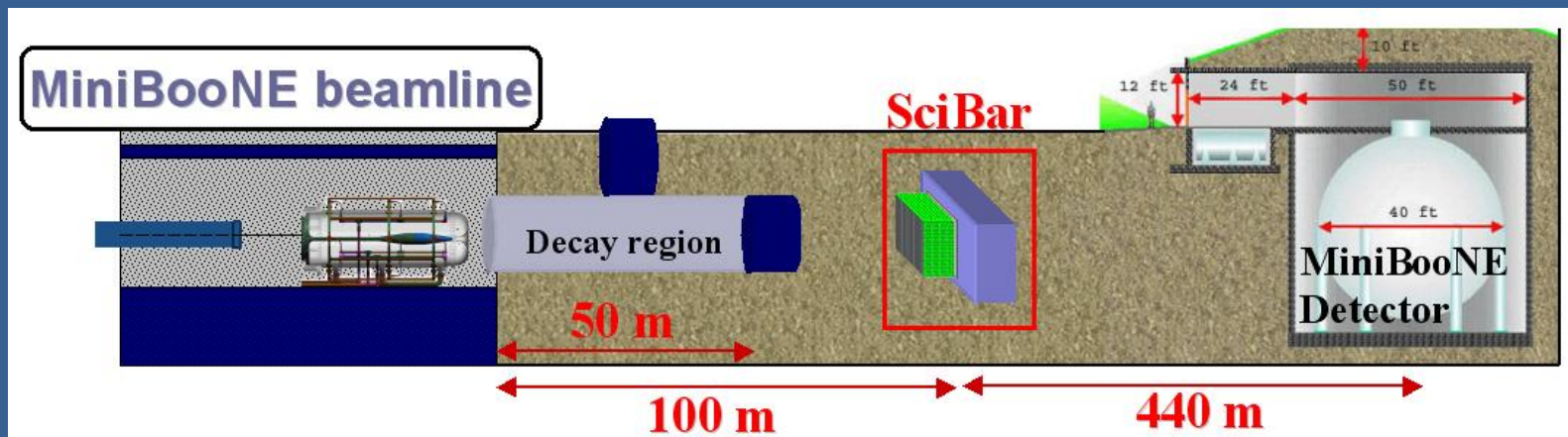


About 70 collaborators from 19 institutions and 6 countries.



# SciBooNE

Placed the preexisting SciBar/EC detectors along with a locally constructed MRD in the Booster Neutrino Beam.



using (mostly)  
pre-existing  
detectors

+

piggy-backing  
on an existing  
beam

=

good physics  
at low cost

# The Booster Neutrino Beam

➤ 8 GeV protons incident on beryllium target. Magnetic horn focuses appropriately signed  $\pi/K$

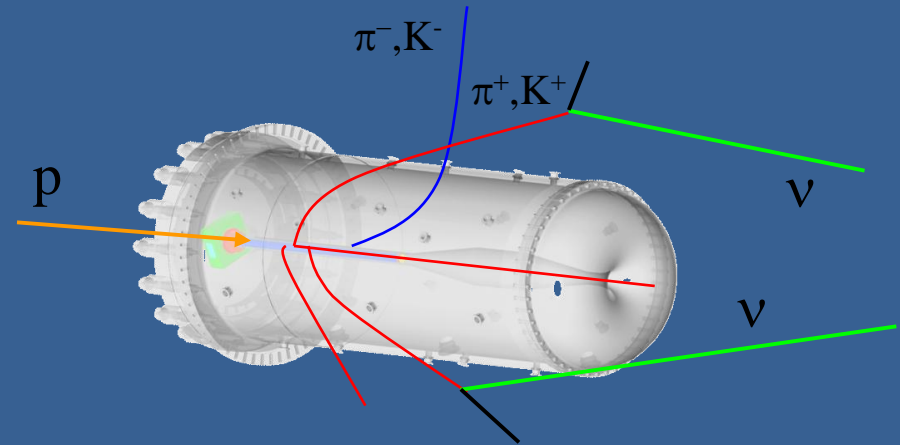
➤ BNL E-910 and CERN HARP data for secondary production predictions

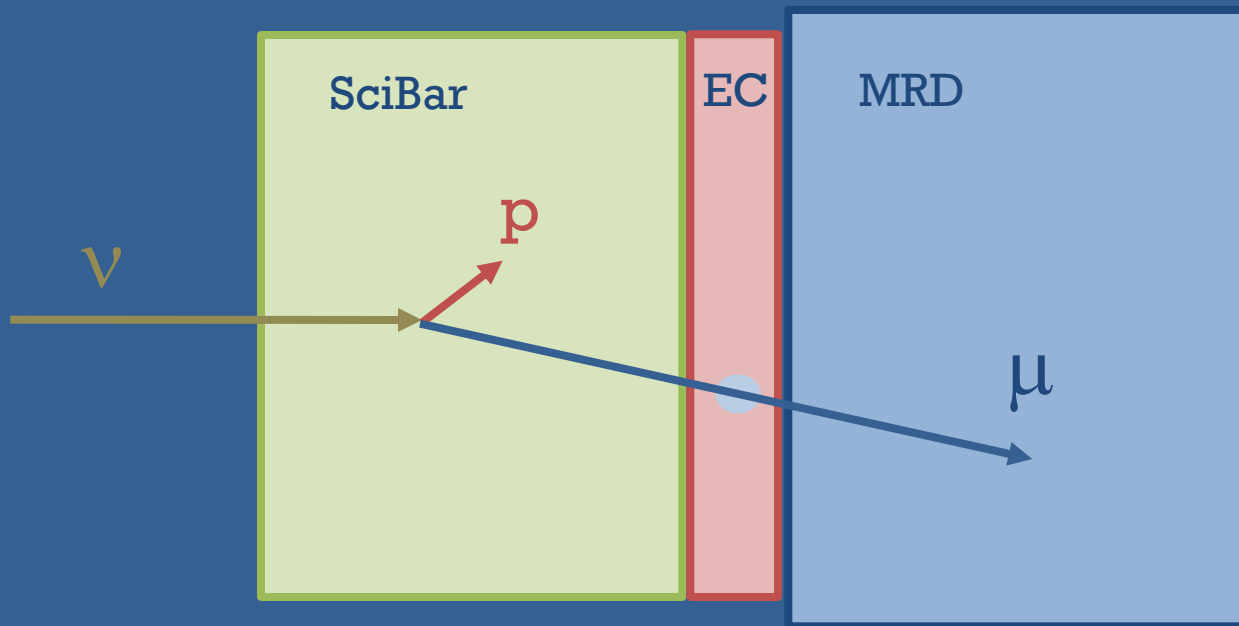
➤ Mean  $\nu$  energy 0.8 GeV  
Peak  $\nu$  energy 0.65 GeV  
→ similar to T2K spectrum

➤  $\nu_e$  contamination  $< 0.6\%$

➤ SciBooNE anticipates

- $\sim 1e20$  POT  $\nu$
- $\sim 1e20$  POT anti- $\nu$





Schematic of  
CCQE event:

$$\nu_{\mu} + n \rightarrow p + \mu$$

$$(\nu_e + n \rightarrow p + e)$$

### Scintillating Bar Detector (SciBar)

- Extruded scintillator with WLS fiber readout
- 224 MA-PMTs x 64 ch/pmt = 14,336 channels
- Scintillator is neutrino target (~15 tons)
- $2.5 \times 1.3 \times 300 \text{ cm}^3$  cells
- Can use  $dE/dx$  to distinguish  $\pi$  from  $p$
- Previously used in K2K experiment

### Muon Range Detector (MRD)

- Steel and scintillator sandwich calorimeter
- Twelve  $9' \times 10' \times 2''$  steel sheets, 48 tons of absorber material
- Can stop virtually all muons of momentum  $\sim 0.9 \text{ GeV}/c$  or lower (within acceptance)
- 362 channels
- Constructed here at Fermilab, mostly out of recycled parts

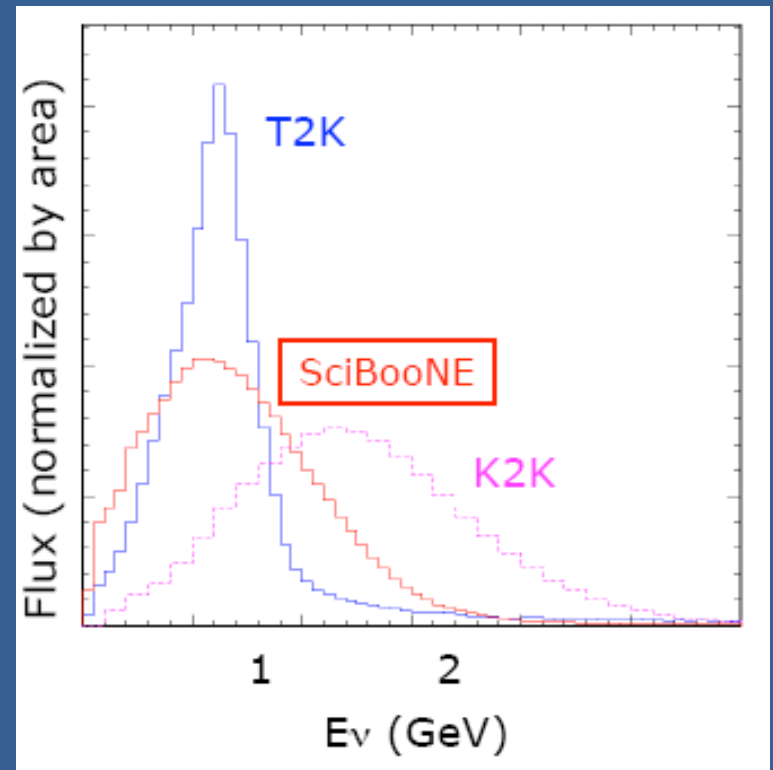
### Electron Catcher (EC)

- “Spaghetti” calorimeter
- 1mm fibers in grooves between lead foils
- 1 vertical and 1 horizontal layer
- $4 \times 4 \text{ cm}^3$  readout from both ends
- 256 channels
- Previously used in K2K, CHORUS

# SciBooNE Physics

SciBooNE will measure neutrino and anti-neutrino cross-sections in the Booster Neutrino Beam.

- Precise knowledge of cross sections necessary to T2K and other experiments
- MiniBooNE near detector; cross-check for BNB
- Anti-neutrino cross section measurements would complement MiniBooNE values. Useful for T2K phase-II.



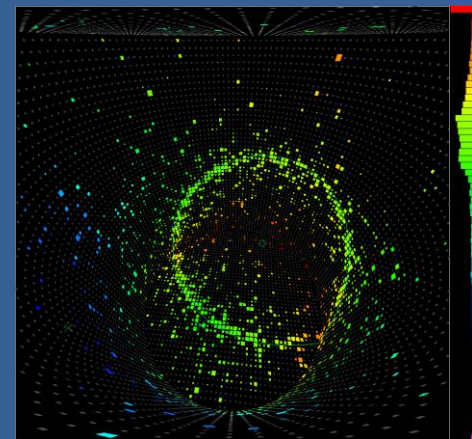
# SciBooNE's fine grained resolution affords certain advantages. For example...

---

You are searching for  $\nu_\mu \rightarrow \nu_e$  oscillations (measuring  $\theta_{13}$ )

Signal: CC-QE interaction  $\nu_e + n \rightarrow e^- + p^+$

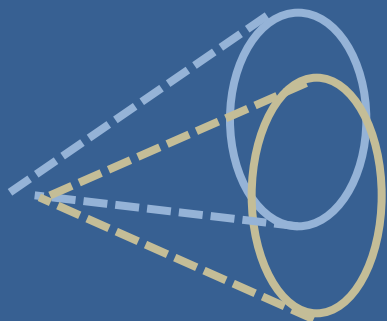
Signature: Cherenkov ring made by the electron



The dominant background is the NC- $\pi^0$  interaction



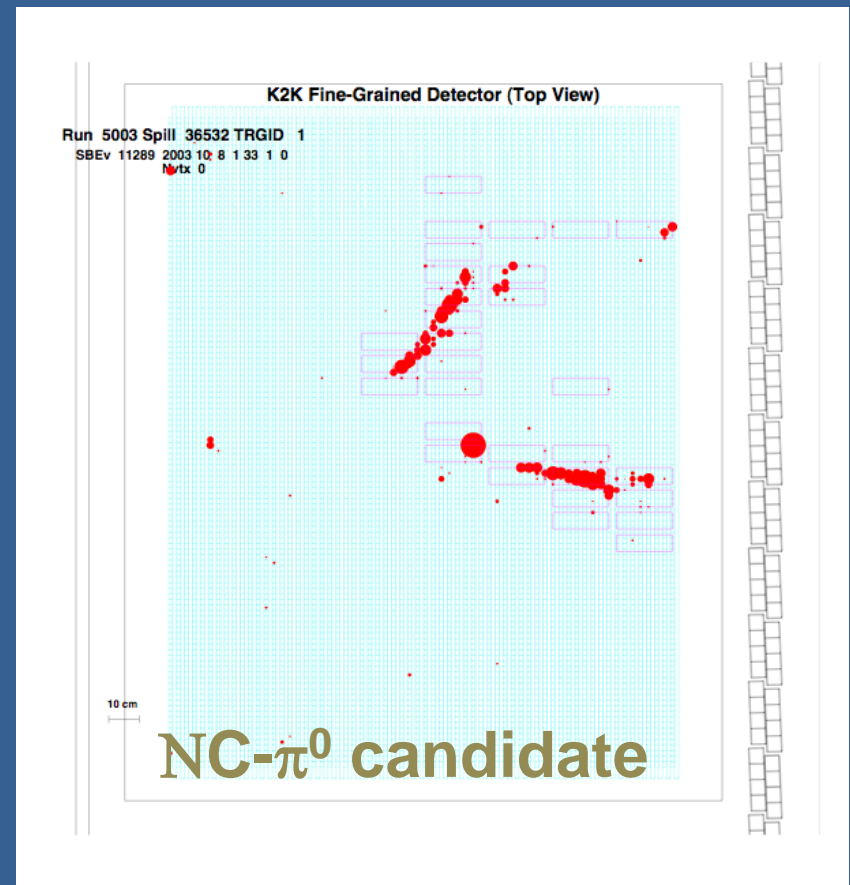
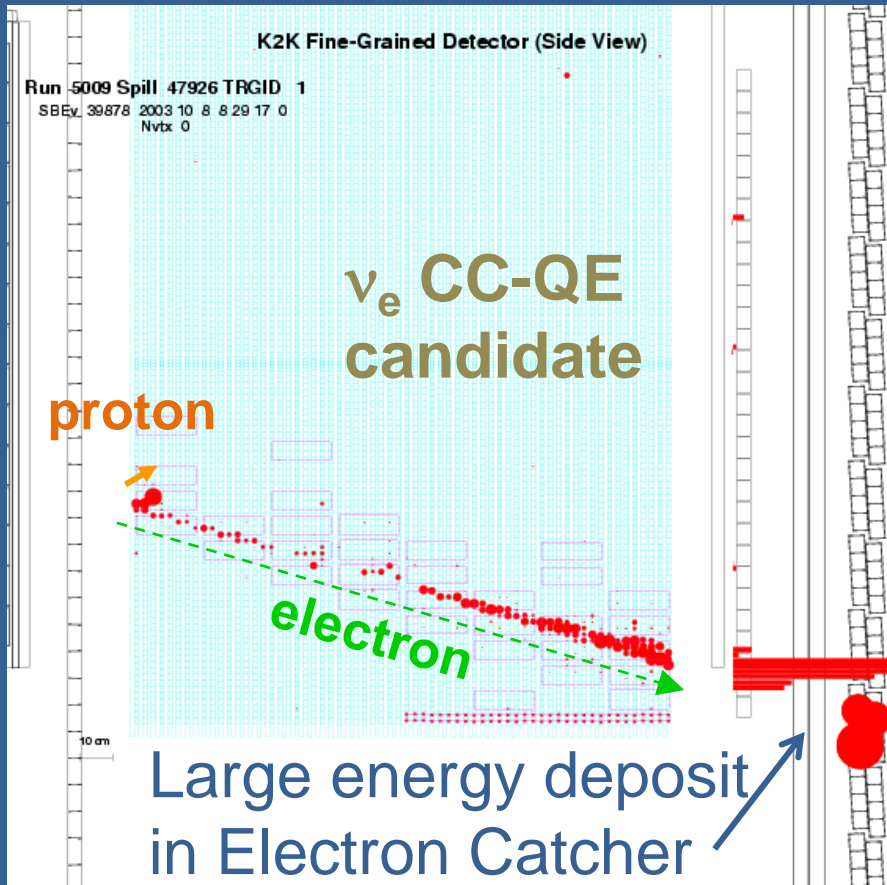
where the pion decays to two photons



The two photons can make a “fuzzy ring” that looks like an electron to your Cherenkov detector.



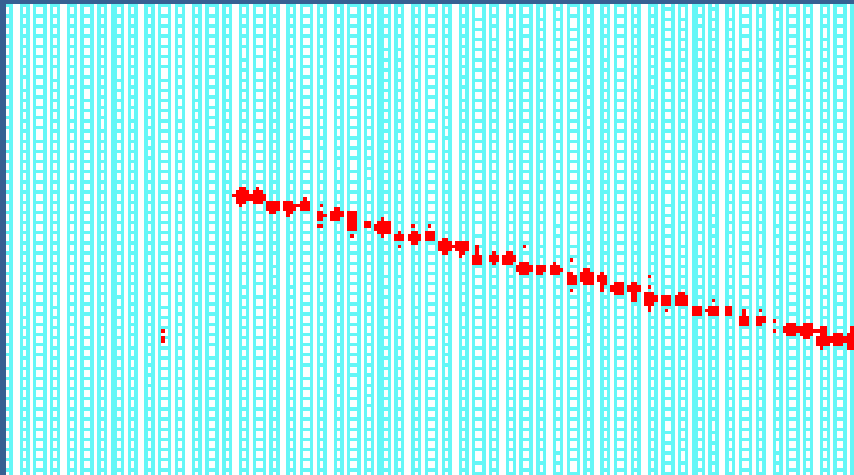
SciBooNE will be able to discriminate between these two channels.



These are event displays of SciBar data from K2K.

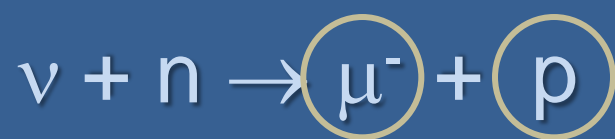
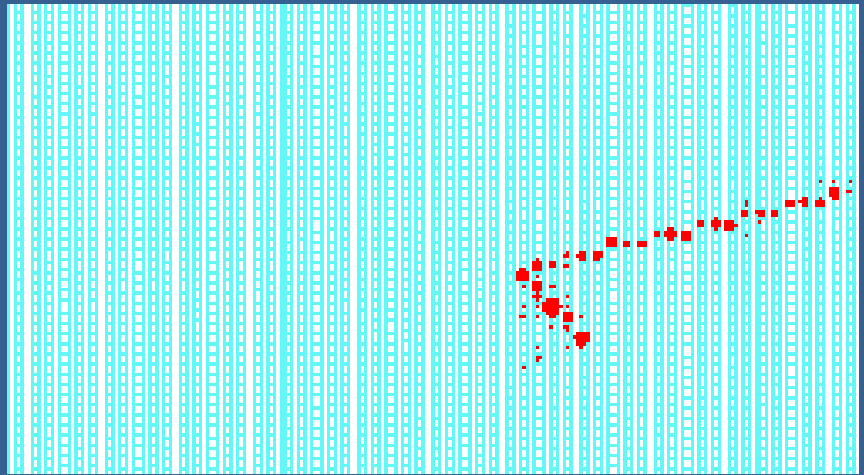
# Another example: BNB wrong sign background

CC  $\nu$ /anti- $\nu$  events typically distinguished by identifying charge of outgoing lepton. MB cannot distinguish  $\nu$ /anti- $\nu$  on an event by event basis.\*



Right Sign

(in anti- $\nu$  mode)



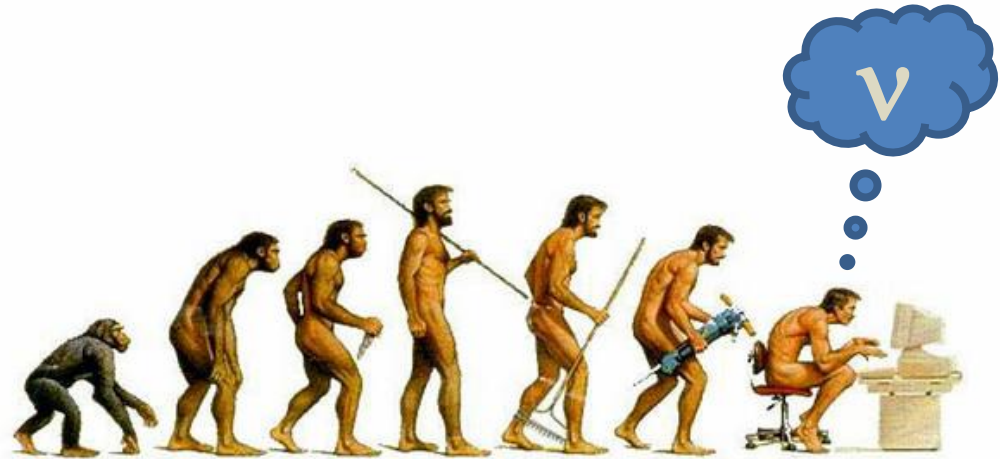
Wrong Sign

In SciBooNE these appear as 1-track vs. 2-track events

# SciBooNE Timeline

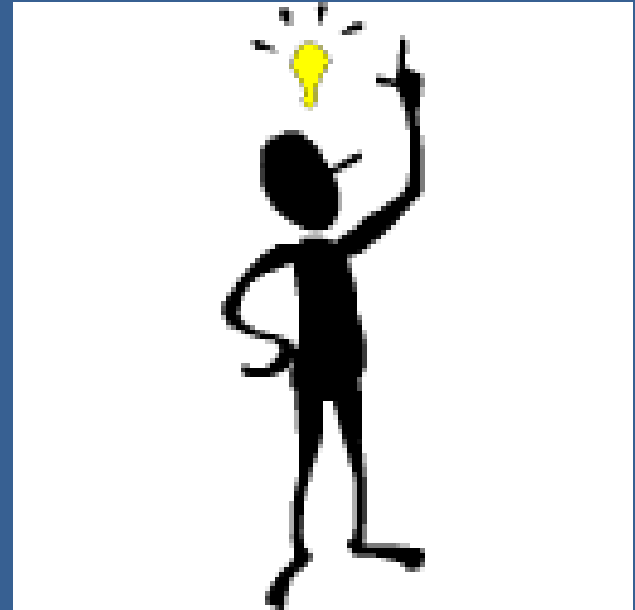
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SciBooNE has evolved a lot over the last two years.



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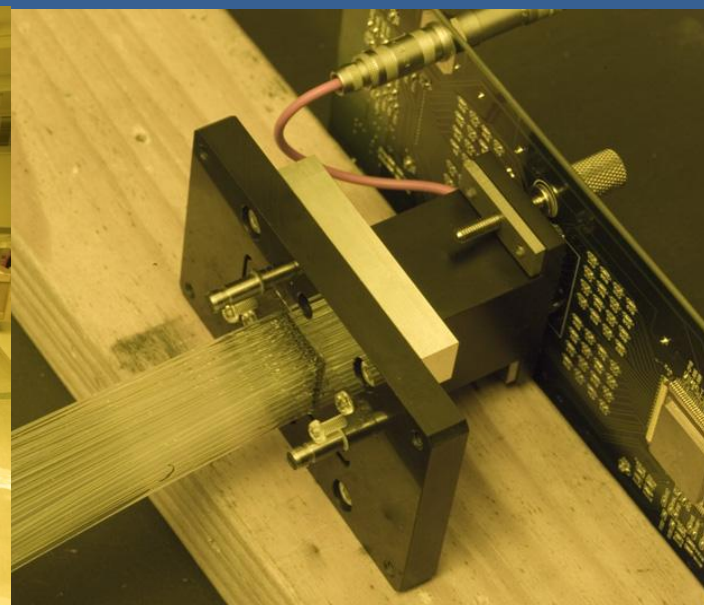
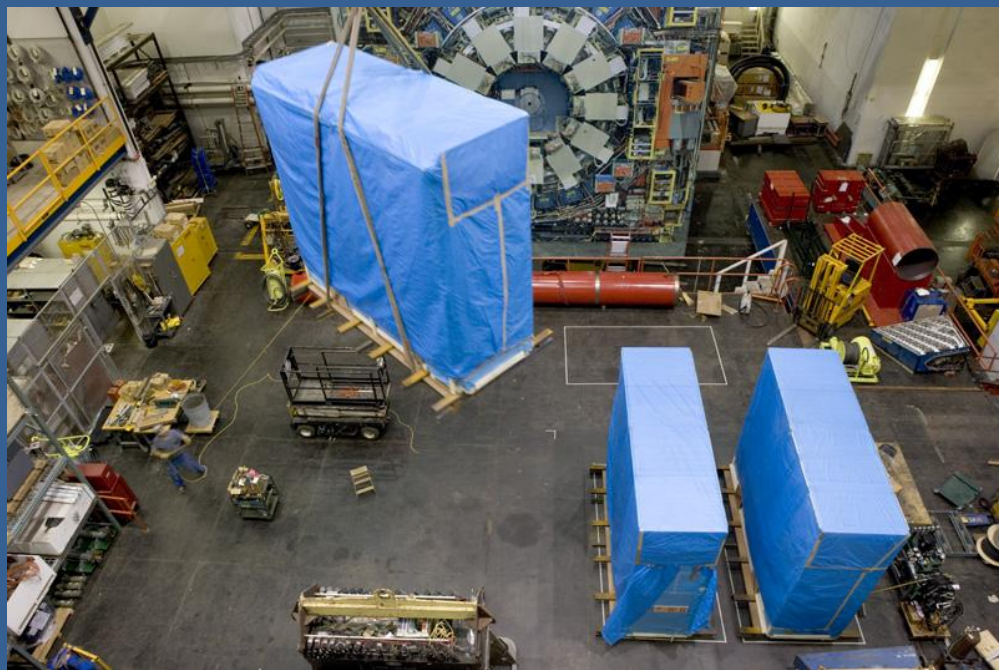
In 2005 an idea  
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20 September 2006



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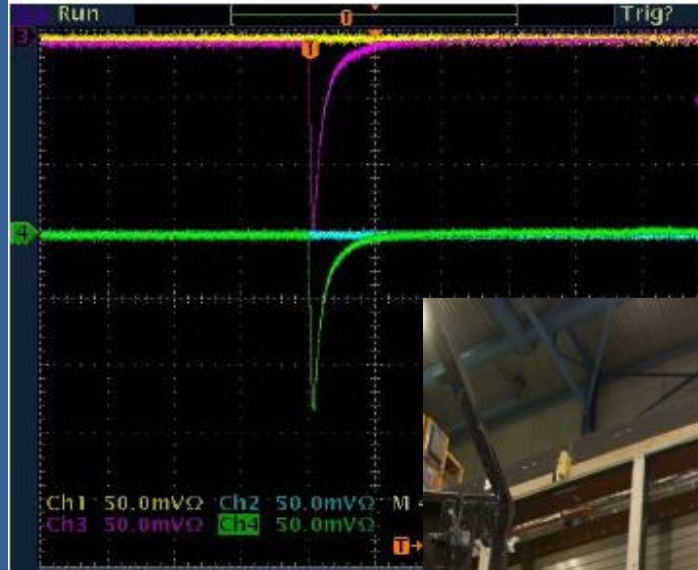
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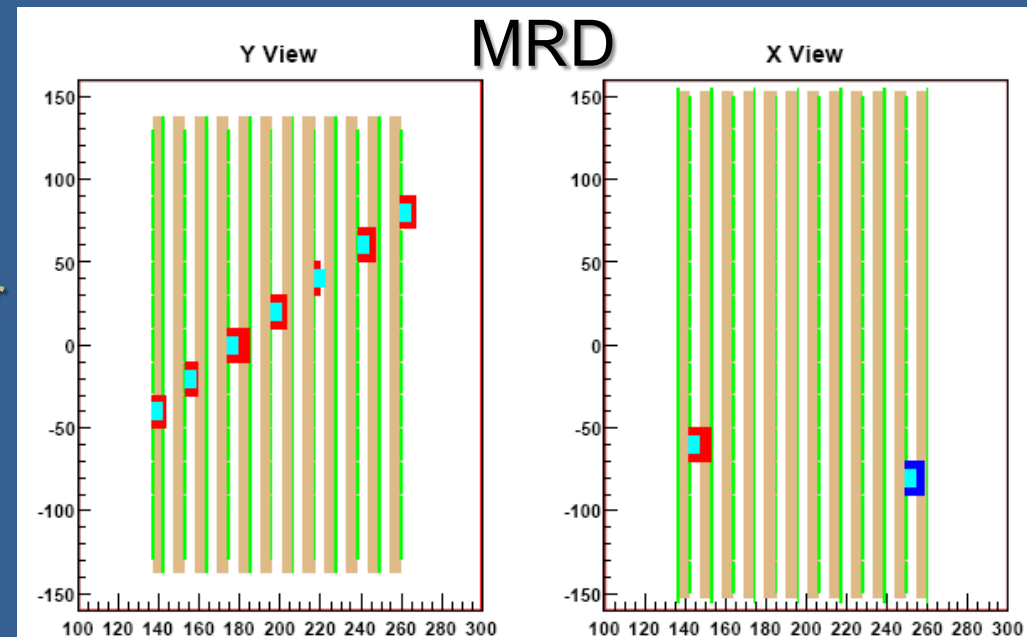
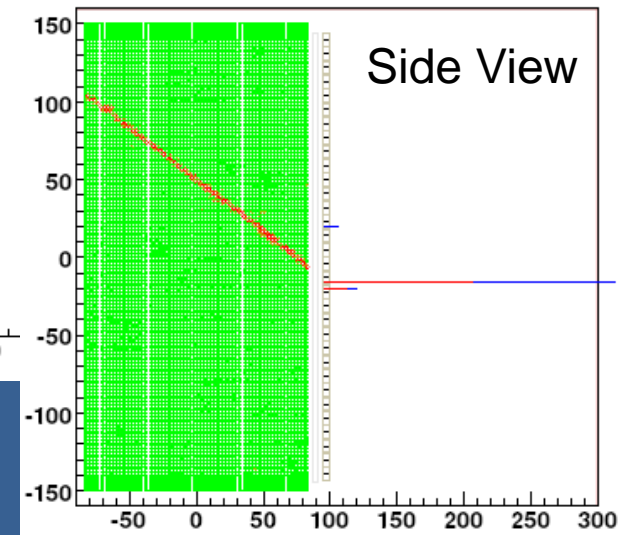
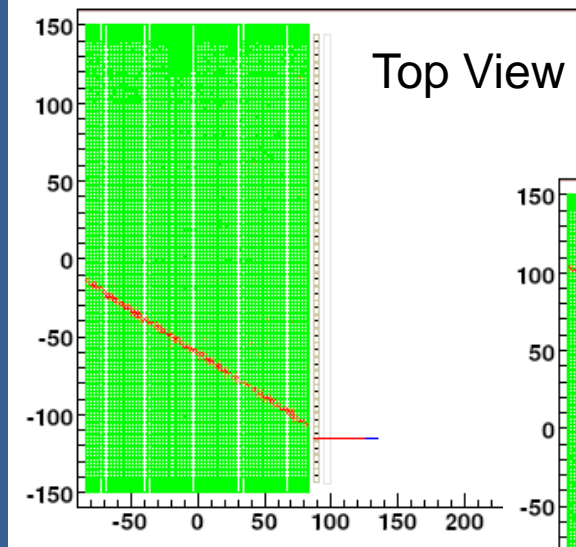




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SciBar/EC



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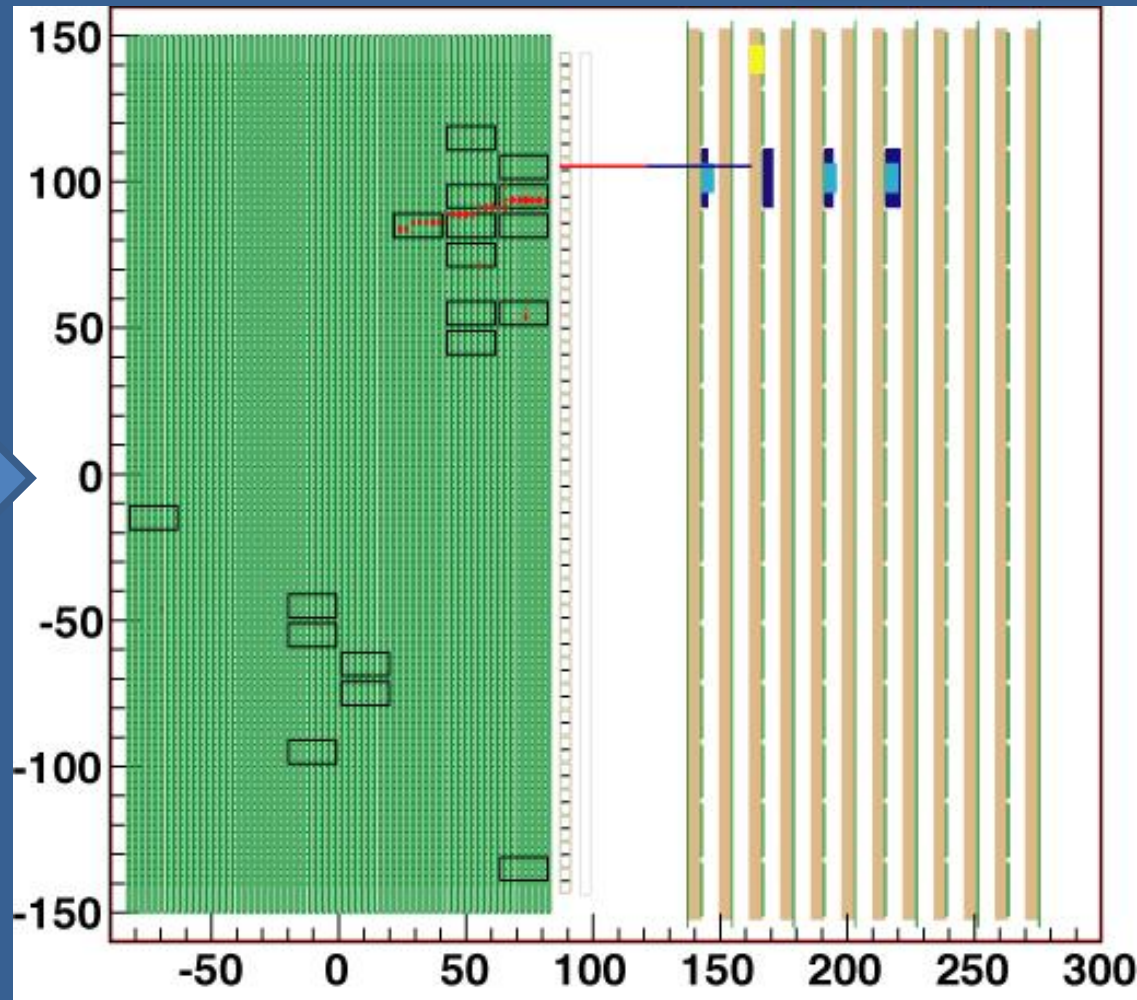




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Top View

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- **May 2007: First events seen in detector !!!**



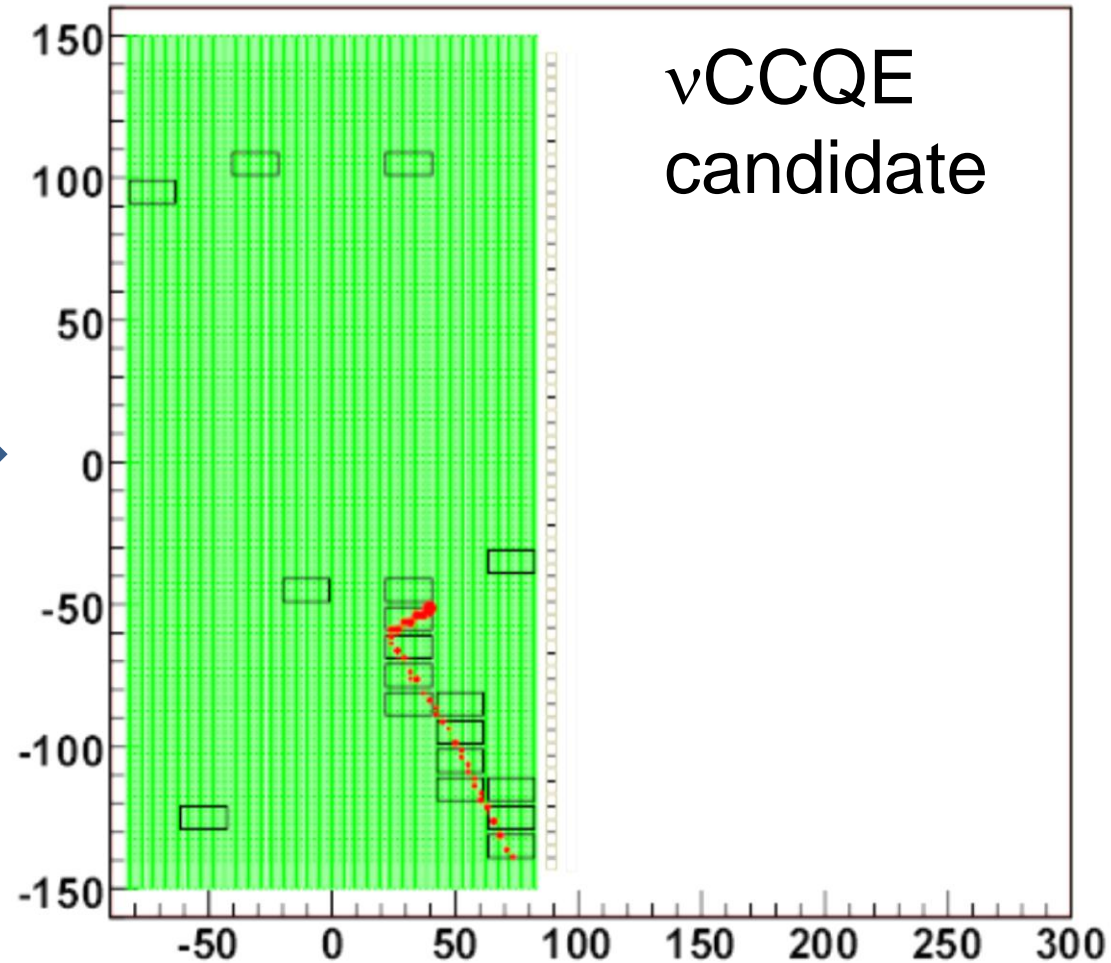
First  $\nu$  event including  
all three subdetectors!  
May 30, 2007 (today)

We are commissioning our detectors now, and have already begun seeing beam data!

Beam  
Direction



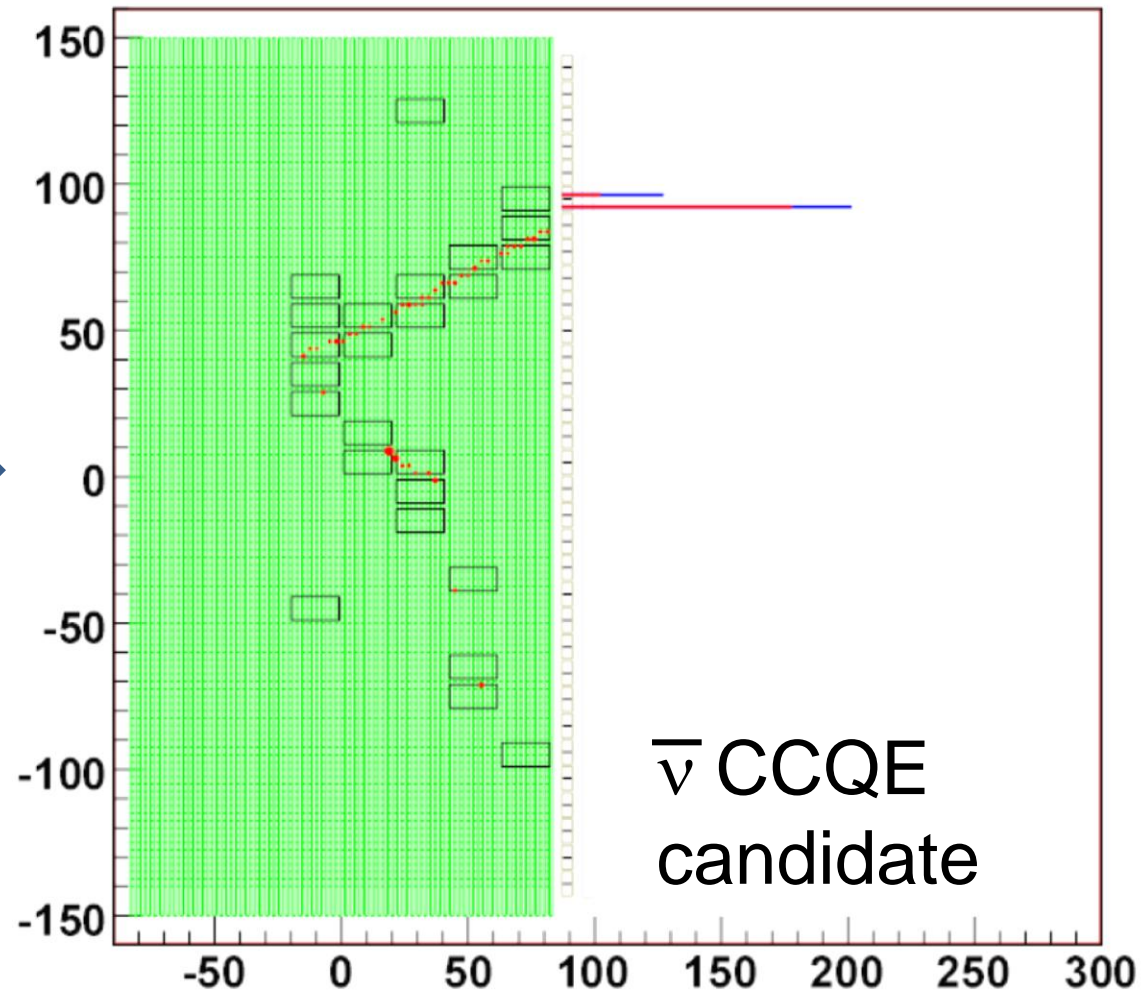
Top View





Top View

Beam  
Direction



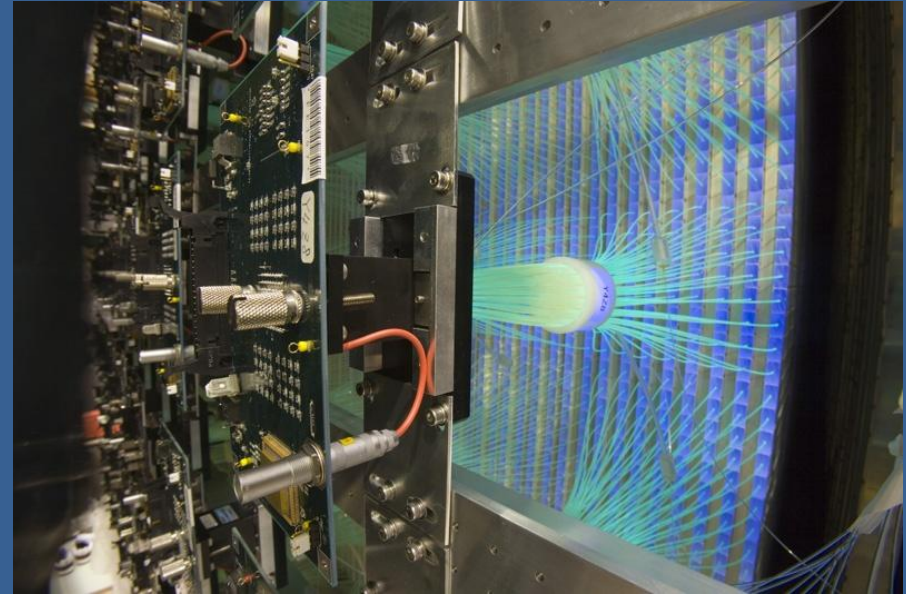
# SciBooNE has come a *long way* in a *short time*

The first collaboration meeting was on 17 Mar 2006, a little over a year ago.

Since then:

- New building has been constructed
- Disassembled, shipped, and reassembled two detector systems
- Designed and constructed a third detector system
- Designed and built a DAQ system
- Bulk of our analysis software already written

Started taking beam data!



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## ➤ Backup Slides

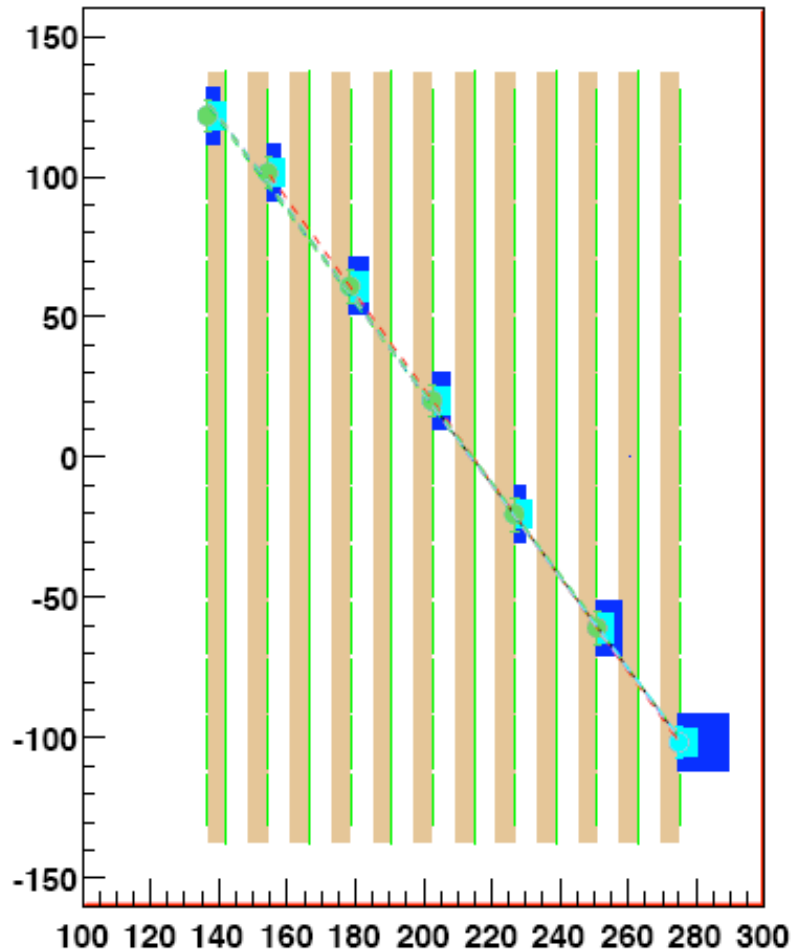
# Physics Motivation

- Neutrino Measurements (~5% precision)
  - CC- $1\pi$  cross section
  - CCQE  $\sigma$ ,  $M_A$  measurement
  - NC $\pi^0$  measurement
  - Search for CC coherent  $\pi$
  - Search for NC coherent  $\pi^0$
  - Search for radiative Delta decay ( $\nu + N \rightarrow \mu + N' + \gamma$ )
  - Intrinsic  $\nu_e$  flux for BNB ( $\nu_\mu \rightarrow \nu_e$  appearance search)
  - Unoscillated  $\Phi_\nu * \sigma$  for BNB ( $\nu_\mu \rightarrow \nu_\mu$  disappearance search)
- Antineutrino Measurements (~10% precision)
  - CCQE measurement, energy dependence of  $\sigma$  &  $M_A$
  - CC- $1\pi$  cross section
  - NC $\pi^0$  measurement, exclusive
  - Search for CC coherent  $\pi$
  - Search for NC coherent  $\pi^0$
  - Search for radiative Delta decay ( $\nu + N \rightarrow \mu + N' + \gamma$ )
  - Energy dependence of  $\nu$  contamination for BNB anti- $\nu$  mode

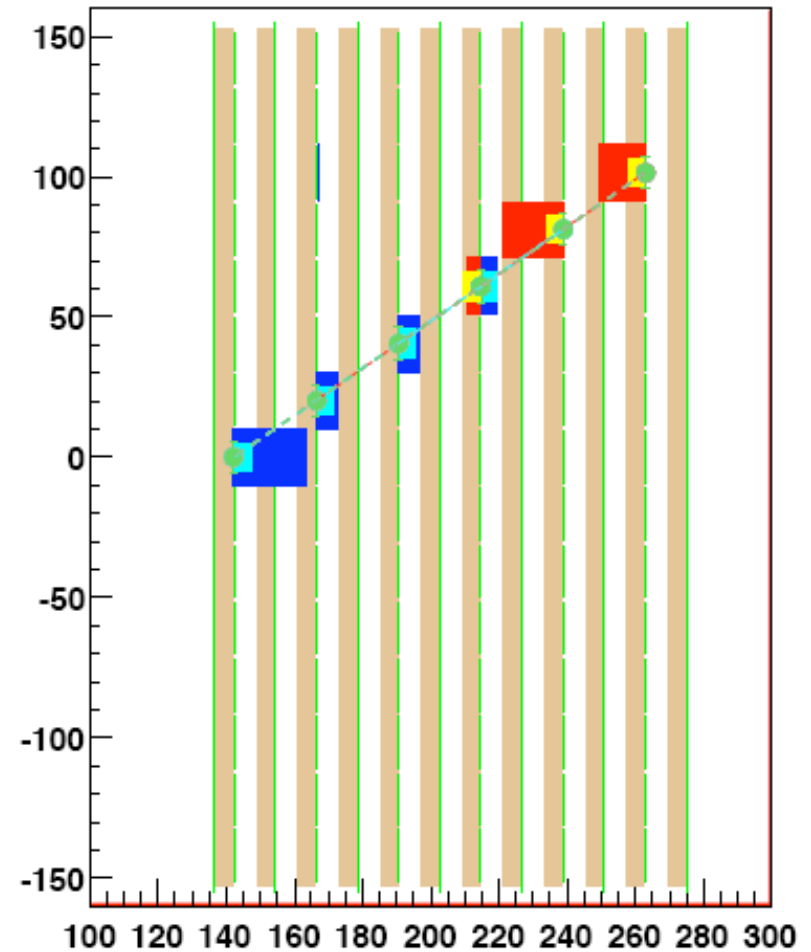


# Cosmic ray in MRD

Y View



X View



# The Leptonic Mixing Matrix

Flavor  
States

Mass  
States

$$\begin{pmatrix} \nu_e \\ \nu_\mu \\ \nu_\tau \end{pmatrix} = \begin{pmatrix} 1 & 0 & 0 \\ 0 & \cos\theta_{23} & \sin\theta_{23} \\ 0 & -\sin\theta_{23} & \cos\theta_{23} \end{pmatrix} \begin{pmatrix} \cos\theta_{13} & 0 & \sin\theta_{13} e^{-i\delta} \\ 0 & 1 & 0 \\ -\sin\theta_{13} e^{i\delta} & 0 & \cos\theta_{13} \end{pmatrix} \begin{pmatrix} \cos\theta_{12} & \sin\theta_{12} & 0 \\ -\sin\theta_{12} & \cos\theta_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix} \begin{pmatrix} \nu_1 \\ \nu_2 \\ \nu_3 \end{pmatrix}$$

“Atmospheric Mixing”

$$\theta_{23} \approx 45^\circ$$

$$\Delta m_{23}^2 \sim 2.5 \times 10^{-3} \text{ eV}^2$$

(SK, K2K, MINOS)

$$\theta_{13} \leq 10^\circ$$

$\theta_{13}$  constrained by  
reactor experiments  
(CHOOZ, Palo Verde),  
but not measured.

$\delta$  is unknown

“Solar Mixing”

$$\theta_{12} \approx 32^\circ$$

$$\Delta m_{12}^2 \sim 8 \times 10^{-5} \text{ eV}^2$$

(SNO, KamLAND)